Ground-based observations of the $\beta$ Cephei CoRoT main target HD 180642: abundance analysis and mode identification

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ABSTRACT

The known $\beta$ Cephei star HD 180642 was observed by the CoRoT satellite in 2007. From the very high-precision light curve, its pulsation frequency spectrum could be derived for the first time (Degroote and collaborators). In this paper, we obtain additional constraints for forthcoming asteroseismic modeling of the target. Our results are based on both extensive ground-based multicolour photometry and high-resolution spectroscopy. We determine $T_{\text{eff}} = 24\,500\pm1000\,$ K and $\log g = 3.45\pm0.15$ dex from spectroscopy. The derived chemical abundances are consistent with those for B stars in the solar neighbourhood, except for a mild nitrogen excess.

Key words. stars: oscillations – stars: early-type – stars: individual: HD 180642 – stars: abundances

1. Introduction

The $\beta$ Cephei stars are a homogeneous group of oscillating B0–B3 stars that have been studied as a class for more than a century. Stankov & Handler (2005) compiled an overview of the observational properties of this group of stars. The oscillations of $\beta$ Cephei stars are explained in terms of the x mechanism that operates on iron-peak elements (Dziembowski & Pamyatnykh 1993; Gautschy & Saio 1993). Given that mainly low-degree low-order pressure and gravity modes are observed, these stars are good potential targets for in-depth asteroseismic studies of the interior structure of massive stars.

In principle, these studies would allow us to calibrate the evolutionary models across the entire main sequence, although most $\beta$ Cephei stars seem to be observed while in the second part of their main-sequence lifetime (Stankov & Handler 2005). Moreover, theory predicts the occurrence of stellar oscillations in B stars after the main sequence (Saio et al. 2006) and oscillations in such stars have indeed been found (Lefever et al. 2007).

Asteroseismic studies have been accomplished for 5 $\beta$ Cephei stars only (see Aerts 2008 for a review). These studies were based on intensive ground-based observing campaigns and led, for the first time, to constraints on the core overshoot and internal rotation of massive B-type stars. Despite these successes, we clearly need to increase the number of asteroseismically determined models for these types of stars. This will become es-
Table 1. Logbook of the photometric observations of HD 180642.

<table>
<thead>
<tr>
<th>Observatory</th>
<th>N</th>
<th>AT</th>
<th>HJD begin</th>
<th>HJD end</th>
<th>filter(s)</th>
<th>( f_i )</th>
</tr>
</thead>
<tbody>
<tr>
<td>La Silla</td>
<td>20</td>
<td>186.7</td>
<td>50562.8</td>
<td>50749.6</td>
<td>UB1, BB2, VG</td>
<td>5.4870(2)</td>
</tr>
<tr>
<td>La Palma</td>
<td>171</td>
<td>1183.8</td>
<td>52421.6</td>
<td>53605.4</td>
<td>UB1, BB2, VG</td>
<td>5.48693(1)</td>
</tr>
<tr>
<td>SPMO</td>
<td>113</td>
<td>5.3</td>
<td>53192.7</td>
<td>53198.0</td>
<td>u, v, b, y</td>
<td>5.494(3)</td>
</tr>
<tr>
<td>KNO</td>
<td>222</td>
<td>359.2</td>
<td>53205.4</td>
<td>53564.6</td>
<td>u, v, b, y</td>
<td>5.48693(3)</td>
</tr>
<tr>
<td>KO CCD</td>
<td>84</td>
<td>319.2</td>
<td>53260.3</td>
<td>53579.5</td>
<td>Johnson V</td>
<td>5.49(1)</td>
</tr>
<tr>
<td>ASAS CCD</td>
<td>261</td>
<td>26.1</td>
<td>53921.4</td>
<td>53947.5</td>
<td>Johnson V</td>
<td>5.4858(3)</td>
</tr>
<tr>
<td>Strömgren</td>
<td>363</td>
<td>2797.6</td>
<td>51979.9</td>
<td>54777.5</td>
<td>Johnson V</td>
<td>5.486915(7)</td>
</tr>
</tbody>
</table>

The number N of datapoints retained for the analysis, the total timespan AT (in days) and the available filters are given. The last column gives the value of the dominant frequency (see Sect. 2.2) expressed in \( \text{d}^{-1} \) with its error between brackets. The HJD is given with respect to \( HJD_0 = 2400000.0 \). Strömgren refers to merged SPMO and SNO Strömgren data. The merged Geneva U, B and V and u, v, b, y data, are denoted as U, B, V data, respectively. The V+ dataset corresponds to the \( V \) data to which the KO Johnson \( V \) data and ASAS Johnson \( V \) data are also added.

2. Photometry

2.1. Data description

The star was monitored with the 1.2m Mercator telescope (MER) at La Palma in the framework of a long-term photometric monitoring programme dedicated to pulsating stars. For our target, we completed observing campaigns of Geneva seven-colour (UB1, BB2, VG) high-precision photometry during two seasons (2002, 2005) with 80-day and 57-day durations, respectively. The star was observed a few times every clear night. In total, we used 171 7-colour Geneva measurements with a time span of 1184 days. Aerts (2000) had already obtained 20 datapoints of the star with the same instrument attached to the 0.7m Swiss telescope at La Silla.

