Polarisation Observations of VY Canis Majoris H$_2$O $^{5}_{32} - ^{4}_{41}$
620.701 GHz Maser Emission with HIFI*


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Abstract

Context. Water vapour maser emission from evolved oxygen-rich stars remains poorly understood. Additional observations, including polarisation studies and simultaneous observation of different maser transitions may ultimately lead to greater insight.

Aims. We have aimed to elucidate the nature and structure of the VY CMa water vapour masers in part by observationally testing a theoretical prediction of the relative strengths of the 620.701 GHz and the 22.235 GHz maser components of ortho H$_2$O.

Methods. In its high-resolution mode (HRS) the Herschel Heterodyne Instrument for the Infrared (HIFI) offers a frequency resolution of 0.125 MHz, corresponding to a line-of-sight velocity of 0.06 km s$^{-1}$, which we employed to obtain the strength and linear polarisation of maser spikes in the spectrum of VY CMa at 620.701 GHz. Simultaneous ground based observations of the 22.235 GHz maser with the Max-Planck-Institut für Radiointernationelle 100-meter telescope at Effelsberg, provided a ratio of 620.701 GHz to 22.235 GHz emission.

Results. We report the first astronomical detection to date of H$_2$O maser emission at 620.701 GHz. In VY CMa both the 620.701 and the 22.235 GHz polarisation are weak. At 620.701 GHz the maser peaks are superposed on what appears to be a broad emission component, jointly ejected from the star. We observed the 620.701 GHz emission at two epochs 21 days apart, both to measure the potential direction of linearly polarised maser components and to obtain a measure of the longevity of these components. Although we do not detect significant polarisation levels in the core of the line, they rise up to approximately 6% in its wings.

Key words. stars: AGB and post-AGB - stars: winds, outflows - supergiants - circumstellar matter - masers - submillimetre

1. Introduction

VY Canis Majoris is a highly luminous, strongly obscured, variable supergiant with a high infrared excess. Only ~1% of the total luminosity is observed at optical wavelengths. The star’s distance has been measured to be $D = 1.1$ kpc (Choi et al., 2008) implying a luminosity $L = 3 \times 10^6 L_\odot$ (Menten et al., 2008). Three thermally emitted mid-infrared water vapour emission lines seen in spectra obtained with the Short Wavelength Spectrometer (SWS) on the Infrared Space Observatory (ISO) indicate a mean radial velocity of order $20 \pm 2$ km s$^{-1}$ and a $25$ km s$^{-1}$ H$_2$O outflow velocity (Neufeld et al., 1999), significantly lower than the $32$ km s$^{-1}$ velocities that Reid & Muhleman (1978) reported for 1612 MHz OH maser outflow.

The effective temperature of the central star, $T_e = 2,800$ K (Monnier et al., 1999), combined with distance and inferred luminosity fixes the star’s radius at $R \sim 10$ AU. Working at $11 \mu$m, Danchi et al. (1994) reported a variable photospheric radius ranging from $9.5$ to $11$ mas, where $10$ mas corresponds to a stellar radius of $10$ AU.

VY CMa is a strong source of $6_{16} - 5_{32}$, 22.235 GHz water vapour maser emission. Recent observations at the Atacama Pathfinder Experiment (APEX) telescope have revealed H$_2$O maser emission also at eight submillimetre frequencies, ranging from 321 to 475 GHz (Menten et al., 2008). Although a theoretical model of Neufeld & Melnick (1991) predicted that the 22.235 GHz transition should be far more luminous than the submillimetre transitions, Menten et al. (2008) find them to be comparable in flux density, or at most a factor of ~6 lower. This suggests a need for further observations that might explain the discrepancies.

One of strongest submillimetre masers predicted by Neufeld & Melnick (1991) is due to the $5_{32} - 4_{41}$, 620.701 GHz transition of ortho water, whose photon luminosity was expected to be roughly 16% of the 22.235 GHz maser. Given the multiple masers observed at 22.235 GHz we expected the 620.701 GHz maser of VY CMa to exhibit similar multiple peaks at photon densities ~16% of those exhibited at 22.235 GHz.

