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Evidence for CO depletion in the inner regions of gas-rich protoplanetary disks

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\textbf{ABSTRACT}

\textbf{Aims.} We investigate the physical properties and spatial distribution of Carbon Monoxide (CO) gas in the disks around the Herbig Ae/Be stars HD 97048 and HD 100546.

\textbf{Methods.} Using high-spectral-resolution 4.588-4.715 \textmu m spectra containing fundamental CO emission taken with CRIRES on the VLT, we probe the circumstellar gas and model the kinematics of the emission lines. By using spectro-astroometry on the spatially resolved targets, we constrain the physical size of the emitting regions in the disks.

\textbf{Results.} We resolve, spectrally and spatially, the emission of the \textsuperscript{13}CO v(1-0) vibrational band and the \textsuperscript{12}CO v = 1 - 0, v = 2 - 1, v = 3 - 2 and v = 4 - 3 vibrational bands in both targets, as well as the \textsuperscript{12}CO v = 5 - 4 band in HD 100546. Modeling of the CO emission with a homogeneous disk in Keplerian motion, yields a best fit with an inner and outer radius of the CO emitting region of 11 and > 100 AU for HD 97048. HD 100546 is not fit well with our model, but we derive a lower limit on the inner radius of 8 AU. The fact that gaseous [OI] emission was previously detected in both targets at significantly smaller radii suggests that CO may be effectively destroyed at small radii in the surface layers of these disks.

\textbf{Key words.} circumstellar matter, stars – pre-main-sequence, stars – HD 97048, stars – HD 100546, protoplanetary disks

1. Introduction

We present a study of warm carbon monoxide (CO) gas in two circumstellar disks surrounding young intermediate-mass stars. CO is commonly detected in disks (Najita et al., 2003), and often used as a tracer of the physical conditions in the disk and the kinematics of the gas. Ro-vibrational transitions of warm (T = 120-1010 K) CO gas have already been found by Mitchell et al. (1990) in 8 out of 9 surveyed embedded infrared sources. Brittain et al. (2007) had a 100 \% detection rate for the ro-vibrational CO transitions in the disks around 9 intermediate-mass stars with optically thick inner disks, while their detection rate is only 1 out of 5 for disks with an optically thin inner disk, a case they explain by UV fluorescence. Blake & Boogert (2004) have detected CO gas in 5/5 Herbig Ae stars and recently, Pontoppidan et al. (2008) have spatially and spectrally resolved the 4.7 \textmu m CO ro vibrational lines in the disks around 3 young solar mass stars with known dust gaps or inner holes, and detected CO well inside the dust gaps in all disks.

HD 97048 (Sp. type A0psh, (van den Ancker et al., 1998)) is a well-studied nearby Herbig Ae star, whose spectral energy distribution (SED) displays a large infrared excess above the emission from the star, commonly explained by emission from warm dust in a circumstellar disk. The disk is a flaring disk according to the classification proposed by Meeus et al. (2001). Lagage et al. (2006) have resolved the disk around HD 97048 in the mid-IR and determined the flaring angle. PAH emission from radii up to 200-300 AU has been detected by van Boekel et al. (2004). Habart et al. (2004) have resolved emission from the 3.43 \textmu m and 3.54 \textmu m diamond bands from within the inner 15 AU in the disk. Circumstellar gas has also been detected, e.g. by Acke & van den Ancker (2006) who detect spatially and spectrally resolved [OI] emission from a few tenths to several tens of AU. HD 97048 is one of the only two Herbig stars where the H\textsubscript{\beta} S(1) line at 17.035 \textmu m has been detected (Martin-Zaidi et al., 2007), coming from warm (T\textsubscript{ex} \leq 570 K) gas within 35 AU.

HD 100546 (Sp. type B9Vne (van den Ancker et al., 1998)) also has a SED characteristic for a flared disk. Its IR spectrum indicates a high crystalline silicate dust fraction in the disk surface layers (Bouwman et al., 2003). There is observational evidence for a disk gap, or ‘second wall’, at \approx 10 AU from the central star. This was derived based on the SED (Bouwman et al., 2003), and confirmed by nulling interferometry (Liu et al., 2007). The [OI] 6300 \AA \ line was studied by Acke & van den Ancker (2006), who also find evidence for a disk gap at that distance, possibly induced by a very low mass stellar companion or planet of \approx 20 M\textsubscript{Jup}, or a planetary system. The latter authors spatially and spectrally resolve the [OI] emission, and determine it to originate from a few tenths to several tens of AU. Geers et al. (2007) have...
Table 1. Astrophysical parameters of the programme stars