With the goal of gathering additional measurements with a higher sampling per night, HD 180642 was observed with the twin Danish six-channel uvby filter photometers attached to the 90cm telescope at Sierra Nevada Observatory (SNO), Spain, and at the 1.5 m telescope at San Pedro Mártir Observatory (SPMO), Mexico. All the data were collected in the four Strömgren uvby filter set. At SNO, a total of 222 high quality datapoints were collected during 1 night in July 2004, 5 nights in June 2005, and 6 clear nights in July 2005. At SPMO, our target was observed during 4 nights in July 2004, resulting in a total of 113 data points.

Johnson \( V \) measurements were also obtained with the 50 cm telescope at the Pizskéstétő Mountain Station of Konkoly Observatory (KO) during 2 nights in September 2004 and 3 nights in July 2005. Since the \( V \) data showed many scattered points, we removed all data deviating by more than 0.05 mag from the fit with the dominant frequency (see Sect. 2.2). We retained a total of 84 datapoints for the analysis. The comparison star used at all sites was HD 181414 (\( V = 7.068 \text{ mag. A2} \)).

During a total of 5 nights in July and August 2006, HD 180642 was also observed in Johnson \( V \) filter with a Princeton Instruments VersArray 1300B CCD detector attached to the 1m RCC telescope (\( f/13.5 \)) at KO. The data were binned such that the total integration time of the binned datapoints was about 3 minutes. In total, we retained 261 datapoints. Since the same comparison stars were monitored not every night in the CCD frame, we reconstructed relative magnitude differences using common bright targets.

Finally, we used the Johnson \( V \) ASAS publically available photometry from Pigulski & Pojmanski (2008). We do not in-
Fig. 1. Phase diagrams for the ultraviolet, blue, and visual merged datasets U, B, V+ (dots) of HD 180642 folded according to the dominant frequency $f_1 = 5.48694 \text{d}^{-1}$. The full gray lines are leastsquares fits including $f_1$, 2$f_1$ and 3$f_1$.

Fig. 2. SCARGLE periodograms calculated from the combined Geneva $U$ and Strömgren $u$ data. The uppermost panel shows the spectral window of the data. All subsequent panels show the periodograms at different stages of prewhitening. From top to bottom: periodogram of the observed light variations, of the data prewhitened with $f_1$ and its two harmonics 2$f_1$ and 3$f_1$, of the data subsequently prewhitened with $f_2^p$, and of the data subsequently prewhitened with $f_3^p$. The three frequencies found are indicated by dashed gray lines. The light gray lines indicate the 4 signal-to-noise (S/N) level. It is computed as the average amplitude over a frequency interval with a width of 2 d$^{-1}$ in an oversampled SCARGLE periodogram obtained after final prewhitening. The highest peak in the third panel corresponds to a one-day alias (1+$f_2^p$) of the true frequency $f_2^p$ (see text).

methods produced identical results within their respective uncertainties, so we list only the outcome obtained for the SCARGLE periodograms. In all datasets and filters, we confirm the dominant frequency. The value of $f_1$ found in each separate dataset is given in the last column of Table 1, where the uncertainty in the last digit is given in parentheses. This frequency uncertainty was calculated to be $\sigma_{f_1} = \sqrt{(6)\sigma_{\text{std}}^2 / \pi \sqrt{N} A_f \Delta T}$ (Montgomery & O'Donoghue 1999), where $\sigma_{\text{std}}$ is the standard deviation of the final residuals, $A_f$ the amplitude of the frequency $f_1$, and $\Delta T$ the total timespan of the observations.

To increase the frequency precision, we combined datasets for given filters whenever possible. We merged the SPMO and SNO Strömgren data (termed Strömgren hereafter), and constructed ultraviolet, blue, and visual ground-based light curves (U,B,V hereafter) by merging the Geneva U and Strömgren $u$, B and $v$, and V and $y$ data, respectively. To combine the data, we calculated the preliminary frequency solution of the individual subsets, checked that the amplitudes of the dominant mode in the Geneva/Strömgren U/u, B/v and V/y are the same within the errors, and realigned the subsets at the same mean brightness level.

We also constructed a more extensive visual band light curve (termed V+ hereafter) by also adding the KO Johnson $V$ data and the ASAS Johnson $V$ data to the Geneva V and Strömgren $y$ data. The results for the dominant frequency of these merged sets are also listed in Table 1. The refined value of the dominant frequency for each of U, B, V, and V+ is $f_1 = 5.48694 \text{d}^{-1}$.

A harmonic fit to the merged U, B, V light curves is shown in Fig. 1 and infers the prominent presence of harmonics of $f_1$. The harmonics of $f_1$ were also derived by a frequency analysis of the residual light curves after prewhitening with $f_1$. We found that 4$f_1$ no longer has a significant amplitude.