2. HIFI Observations at 620.701 GHz

The Herschel Heterodyne Instrument for the Far Infrared (HIFI) covers seven frequency bands, ranging from 488.1 to 1901.8 GHz (de Graauw et al., 2010). The 620.701 GHz radiation of the $5_{32} - 4_{41}$ maser of ortho H$_2$O is observed in HIFI’s Band 1B, which nominally covers the range from 562.6 to 628.4 GHz. Two channels cover each of the frequency bands, one sensitive to linearly polarised radiation roughly parallel to the spacecraft horizontal (H) direction, the other roughly parallel to the vertical direction (V). For Band 1B the H direction of polarisation is at an angle of 82.5° relative to the spacecraft V axis, while the V direction of polarisation is at an angle of -7.5° to that axis. Our data were obtained in the HIFI high-resolution spectroscopy (HRS)
mode with spectral resolution 0.125 MHz, or line-of-sight velocity resolution 0.06 km s\(^{-1}\).

The H and V beams on HIFI are not fully coincident. In Band 1B they are separated by \(\sim 19\%\) of the \(\sim 34.4''\) full-width-half-power beam diameter, the offset between the two beams is \(\sim 6.6''\). VY CMa is a spatially unresolved source at these frequencies (Decin et al., 2006). In our observations the star was positioned half-way between beam centres, i.e., displaced \(\sim 3.3''\) relative to the centre of each beam, although a random \(\sim 1.8''\) pointing error can slightly increase alignment uncertainty.

In order to determine the orientation on the sky of any observed linear polarisation, a source has to be viewed at least at two different rotation angles relative to the telescope. The Herschel Space Observatory (Pilbratt et al., 2010), however, cannot be rotated without producing undesirable thermal drifts. A rotation is best achieved by observing a target at two epochs separated by a number of weeks. The further the target lies above the ecliptic plane, the faster is the rotation produced. In our observations, an interval of three weeks between two sightings of VY CMa resulted in a rotation of the telescope of \(\sim 16°\) about the line of sight to the star.

We first observed the star on March 21, 2010 in two contiguous segments, lasting 10,139 seconds apiece, for a total observing time of 5.633 hours. The first segment began at 08:48:38.0 UT; the second started at 11:45:00 UT, 7 minutes and 23 seconds after the first had terminated. Our second epoch of observations started at 17:02:33 UT on April 11, 2010 and again ran for 10,139 seconds or 2.816 hours. The system temperature for these observations ranged between approximately 80 K and 100 K with respective RMS noise levels for the first and second sets of observations of 6.1 and 5.3 Jy for the H channel, and 4.3 and 4.5 Jy for the V channel at the resolution of 0.125 MHz.

3. Polarisation Analysis of the 620.701 GHz Maser Components

Linear polarisation measurements are often optimised to determine the Stokes \(Q\) and \(U\) parameters independently from each other (Li et al., 2008; Hezareh & Houde, 2010). If we define four intensity measurements \(I_0\) within the equatorial system with \(q\) set to 0° when pointing north and increasing eastwards, then

\[
\begin{align*}
I &= (I_0 + I_{45} + I_{90} + I_{135})/2 \\
Q &= I_0 - I_{90} \\
U &= I_{45} - I_{135}.
\end{align*}
\]

Although the state of linear polarisation is completely defined with these equations, it is also commonly expressed with the polarisation fraction \(p = \sqrt{Q^2 + U^2}/I\) and angle \(\theta = 0.5\arctan(U/I)\). Alternatively, we can also write

\[
\begin{align*}
Q &= pl\cos(2\theta) \\
U &= pl\sin(2\theta).
\end{align*}
\]

With HIFI we do not generally have access to four independent measurements precisely sampled with a spacing of 45°, as in Equations (1)-(3). Instead, we may have access to \(N\) irregularly spaced observing epochs with the HIFI vertical polarisation axis (V) oriented at an angle \(\phi_v\) relative to north \((i = 1, 2, \ldots, N)\) and the horizontal polarisation axis (H) at \(\phi_h + 90°\), each observation yielding a pair of intensities \(I_i\) and \(I_{ih}\).