<table>
<thead>
<tr>
<th>HD</th>
<th>log Teff</th>
<th>log Lbol</th>
<th>M/M⊙</th>
<th>Dist. (pc)</th>
<th>i</th>
<th>PA</th>
</tr>
</thead>
<tbody>
<tr>
<td>97048</td>
<td>4.00</td>
<td>4.28</td>
<td>2.5</td>
<td>180±9</td>
<td>42.5±1.6</td>
<td>160±0.8</td>
</tr>
<tr>
<td>100546</td>
<td>4.02</td>
<td>4.46</td>
<td>2.81</td>
<td>10±3</td>
<td>42.5±1.6</td>
<td>145±5.3</td>
</tr>
</tbody>
</table>

a van den Ancker et al. 1998 b Lagage et al. 2006 c Acke & van den Ancker 2006 d Ardila et al. 2007

Table 2. Log of spectroscopic observations at June 16th 2007

<table>
<thead>
<tr>
<th>Object</th>
<th>UT start hh:mm:ss</th>
<th>Ref. wavelength [nm]</th>
<th>N exp</th>
<th>Exp time [s]</th>
<th>PA range [°]</th>
</tr>
</thead>
<tbody>
<tr>
<td>HD 97048</td>
<td>00:41:15</td>
<td>4662.1</td>
<td>20</td>
<td>15</td>
<td>41.1 - 45.6</td>
</tr>
<tr>
<td>HD 97048</td>
<td>00:58:43</td>
<td>4676.2</td>
<td>20</td>
<td>15</td>
<td>46.0 - 50.3</td>
</tr>
<tr>
<td>HIP 052419</td>
<td>02:36:14</td>
<td>4662.1</td>
<td>2</td>
<td>15</td>
<td>85.1 - 85.3</td>
</tr>
<tr>
<td>HIP 061585</td>
<td>03:52:30</td>
<td>4676.1</td>
<td>2</td>
<td>15</td>
<td>73.1 - 73.3</td>
</tr>
</tbody>
</table>

A sub-pixel wavelength shift was manually added to reduce telluric residuals (i.e. spikes due to minor wavelength mismatches). Some telluric absorption lines are fully saturated, causing problems with the division of the spectra (i.e. division by 0). In the further data reduction, we will disregard these areas. We have shifted the spectrum to center the \( v = 2 - 1 \) lines on their rest wavelength. The spectra are shown in Fig. 1.

3. Observational Analysis

To check whether the targets are spatially resolved, we perform spectro-astrometry on the reduced 2D frames. We determine the location of the photocenter of the PSF and the spatial FWHM as function of wavelength (or equivalently velocity for each line), by fitting a Gaussian to the spatial profile in each spectral column. The relative position of the photocenter to the continuum, the Spatial Peak Position (SPP), is a measure for the spatial offset of the emitting region compared to the continuum. Our observations are made with the slit positioned over a small range in position angles. We note that Pontoppidan et al. (2008) have observed three nearby T Tauri stars at 6 different PA’s with CRIRES and the same observational set-up as employed by us, and find, after pairwise subtracting parallel and anti-parallel spectro-astrometric signals, that no telluric feature produces false spectro-astrometric signals in excess of 0.2-0.5 mas. Since we do not possess anti-parallel spectra, we show the telluric absorption in Figs. 2 and 3, and note, by comparing with CO lines un-affected by telluric absorption, that their effect on the astrometric signal is small. We compare the SPP in the CO lines to the SPP of the underlying continuum (SPP=0 by definition). If the target is resolved, and the photocenter of the CO emission region is offset to that of the dust continuum emission from the disk, the SPP is expected to be non-zero. The photocenter displacement can be derived from the measured SPP if we correct for dilution by the continuum flux, by multiplying the SPP with a factor \( 1 + F_c(v)/F_l(v) \), with \( F_c(v) \) and \( F_l(v) \) the continuum and line flux at velocity \( v \) (Pontoppidan et al., 2008).