In the periodograms of the residuals of the individual and merged datasets, the highest peak occurs at low frequency between 5 and 9 d$^{-1}$. In the individual datasets, none of the candidate frequencies reaches a high enough amplitude to be significant, where we adopt the criterion derived empirically by Breger et al. (1993) and studied further by Kuschnig et al. (1997), who studied the variability of HST guide-stars. They concluded that a peak reaching at least 4 times the noise level in the periodograms has only a 0.1% chance of having been produced by noise. These results were confirmed by De Cat & Cuypers (2003) for the case of g-mode pulsations studied with single-site data. The noise
Table 2. Amplitudes and phases of the least-squares fit to the merged U, B and V+ light curves of HD 180642.

<table>
<thead>
<tr>
<th>Freq (d⁻¹)</th>
<th>U &amp; u Amp (mmag)</th>
<th>Phase (rad)</th>
<th>B &amp; v Amp (mmag)</th>
<th>Phase (rad)</th>
<th>V &amp; y &amp; V Amp (mmag)</th>
<th>Phase (rad)</th>
</tr>
</thead>
<tbody>
<tr>
<td>f₁ = 5.48694</td>
<td>82.7(7)</td>
<td>2.884(9)</td>
<td>48.1(5)</td>
<td>2.88(1)</td>
<td>38.1(5)</td>
<td>2.86(1)</td>
</tr>
<tr>
<td>2 f₁ = 10.97388</td>
<td>8.2(7)</td>
<td>1.48(9)</td>
<td>6.6(5)</td>
<td>1.52(8)</td>
<td>5.3(5)</td>
<td>1.4(1)</td>
</tr>
<tr>
<td>3 f₁ = 16.46082</td>
<td>4.3(7)</td>
<td>2.0(2)</td>
<td>4.3(5)</td>
<td>2.1(1)</td>
<td>3.6(5)</td>
<td>1.9(2)</td>
</tr>
<tr>
<td>f² = 0.30818</td>
<td>4.9(7)</td>
<td>0.9(2)</td>
<td>3.1(5)</td>
<td>1.1(2)</td>
<td>1.6(5)</td>
<td>0.9(4)</td>
</tr>
<tr>
<td>f³ = 7.36673</td>
<td>4.5(7)</td>
<td>5.4(2)</td>
<td>3.1(5)</td>
<td>5.4(2)</td>
<td>2.2(5)</td>
<td>5.7(3)</td>
</tr>
</tbody>
</table>

Errors in units of the last digit are given in brackets. The amplitude is expressed in millimag, and the phase is given in radians. The adopted reference epoch for $\varphi = 0.0$ corresponds to HJD 2447963.2564. For each combined dataset the total variance reduction of the quintu-periodic model and the root mean square (rms) of the residuals is reported.

In Section 3.3, we used the frequencies found in Table 2 to prewhiten the light curves of U, B, and V+ data. The results show that the central peak is seen in the ground-based data and not in the CoRoT light curve. The frequency values found in the space photometry are $f² = 0.29917(9)$ d⁻¹ and $f³ = 7.3586(2)$ d⁻¹. In the CoRoT periodogram, the non-adiabatic computations were performed on two grids of theoretical models for which the adopted metal mixtures, were chosen instead of OPAL opacity tables (Iglesias & Rogers 1996). Grey atmosphere models were used instead of Kurucz atmospheric models. No overshooting was assumed in grid 1 whereas different values of the overshooting parameter (between 0 and 0.5 local pressure scale heights) were explored.

2.3. Mode identification

To identify the modes, we applied the usual method of photometric amplitude ratios, using the formalism of Dupret et al. (2003). This technique has already been successfully applied to the pressure modes in $\beta$ Cephei stars (e.g. Aerts et al. 2006) and to the gravity modes in slowly pulsating B stars (De Cat et al. 2007).

We computed the non-adiabatic eigenfunctions and eigenfrequencies required for modes with $\ell \leq 4$ by using the code called MAD (Dupret et al. 2001). Ground-based data are currently not precise enough to detect higher degree modes since cancellation effects leave a very tight signal that is below the detection threshold. The non-adiabatic computations were performed on two grids of theoretical models for which the adopted input physics were different, as described below. All the models were computed with the Code Liégeois d’Évolution Stellaire (CLES, Scuflaire et al. 2008).

The first grid (grid 1 hereafter) was adopted from De Cat et al. (2007) and we refer to this paper for details of the input physics. The second grid (grid 2 hereafter) was computed using different opacity tables and different atmosphere models. For grid 2, OP opacity tables (Seaton 2005), assuming Asplund et al. (2005) metal mixtures, were chosen instead of OPAL opacity tables (Iglesias & Rogers 1996). Grey atmosphere models were used instead of Kurucz atmospheric models. No overshooting was assumed in grid 1 whereas different values of the overshooting parameter (between 0 and 0.5 local pressure scale heights) were explored.
considered in grid 2. Finally, grid 2 covers higher masses (up to 18 $M_\odot$) than grid 1 (up to 15 $M_\odot$).