These \(2N\) intensities yield the polarisation if we first define

\[
d_i = I_i - I_{ih}.
\]

Introducing a vector \(d\) with elements \(d_i/\sigma_i\), where \(d_i\) is the intensity difference of Equation (6) and \(\sigma_i\) its uncertainty (i.e., the corresponding noise intensity), we can write the matrix equation

\[
d = A \cdot s,
\]

where the matrix \(A\) and vector \(s\) are given by

\[
A = \begin{pmatrix}
\cos(2\phi_1)/\sigma_1 & \sin(2\phi_1)/\sigma_1 \\
\vdots & \vdots \\
\cos(2\phi_N)/\sigma_N & \sin(2\phi_N)/\sigma_N
\end{pmatrix}
\]

\[
s = \begin{pmatrix}
Q \\
U
\end{pmatrix}.
\]

The vector \(s\), and therefore the Stokes parameters, are easily obtained by inverting Equation (7) using the pseudo-inverse matrix of \(A\). We then find

\[
s = (A^T A)^{-1} A^T \cdot d.
\]

The elements of the covariance matrix \(C = (A^T A)^{-1}\) can be shown to yield the corresponding uncertainties on the estimated Stokes parameters with

\[
C = \begin{pmatrix}
\sigma_{QQ}^2 & \sigma_{QU}^2 \\
\sigma_{UQ}^2 & \sigma_{UU}^2
\end{pmatrix},
\]

where \(\sigma_{ij}^2 = \langle s_is_j \rangle - \langle s_i \rangle \langle s_j \rangle\) with \(i, j = Q\) or \(U\) and \(s_i\) the appropriate component of the vector \(s\). It is also possible to determine the uncertainties \(\sigma_p\) and \(\sigma_\theta\) for the polarisation fraction and angle, respectively, with

\[
\sigma_p^2 \approx \frac{1}{P^2 I^2} (Q^2\sigma_{QQ}^2 + U^2\sigma_{UU}^2 + 2QU\sigma_{QU}^2),
\]

\[
\sigma_\theta^2 \approx \frac{1}{4P^2 I^2} (Q^2\sigma_{UU}^2 + U^2\sigma_{QQ}^2 - 2QU\sigma_{QU}^2).
\]

We have applied this polarisation analysis to our observations of the 620.701 GHz \(\text{H}_2\text{O}\) 532 – 441 ortho-transition. Our results shown in Figure 1 were obtained from the two aforementioned observations made at epochs of corresponding position angles of 261.27° and 277.46°.

Although there are no obvious strong polarisation signals from the maser emission peaks, we clearly detect polarisation levels ranging from \(p \approx 1.5\%\) to \(p \approx 6\%\) in regions of significant line intensity (i.e., from approximately -5 to 45 km s\(^{-1}\)). Furthermore, the observed anti-correlation of the polarisation fraction with the Stokes \(I\) intensity is similar to previous ground-based polarisation observations (Girart et al., 2004; Hezareh & Houde, 2010) aimed at the detection of the Goldreich-Kylafis effect in non-masing molecular lines (Goldreich & Kylafis, 1981; Cortes et al., 2005), which appears to have first been detected in evolved stars (Glenn et al., 1997). We will discuss the relevance of the Goldreich-Kylafis effect for our observations in section 7 below.
Figure 1. Polarisation spectrum of the 620.701 GHz H$_2$O $5_{32}-4_{41}$ ortho-transition for VY-CMa. Shown are, from top to bottom, the Stokes $I$ and polarisation fraction (symbols and right axis), the polarization angle, and the (unsigned) polarised flux. The polarisation fraction and flux were corrected for positive bias but not for instrumental polarisation. A simple linear baseline was removed from each $V$ and $H$ channel observation and all polarisation vectors shown satisfy $p > 3\sigma_p$ and $I > 10\sigma_I$ at the corresponding velocities, where $\sigma_I$ is the uncertainty (or noise level) on the Stokes $I$ spectrum (i.e., the average of the $IV+IH$ spectra measured at the two epochs). The polarisation fraction and angle have their spectral resolution reduced by a factor of 20 compared to Stokes $I$, which is at the maximum resolution.