The Gaussian fit to the PSF also yields a spatial FWHM. If the continuum is unresolved, this quantity is a measure for the achieved effective spatial resolution in the direction parallel to the slit. When resolved, the FWHM allows us to determine the spatial extent of the emission. Outside the spectral lines it then is a direct probe for the size of the continuum emitting disk, whilst inside the lines it is a function of both the line- and continuum emission, and its value a lower limit to the CO emission. Thus, measuring the SPP and the FWHM of the PSF in the continuum and the CO lines allows us to determine [1] the size of the continuum emitting region, and [2] the spatial offset of each CO line with respect to the continuum. This procedure is shown in Figs. 2 and 3.

4. Results

We detect in the spectrum of HD 97048 the \( ^{13} \text{CO} \) \( v = 1 - 0 \), as well as in the \( ^{12} \text{CO} \) \( v = 1 - 0 \) to \( v = 4 - 3 \) vibrational bands. All lines are spectrally resolved with a remarkably constant FWHM of 15 km s\(^{-1}\). The Half Width at Zero Intensity (HWZI) is more difficult to estimate due to the continuum S/N, but we make a conservative estimate of 15 km s\(^{-1}\) corresponding to an inner radius of 4.5 AU, assuming the stellar parameters and inclination as given in Table 1. Average line profiles and astrometric signals of the \( ^{12} \text{CO} \) \( v = 1 - 0 \) to \( v = 3 - 2 \) are shown in Fig. 5 and Table 3. The 4.6 μm continuum emission of the disk around HD
97048 is resolved. Its FWHM of 0.231" ± 0.004" corresponds to 23 ± 3 AU, after correcting for the FWHM of the PSF. The latter (0.192" ± 0.001") was estimated from 10 unresolved telluric standards, all sufficiently bright for the AO loop to close at maximum frequency, observed with the same instrumental settings, and within 2 days of our observations. The red- and blue-shifted wings of the SPP form a sinusoidal signal, typical for an inclined circumstellar disk.

In the spectrum of HD 100546 we detect the 13CO ν = 1 – 0, as well as in the 12CO ν = 1 – 0 to ν = 5 – 4 vibrational bands. All lines are spectrally resolved and have a FWHM of 15 km s⁻¹, irrespective of the transition. A conservative estimate for the HWZI is 17.5 km s⁻¹ corresponding to an inner radius of 3.3 AU. The average line profiles and astrometric signals of the 12CO ν = 1 – 0 to ν = 4 – 3 transitions are shown in Fig. 7 and their values in Table 3. The 4.6 μm continuum emission of the disk around HD 100546 is resolved. Its FWHM of 0.218" ± 0.002" corresponds to 11 ± 1 AU, after correcting for the FWHM of the PSF. The SPP is very asymmetric, i.e. the negative shift of the blue wing in the SPP is larger than the positive SPP of the red wing. This asymmetry is also seen in the FWHM and is more pronounced for the lower vibrational transitions.

5. A Simple Model

A powerful method to determine the extent of the emitting gas is to model the kinematics of the line- and astrometric profiles. We construct the average profiles for the ν = 1 – 0 to ν = 4 – 3 vibrational bands in Figs. 5 and 7, and fit these to the model shown in Fig. 4. In this model, the gas is in Keplerian orbit in a flat disk with known inclination and PA around a star with known stellar mass and distance. Further, the intensity of the emission decreases as I(R) = I₀₉(R/R₀)⁻α, with I₀ the intensity at the inner radius R₀, and R the radial distance from the star (Carmona et al., 2007). The astrometric signal is consequently determined by fitting a Gaussian to the simulated PSF. Because the stellar parameters are all well constrained, the free parameters are α, R₀, and R_out. α describes the change in intensity of the CO emission as function of radius, where α = 2 reflects an intensity directly proportional to the stellar radiation field assuming the surface density has no radial dependence. We compare the line profile, the SPP and the FWHM with their simulated counterparts for the parameter-space 1AU ≤ R₀ ≤ 29AU, 25AU ≤ R_out ≤ 200AU and 0 ≤ α ≤ 5.5, and quantify the best fit by minimizing the reduced χ². We estimate the variance on the data from the con-
Fundamental CO emission from Herbig stars, and gaps therein, have been detected in more circumstellar disks, e.g. in SR 21 with an $R_{\text{in}}$ of 7.6 ± 0.4 AU for the 12CO (v=1-0) line (Pontoppidan et al., 2008), while there is an inner dust component at $\leq 0.5$ AU, followed by a dust gap up to 18 AU (Brown et al., 2007). Goto et al. (2006) find an inner clearing in the CO emission at 11 ± 2 AU around HD 141569, and suggest the cavity most likely is cleared by photo-evaporation working together with viscous accretion. Brittain et al. (2007) observe the same star, find CO emission starting between 6 and 15 AU, and explain the emission as due to UV fluorescence. Blake & Boogert (2004) study a sample of 5 Herbig stars, not known to have disk gaps or inner holes, and find that the CO gas traces both collisional driven emission in the inner (< 0.5 - 1 AU) disk, as well as CO gas excited by resonance fluorescence (R < 50-100 AU). These observations suggest that CO depletion in the inner disk is associated with large disk gaps or inner holes. Finally, in all cases discussed above (except SR 21 for which no data exists), the [OI] emission, when detected, arises at similar or higher velocities than the CO fundamental ro-vibrational emission.