The computation of grid 2 was motivated by the following reasons. Since the paper by De Cat et al. (2007), work on the modeling of $\delta$ Cephei stars pointed out that OP opacity tables can account for the excitation of modes observed in studied objects, while OPAL opacity tables fail (Miglio et al. 2007, Briquet et al. 2007, Dziembowski & Pamyatnykh 2008). In addition, closer agreement between theory and observations is also achieved if overshooting is present (Briquet et al. 2007, Aerts 2008).

We obtained the same conclusions regardless of the grid considered. In the following, we present our mode identification outcome for grid 2. We used the models whose stellar parameters are compatible, within a 3$\sigma$ error, with those determined from spectroscopy in Sect. 3.2: $T_{\text{eff}} = 24,500 \pm 1,000$ K and $\log g = 3.45 \pm 0.15$.

The outcome for $f_1$ is displayed in the top panel of Fig. 3. It confirms that the dominant mode is a radial one as already ascertained by Aerts (2000). To identify the $\ell$-values of the two observed low-amplitude modes, we then considered models only in the 3$\sigma$ error box that also fit $f_1$ as a radial mode. The amplitude ratios of $f_p$ and $f'_p$ are plotted in the middle and bottom panels of Fig. 3, respectively. We conclude that $f_p$ is unambiguously identified as a g-mode pulsation with $\ell = 3$. Moreover, $f_p$ corresponds to $\ell = 0$ or 3, with a preference for $\ell = 0$.

It is noteworthy that a low frequency such as 0.30818 d$^{-1}$, which corresponds to a high-order g mode, is not yet a common feature observed for most of the $\delta$ Cephei stars. However, it is certainly not exceptional. For several of the $\delta$ Cephei stars that have indeed been the subject of intensive observing campaigns, low frequencies have been found, e.g., in 19 Mon (Balona et al. 2002: 0.17019 d$^{-1}$), v Eri (Handler et al. 2004: 0.43218 d$^{-1}$), and 12 Lac (Handler et al. 2006: 0.35529 d$^{-1}$).

3. Spectroscopy

3.1. Data description

In addition to the photometric datasets, we have at our disposal 262 useful (S/N $>$ 80) high-resolution spectra gathered...
with three instruments. HD 180642 was observed during 15
nights (25 June–4 July; 16–20 July 2007) with FEROS at the
2.2m ESO/MPI telescope at La Silla and during 12 nights
(13–24 June 2007) with SOPHIE at the 1.93m telescope at
OHP in the framework of the CoRoT ground-based Large
Programme (Uytterhoeven & Poretti 2007; Uytterhoeven et al.
2008). Additionally, 22 Aurèlie spectra at the 1.52m telescope
(OHP) and 11 FEROS spectra were already taken in 2005 and
2006 (14, 17, 19, and 22 June 2005 at OHP; 25–26 May 2006 at
ESO). An overview and logbook of the spectroscopic observa-
tions are given in Table 3.

We carefully calculated the integration times that we selected
to ensure a S/N ratio of about 200 near the wavelength 4550 Å,
which is the position of a prominent Si III triplet well suited to
our line-profile variability study (Aerts & De Cat 2003). These
silicon lines are indeed sufficiently strong without being much
affected by blending. Moreover, they are dominated by temper-
bature broadening, such that the intrinsic profile can be modelled
by a Gaussian. This simplifies the modeling of the line-profile
variations for mode identification purposes (De Ridder et al.
2002).

We reduced the FEROS spectra using an improved version
of the standard FEROS pipeline, written in MIDAS, developed
by Rainer (2003). The main improvements of this pipeline con-
cern the blaze and flat-field correction of the spectra, by using an
accurate definition of the blaze function extracted from a well-
exposed spectrum of a hot star. The SOPHIE spectra were ex-
tracted and automatically reduced in real-time by a reduction
package adapted from HARPS. Since there is no pipeline re-
duction available for the Aurèlie spectrograph, we used standard
reduction procedures with IRAF. After a correction to the he-
liocentric frame, the spectra were manually normalised using a
cubic spline fit.

3.2. Abundance analysis

The non-local thermodynamic equilibrium (NLTE) abund-
cances of helium and the dominant metals have been cal-
culated using the latest versions of the line-formation codes
DETAIL/SURFACE (Butler & Giddings 1985; Giddings 1981)
and plane-parallel, fully line-blanketed atmospheric models with
a solar helium abundance (Kurucz 1993). Our analysis is based
on a mean FEROS spectrum created by coadding 11 individual
exposures obtained in May 2006. All exposures were transferred
to the laboratory rest frame prior to this operation. Curve-of-
growth techniques were used to determine the abundances using
the equivalent widths of a set of unblended lines. The reader is
referred to Morel et al. (2006) for complete details of the
methodology used to derive the elemental abundances.