4. Results

Several instrumental capabilities of HIFI and the observations obtained with them may be noted:

(i) The two orthogonally polarised HIFI receivers are well matched and extremely stable. But observations of extended sources need to be conducted with caution. A slight misalignment of H and V receivers can lead to a "false polarisation" that reverses polarity at half-year intervals (see Appendix A).

(ii) The misalignment of the HIFI receivers does not appear to affect observations of unresolved sources. Our observations realised at two observing epochs indicate that instrumental polarisation, which could be in part due to errors in the relative calibration between the two receiver chains of Band 1B, cannot exceed a measure of order 1 – 2% (see section 7 below).

(iii) The polarisation of VY CMa is not significant near the peak of the 620.701 GHz H$_2$O $5_{32}-4_{41}$ line, but rises up to $\sim 6\%$ in the wings of the spectrum in a manner consistent with polarisation due to the Goldreich-Kylafis effect discussed in greater detail in section 7 below.

(iv) The stability of the 620.701 GHz masers is remarkable. The variation over a three week period is $\lesssim 1\%$ (see Figure 2).

(v) As Figure 3 shows, the spectral profile of the 620.701 GHz and 22.235 GHz masers appears remarkably similar, except that the relative expansion velocities between peaks, along the line of sight, is a factor of 2.3 greater at 620.701 GHz than for contemporaneous observations at 22.235 GHz (see Appendix B).

(vi) Assuming the 620.701 GHz masers to sit atop a broad pedestal with flux density $\sim 400$ Jy, we find the main 620.701 GHz maser peak flux density to be $\sim 800$ Jy, compared to $\sim 1900$ Jy at 22.235 GHz. The velocity spread of the expanding 620.701 GHz maser peaks is $\sim 2.3$ times wider than at 22.235 GHz, implying a 620.701 GHz photon luminosity roughly 2.4 times lower than at 22.235 GHz — comparable in luminosity to the submillimetre masers observed by Menten et al. (2008), and about 2.5 times more luminous than Neufeld & Melnick (1991) predicted.

5. Discussion

A spherically symmetric model for the outflows of VY CMa developed by Decin et al. (2006) envisions dust formation in the cooling outflow from the star at distances $\sim 10\, R_*$. Radiation pressure then accelerates the dust, and with it also the ambient gas, from velocities $\lesssim 5\, \text{km}\, \text{s}^{-1}$ to $\sim 25\, \text{km}\, \text{s}^{-1}$ at a distance of $20\, R_*$. The kinetic temperature at $20\, R_*$ is of the order of $\sim 1000$ K, sufficient to excite both 620.701 and 22.235 GHz masers. We propose that the masing outflows we observe propagate along several mutually common directions from the star, and that the faster expanding 620.701 GHz masers lie at greater distances from the star than the central cluster of masers peaking at 22.235 GHz. We picture this outflow consisting of gas ejected obliquely to the line of sight to VY CMa but fanning out over a considerable angular width to account for the observed range of outflow velocities. Common ejection of both the denser
Nevertheless, a feature common to all the submillimetre spectra of Menten et al. (2008) is a broad pedestal that diminishes in strength with rising upper-excitation temperature $E_u$ of the observed transitions. For $E_u \approx 725$ K it remains strong with wings extending over a full width $\gtrsim 40$ km s$^{-1}$, whereas at $E_u = 1861$ K, the pedestal has entirely disappeared. The radiation emitted in the pedestal thus appears to be due to collisional excitation by gas flowing out from the star at speeds $\gtrsim 20$ km s$^{-1}$ close to the terminal velocity in the model of Decin et al. (2006). The pedestal we observe for the transition at 620.701 GHz, whose $E_u = 732$ K, has the same broad wings extending roughly symmetrically to either side of the maser outflow peaks.

Low $E_u$ levels thus excited can then radiate in three ways, spontaneously, through stimulated emission, or induced by collisions. As Neufeld & Melnick (1991) point out, stimulated emission leads to masing only when the ratio of water vapor density to the velocity gradient along the line of sight is sufficiently high. We interpret the maser peaks in our 620.701 GHz spectrum as emanating from denser clumps in the VY CMa outflow within which the velocity gradient is low, and the 620.701 GHz pedestal as emission from an ambient medium characterized by lower density and/or higher velocity gradients, and thus unable to radiate significantly through stimulated emission to sustain maser amplification. This ambient medium, and its corresponding radiation, will bear the imprint of the Goldreich-Kylafis (GK) effect when also subjected to anisotropic radiation or optical depths, and permeated by a weak magnetic field.