Contrary to the above, the lack of CO emission at small radii in HD 97048 and HD 100546 can not be caused by a gas-free inner disk. [OI] 6300 Å emission, a tracer of OH molecular gas (Acke et al., 2005) in the surface layers of flared disks, and induced by the stellar UV field, has been observed for both stars by
Fig. 5. *Top* The average line profiles of the HD 97048 $v = 1 \rightarrow 0$, $v = 2 \rightarrow 1$ and $v = 3 \rightarrow 2$ transitions, created with 5 ($5 < J < 7$), 12 ($5 < J < 14$) and 7 ($14 < J < 21$) lines respectively, are shown with the black histogram. The blue wing of the $v = 1 \rightarrow 0$ transition is contaminated by telluric CO lines, shown with dots. Over plotted is the 6300 Å [OI] emission (red dotted line), and the model fit to the CO emission with the solid red line. The corresponding, continuum corrected (section 3, the uncorrected SPP is shown with the dotted line), average astrometric signal. The model is over plotted in red. *Middle* The average spatial FWHM in black, with the model over plotted in red.

**Fig. 6.** $\chi^2_{\text{red}}$ surfaces for the best fitting disk model for, from left to right, the line profile, SPP and FWHM of the $v = 2 \rightarrow 1$ transition of HD 97048. The minimal $\chi^2_{\text{red}}$ for each fit is given in the panels, and the 1, 2 and 3 $\sigma$ lines represent $\geq 68\%$, $\geq 90\%$ and $\geq 99\%$ confidence levels. The best fit parameters for each panel are shown by a red circle. To better show the $\chi^2_{\text{red}}$ surface landscape we have overplotted the $\chi^2_{\text{red}} = 5, 10$ (left two panels, $\chi^2_{\text{red}} = 1.5, 1.7$ in the right panel) contours.

Acke & van den Ancker (2006), with 80% of the emission originating between 0.8 < R < 20 AU. The SEDs of our targets also display large NIR excesses, typical for hot dust close to the star (Acke & van den Ancker, 2004; Bouwman et al., 2003). In the cases of HD 97048 and HD 100546, the absence of CO emission at Small radii can thus not be attributed to the paucity of matter in the inner disk, nor to the absence of UV flux. Disk holes and/or large gaps therefore are not the only cause of CO depletion in the inner disk. Here, we suggest that CO is efficiently destroyed in the inner disk. A possible mechanism is the photo dissociation of CO and the subsequent formation of simple organic species such as C2H2, HCN and CH4 (Agundez et al., 2008). Further analysis (e.g. observations of hot transitions of carbon-bearing molecules) is needed to determine the reason why the CO gas is not detected in the inner regions, and may offer us a glimpse of the molecular chemistry of these inner disks.

The similar line widths and shapes for all vibrational states in HD 97048, as well as the similar continuum corrected SPP, suggest a common excitation mechanism and location. The best fit parameter for the radial dependency of the CO emission intensity $\alpha = 3.0$ can guide models that aim to explain the highly excited vibrational line emission, such as UV fluorescence. The similar line shape, continuum corrected SPP and FWHM of the higher ($v = 3 \rightarrow 2$ and $v = 4 \rightarrow 3$) vibrational transitions in HD 100546 also suggest a common excitation mechanism and location, but the $v = 2 \rightarrow 1$ and especially $v = 1 \rightarrow 0$ vibrational transitions deviate significantly. The asymmetry in the data on HD 100546, and the poor fit to the model, point to a more complicated disk geometry, in agreement with finding of Bouwman et al. (2003) and Acke & van den Ancker (2006).

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References