3.2.1. Atmospheric parameters

A standard, iterative scheme was used to self-consistently de-
rive the atmospheric parameters solely by spectroscopic means:
\(T_{\text{eff}}\) was determined from the Si II/III/IV ionisation balance, \(\log g\)
from fitting the collisionally-broadened wings of the Balmer
lines, and the microturbulent velocity, \(\xi\), by requiring that the
abundances corresponding to the O I I features were independent
of the line strength. We obtain: \(T_{\text{eff}}=24,500\pm1000\ \text{K}, \log g=3.45\pm0.15\ \text{dex}\) and \(\xi=12\pm3\ \text{km s}^{-1}\). As can be seen in
Fig. 4, satisfactory fits are obtained for all Balmer lines con-
sidered. Within the uncertainties, identical values are obtained
when analysing these data with the unified code FASTWIND

![Fig. 4. Comparison between the synthetic (dashed line, \(T_{\text{eff}}=24,500\ \text{K}, \log g=3.45\ \text{dex}\)) and observed (solid line) Balmer line profiles. The synthetic spectra have been convolved with a rotational profile with \(v=44\ \text{km s}^{-1}\). No attempts have been made to fit the other spectral lines.](image)

3.2.2. Elemental abundances

The abundances computed for \(T_{\text{eff}}=24,500\ \text{K}, \log g=3.45\ \text{dex}\), and \(\xi=12\ \text{km s}^{-1}\) are given in Table 4, and compared with the
standard solar mixture of Grevesse & Sauval (1998) and values
derived from time-dependent, three-dimensional hydrody-
Table 4. Mean NLTE abundances (on the scale in which log $\epsilon$(H)=12) and total 1-$\sigma$ uncertainties (in brackets), in units of the last digits.

<table>
<thead>
<tr>
<th></th>
<th>HD 180642</th>
<th>B stars</th>
<th>Sun 1-D</th>
<th>Sun 3-D</th>
</tr>
</thead>
<tbody>
<tr>
<td>He/H(4)</td>
<td>0.088(18)</td>
<td>-0.089</td>
<td>0.085(1)</td>
<td>0.085(2)</td>
</tr>
<tr>
<td>log e(C)(9)</td>
<td>8.21(10)</td>
<td>-8.20</td>
<td>8.52(6)</td>
<td>8.39(5)</td>
</tr>
<tr>
<td>log e(N)(21)</td>
<td>8.00(19)</td>
<td>-7.79</td>
<td>7.92(6)</td>
<td>7.78(6)</td>
</tr>
<tr>
<td>log e(O)(25)</td>
<td>8.53(14)</td>
<td>-8.58</td>
<td>8.83(6)</td>
<td>8.66(5)</td>
</tr>
<tr>
<td>log e(Ne)(6)</td>
<td>7.86(16)</td>
<td>-8.07</td>
<td>8.06(6)</td>
<td>7.84(6)</td>
</tr>
<tr>
<td>log e(Mg)(1)</td>
<td>7.34(20)</td>
<td>-7.48</td>
<td>7.58(5)</td>
<td>7.53(9)</td>
</tr>
<tr>
<td>log e(Al)(3)</td>
<td>6.22(15)</td>
<td>-6.08</td>
<td>6.47(7)</td>
<td>6.37(6)</td>
</tr>
<tr>
<td>log e(Si)(7)</td>
<td>7.19(19)</td>
<td>-7.20</td>
<td>7.55(5)</td>
<td>7.51(4)</td>
</tr>
<tr>
<td>log e(S)(4)</td>
<td>7.10(34)</td>
<td>-7.19</td>
<td>7.33(11)</td>
<td>7.14(5)</td>
</tr>
<tr>
<td>log e(Fe)(21)</td>
<td>7.34(21)</td>
<td>-7.36</td>
<td>7.50(5)</td>
<td>7.45(5)</td>
</tr>
<tr>
<td>[N/C]</td>
<td>-0.21(22)</td>
<td>-0.41</td>
<td>-0.60(9)</td>
<td>-0.61(8)</td>
</tr>
<tr>
<td>[N/O]</td>
<td>-0.53(24)</td>
<td>-0.79</td>
<td>-0.91(9)</td>
<td>-0.88(8)</td>
</tr>
<tr>
<td>Z</td>
<td>0.0099(16)</td>
<td>-0.0108</td>
<td>0.0172(12)</td>
<td>0.0124(7)</td>
</tr>
</tbody>
</table>

The number of spectral lines used is given in brackets in the first column. For comparison purposes, we provide typical values found for early B dwarfs in the solar neighbourhood (Morel 2009, and references therein; B stars), the standard solar composition of Grevesse & Sauval (1998; Sun 1-D), and updated values derived from three-dimensional hydrodynamical models (Asplund et al. 2005; Sun 3-D). We define [N/C] and [N/O] as $\log e$(N)/(C) and $\log e$(N)/(O), respectively. The metallicity, Z, is given in the last row. To compute this quantity, we assumed the abundances of Grevesse & Sauval (1998) for the trace elements not under study.