Although the low polarization levels we detect are expected for water masers (Watson, 2009), the rise in polarization seen in the wings of the spectrum in the top panel of Figure 1 deserves careful consideration. If we assume that the masers simply maser-amplify the seed radiation emanating from the ambient gas responsible for the pedestal component, then we may have a qualitative explanation for the shape of the polarization spectrum. That is, the polarization fraction should be approximately the same at the position of, and between, maser peaks, as observed in Figure 1, and should rise in the wings of the line where the optical depth of the pedestal component is lowest, in a manner consistent with the GK effect (Cortes et al., 2005).

Although maser-amplification of GK polarisation could lead to different levels of polarisation at the maser peaks, other factors (e.g., magnetic field orientation) could also suppress this effect. We also note that in cases where the pedestal component was highly saturated, the level of polarization measured in its core would provide a measure of the amount of instrumental polarisation (Hezareh & Houde, 2010). Although this scenario probably does not perfectly fit our observations, the polarization fraction measured in the vicinity of the line core suggests that the instrumental polarisation is approximately 1 - 2%.

Conclusions

We report the first astronomical observations of the 620.701 GHz $5_{22} - 4_{11}$ submillimetre maser of ortho H$_2$O. The maser peaks in VY CMa show weak linear polarisation at levels consistent with the Goldreich-Kylafis effect across the spectral line profile.

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References

Appendix A: A Cautionary Note on Observations of Extended Sources
As part of the Performance Verification of HIFI in space, spectroscopic observations of the $^{12}$C$^{16}$O $J = 5 \rightarrow 4$ 576.3 GHz transition were obtained at the position of the chemically-active bow shock B1 driven by the LDN 1157 Class 0 protostar (Bachiller et al., 2001). Initial observations in HIFI’s Band 1B were made on August 1, 2009, a second set was obtained 186 days later on February 4, 2010. The spectral shape of the source remained essentially unchanged but initially the signal from the high-frequency edge of the line appeared stronger in the V direction, when the H and V peaks were matched, whereas half a year later it appeared stronger by the same ratio in the H direction.

This “false polarisation” effect was due to the small misalignment of the H and V beams described earlier, which imaged slightly different portions of a scene onto the two receivers (Attard et al., 2008). Unless otherwise specified, observations of a targeted source are centred on a position halfway between the two beam centres. Because the viewing direction of Herschel at all times is constrained by the need to keep the plane of the sun-shield roughly perpendicular to the radius vector to the Sun, sources close to the ecliptic plane can only be viewed at half year intervals. Over this interval the telescope aperture rotates $180^\circ$ about its viewing direction on the sky. Thus, portions of LDN 1157 B1 initially viewed by the V channel were viewed, half a year later, by the H channel and vice versa. It is clear that this effect was not due to linear polarisation because a $180^\circ$ rotation of the telescope leaves the polarisation direction unchanged.

For VY CMa, where the entire source was relatively well centred on each beam, this effect was negligible. But for linear polarisation observations of an extended source the best strategy may be to make a small map with HIFI over the area of interest, where the spectra obtained in the H and V channels at each position of the source, can be individually and directly compared.

Appendix B: Ground-based Observations at 22.235 GHz
Overlapping with the second epoch of 620.701 GHz observations, we observed the VY CMa 22.235 GHz maser with the Effelsberg 100-meter telescope of the Max-Planck-Institut für Radioastronomie. Those observations were begun on April 11, 2010 at 16:40 UT and lasted until 17:20 UT and are shown in Figure 3. Data in two linear polarisations were obtained with the two channels of the K-band or 1.3 cm receiver at the primary
focus. The frequency resolution was 6.104 kHz, corresponding to a velocity resolution of 0.082 km s$^{-1}$. Although these observations were realised at only one epoch and are, therefore, insufficient for a full characterisation of the linear polarisation state, a search by Vlemmings et al. (2002) found that the linear polarisation of the VY CMa 22.235 GHz masers is well below 1%.