Table 5. Abundance and metallicity differences when adopting the microturbulent velocity yielded by the N ii lines ($\xi$=7 km s$^{-1}$), instead of the adopted value estimated from the O ii features ($\xi$=12 km s$^{-1}$).

<table>
<thead>
<tr>
<th></th>
<th>from 12 to 7 km s$^{-1}$</th>
</tr>
</thead>
<tbody>
<tr>
<td>$\Delta e$</td>
<td></td>
</tr>
<tr>
<td>$\Delta$He/H</td>
<td>+0.025</td>
</tr>
<tr>
<td>$\Delta$log e(C)</td>
<td>+0.04</td>
</tr>
<tr>
<td>$\Delta$log e(N)</td>
<td>+0.13</td>
</tr>
<tr>
<td>$\Delta$log e(O)</td>
<td>+0.16</td>
</tr>
<tr>
<td>$\Delta$log e(Fe)</td>
<td>+0.06</td>
</tr>
<tr>
<td>$\Delta$log e(S)</td>
<td>+0.02</td>
</tr>
<tr>
<td>$\Delta$log e(Al)</td>
<td>+0.11</td>
</tr>
<tr>
<td>$\Delta$log e(E)</td>
<td>+0.10</td>
</tr>
<tr>
<td>$\Delta$log e(Al)</td>
<td>+0.15</td>
</tr>
<tr>
<td>[N/C]</td>
<td>+0.09</td>
</tr>
<tr>
<td>[N/O]</td>
<td>-0.03</td>
</tr>
<tr>
<td>Z</td>
<td>+0.0027</td>
</tr>
</tbody>
</table>

3.3. Frequency analysis

To increase the S/N of the spectra and improve the detection of the line-profile variations, we combined several line profiles as follows. First, cross-correlated profiles were computed using the leastsquares deconvolution (LSD) method (Donati et al. 1999) by means of a mask consisting of lines from the VALD database (Piskunov et al. 1995; Ryabchikova et al. 1999; Kupka et al. 1999) for $T_{\text{eff}} = 24,000$ K and log $g = 3.50$. All elements were taken into account, except He and H. This resulted in combining about 570 lines for the FEROS and SOPHIE spectra and only 78 lines for the Aurelie spectra, leading to average S/N ratios of 600, 400, and 90, respectively.

Second, to avoid the effects of the different behaviour of the individual lines during the pulsation cycle, which was ignored in the individual method, we computed mean spectra of sufficiently strong and unblended lines originating from the same element and ionisation state. The averaged spectrum was computed for He i lines, N ii lines, O ii lines, and Si iii lines. The use of LSD profiles or different elements did not reveal additional periodicities. In what follows, we describe only our frequency analysis of the combination of the two deepest silicon lines of the Si iii triplet around 4567 Å, of an average S/N ratio of 350.

To perform our frequency analysis, we used the software package FAMIAS (Zima 2008). We first examined the first three moments $< v^1 >$, $< v^2 >$ and $< v^3 >$ (see Aerts et al. 1992 for a definition) of the combined silicon lines. Since the line profiles move significantly because of pulsation, the integration limits for computing the moments were dynamically chosen by sigma clipping to avoid the noisy continuum (see Zima 2008).

Harmonics up to the 5th of the dominant frequency are of significant amplitude in $< v^1 >$. After removing these, we do find four additional significant frequencies in the residuals of $< v^1 >$. All frequencies retained, except the last one, have an amplitude larger than the 4 S/N-level, computed in the residuals, in a 2 d$^{-1}$ interval centered on the frequency of interest. The reason for retaining the last peak is that it is a one-day-alias of a combination frequency. In our frequency solution, we use the value of the combination frequency. The detected frequency values and their corresponding amplitudes and phases of the leastsquares sine fits to $< v^1 >$ are listed in Table 6. These frequencies were also found in the CeRoT light curve (Degroote et al. 2009), albeit not with exactly the same values. We already discussed a prob-

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1 FAMIAS has been developed in the framework of the FP6 European Coordination Action HELAS – http://www.helas-eu.org/

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Fig. 5. Phase diagram of the equivalent width of the \( \text{Si} \)\textsc{iii} 4552 Å line, and of the first three velocity moments, for \( f_1 = 5.4869 \text{ d}^{-1} \). The dashed lines represent our best model for the moment variations (see Sect. 3.4 for more explanations).

Table 6. Frequencies, amplitudes, and phases of the least squares sine fits to the observed first moment.

<table>
<thead>
<tr>
<th>ID</th>
<th>Frequency (d(^{-1}))</th>
<th>Amplitude (km s(^{-1}))</th>
<th>Phase (rad)</th>
</tr>
</thead>
<tbody>
<tr>
<td>( f_1 )</td>
<td>5.4869(1)</td>
<td>38.8(2)</td>
<td>3.139(6)</td>
</tr>
<tr>
<td>2 ( f_1 )</td>
<td>10.9738(1)</td>
<td>3.0(2)</td>
<td>5.45(8)</td>
</tr>
<tr>
<td>3 ( f_1 )</td>
<td>16.4607(1)</td>
<td>4.7(2)</td>
<td>3.30(5)</td>
</tr>
<tr>
<td>4 ( f_1 )</td>
<td>21.9476(1)</td>
<td>3.6(2)</td>
<td>5.30(7)</td>
</tr>
<tr>
<td>5 ( f_1 )</td>
<td>27.4345(2)</td>
<td>1.7(2)</td>
<td>0.8(1)</td>
</tr>
<tr>
<td>( f_2 )</td>
<td>8.4079(2)</td>
<td>1.5(2)</td>
<td>4.2(1)</td>
</tr>
<tr>
<td>( f_3 )</td>
<td>0.3046(2)</td>
<td>1.1(2)</td>
<td>2.0(2)</td>
</tr>
<tr>
<td>( f_4 )</td>
<td>7.1037(3)</td>
<td>1.0(2)</td>
<td>2.7(2)</td>
</tr>
<tr>
<td>( f_1 + f_3 )</td>
<td>13.8948(3)</td>
<td>0.8(2)</td>
<td>6.3(3)</td>
</tr>
</tbody>
</table>

The error estimates in units of the last digit are given in brackets. The adopted reference epoch for \( \psi = 0.0 \) corresponds to HJD 2447963.2564.

A first indicator of the shock propagation is the evolution in the line profile, which exhibits a doubling phenomenon when the front shock enters the line forming region, following the Schwarzschild mechanism (Schwarzschild 1952). The second indicator is a phase lag between line forming regions, due to the finite shock front velocity, which is present in progressive waves. For HD 180642, no line-doubling behaviour is seen in the profiles, but it could be hidden by the presence of the low-amplitude modes. No phase lag is also observed between variations associated with different ions. We therefore conclude that the presence of shocks in HD 180642 is not established, nor can be rejected. The observation of the star during a few consecutive cycles would allow us to prevent the beat phenomenon with the non-radial modes and identify the same behaviour as observed for BW Vul and \( \sigma \) Sco.

The behaviour of \( v^2 \) (see middle panel of Fig. 5) also deserves some comments. It clearly cannot be described by a double sine, as linear theory predicts for radial pulsators (Aerts et al. 1992). In addition, the maxima (minima) at phases 0.05 and 0.45 (at phases 0.25 and 0.80) are different. This difference cannot be attributed to a non-linear velocity at the stellar surface only but it reflects that temperature variations also take place. By taking into account temperature effects, \( v^2 \) can be satisfactorily modelled, as shown in Sect. 3.4. This is because the second moment depends on the thermal width of the local profile, which varies with temperature. Similar behaviour, but less prominent, was already observed for \( \delta \) Ceti (Aerts et al. 1992) and \( \xi^1 \) CMa (Saesen et al. 2006).

The temperature variations caused by the large-amplitude radial mode are also revealed in the equivalent width of the spectral lines. As shown in Fig. 5, the equivalent width of the \( \text{Si} \textsc{iii} \) 4552 Å line varies with the dominant frequency and two

\[ \text{Fig. 5. Phase diagram of the equivalent width of the Si} \textsc{iii} 4552 \text{Å line, and of the first three velocity moments, for } f_1 = 5.4869 \text{ d}^{-1}. \]
significant harmonics. We point out that this strong EW variability is not common among pulsating B-type stars. We refer to both De Ridder et al. (2002) and De Cat (2002) for typical examples of β Cephei stars and slowly pulsating B stars, respectively. However, this behaviour is certainly not unique, and is observed in several stars with a high-amplitude mode. For β Cephei stars, it is also the case for BW Vul (Mathias et al. 1998), σ Sco (Mathias et al. 1991), δ Ceti (Aerts et al. 1992), ν Eridani (Aerts et al. 2004), and the prototype β Cephei itself (Telting et al. 1997, Catanzaro & Leone 2008). All these targets are monoperiodic radial pulsators or multiperiodic pulsators with a highly dominant radial mode, as is the case for HD 180642. A different case concerns the β Cephei star 12 Lac (Desmet et al. 2009). This latter object, which also is of relatively high amplitude, shows an equivalent width that varies with its two dominant modes but the highest amplitude mode is identified as an (l, m) = (1, 1).

A frequency search was carried out for the spectra by means of a two-dimensional Fourier analysis available in FAMIAS. Once the frequencies are selected, a non-linear multi-periodic least-squares fit of a sum of sinusoidals is computed with the Levenberg-Marquardt algorithm. This fitting is applied for every bin of the spectrum separately according to the formula $Z + \sum_i A_i \sin(2\pi(f_i + \phi_i))$, where Z is the zeropoint, and $A_i$, $f_i$, and $\phi_i$ are the amplitude, frequency, and phase of the $i$-th frequency, respectively. The amplitude and phase distributions across the combined silicon lines are shown, in Fig. 6, for the dominant radial mode and, in Fig. 7, for the other frequencies detected.

We emphasize that the 2D frequency analysis detected one additional frequency, which was also observed in the CoRoT photometry ($f_0 = 6.3248\, \text{d}^{-1}$), compared to the 1D analysis of the moments, which are integrated quantities. Our Fourier analysis detected a one-day-alias of $f_0$ and we needed the CoRoT data to pinpoint the true periodicity. Additional frequencies were also discovered by a 2D frequency search in the case of several other β Cephei stars (see Telting et al. 1997 for β Cephei; Schrijvers et al. 2004 for ε Centauri; Briquet et al. 2005 for θ Ophiuchi), pointing out the necessity of performing this kind of analysis, in addition to a frequency search in the radial velocity or higher-order moments.
3.4. Mode identification

The moment method (Briquet & Aerts 2003) and the Fourier parameter fit (FPF) method (Zima 2006) are spectroscopic mode identification techniques appropriate to identifying both the degree \( \ell \) and the azimuthal order \( m \) of the oscillations of main-sequence pulsators hotter than the Sun. Several successful applications, in particular to \( \beta \) Cephei stars, are available in the literature (e.g. Briquet et al. 2005 for \( \theta \) Oph, Mazumdar et al. 2006 for \( \beta \) CMa, Desmet et al. 2009 for 12 Lac).

As explained and illustrated in Zima (2006), the FPF technique, based on mono-mode line profiles, cannot be applied in case of a large pulsation velocity relative to the projected rotational velocity, i.e., if the radial velocity amplitude is above 0.2 \( v \sin i \), as for HD 180642. Instead of using amplitude and phase distributions for mono-mode line profiles, we thus considered the same quantities but computed them for a huge grid of multiperiodic line profile time series. We fixed the dominant mode as radial and its velocity field at the stellar surface was modelled as the superposition of the velocity fields of five linear radial modes.

To model the local temperature and surface gravity variations, we adopted the following empirical approach. At each time of observation, the equivalent width (EW) and the width \( \sigma \) of an intrinsic Gaussian profile were chosen to match the observed EW variations and second moment variations. We point out that this model assumes that the temperature and gravity variations are due solely to the dominant mode, which is a reasonable assumption in view of the very low amplitudes of the other modes. The theoretical moment variations in this simple model are shown in Fig. 5. The amplitude and phase distributions are displayed in Fig. 6. Our modeling of the non-linear radial mode reproduces the observed behaviour quite well, especially for the first four harmonics.
Table 7. The bestfit solutions of the spectroscopic mode identification for the mode with frequency 8.4079 d⁻¹ determined by the adapted discriminant Σ, based on the definition in Aerts (1996).

| (ℓ, |m|) | ℓ | v sin i | A_p | v_r,max | v_t,max | Σ |
|------|----|---------|-----|--------|--------|----|
| (3,2) | 3 | 81.5 | 38.5 | 77.35 | 30.4 | 3.0 | 0.93 |
| (3,1) | 3 | 30.5 | 39.5 | 25.69 | 11.4 | 1.3 | 4.13 |
| (1,1) | 1 | 14.5 | 29.9 | 31.73 | 11.0 | 0.4 | 4.41 |
| (2,1) | 2 | 14.0 | 25.2 | 25.14 | 9.7 | 0.7 | 4.42 |
| (2,2) | 2 | 47.5 | 38.4 | 21.71 | 8.4 | 0.6 | 4.48 |
| (3,3) | 3 | 66.5 | 23.8 | 35.52 | 14.8 | 1.7 | 4.59 |
| (3,0) | 3 | 23.0 | 1 | 26.90 | 20.0 | 1.7 | 35.63 |
| (1,0) | 1 | 83.5 | 1 | 49.63 | 24.2 | 1.0 | 35.65 |
| (2,0) | 2 | 63.0 | 1 | 37.88 | 23.8 | 1.4 | 35.65 |
| (0,0) | 0 | 90.0 | 0 | 7.72 | 2.2 | 0.0 | 36.45 |
| (4,0) | 4 | 3.0 | 1 | 101.88 | 86.0 | 9.1 | 49.94 |

The inclination angle i is expressed in degrees; v sin i is the projected rotational velocity, expressed in km s⁻¹; A_p is the amplitude of the radial part of the pulsation velocity, expressed in km s⁻¹; and v_r,max and v_t,max are, respectively, the maximum radial and tangential surface velocity due to the mode, expressed in km s⁻¹.

To identify the wavenumbers (ℓ, m) of the low-amplitude modes, we added their pulsational velocity field to the one of the radial mode modelled as described above, and, using a χ² value as in Zima (2006), we searched for the solutions that fit the observed amplitude and phase distributions computed from multi-mode line profiles best. Unfortunately, no good fit could be achieved for any of the modes. Next, we used the moment method for multiperiodic stars (version of Briquet & Aerts 2003) also adapted to adequately model the dominant radial mode, but, again, the outcome was inconclusive.

Therefore, we tried to identify the modes individually, following the version of the moment method of Aerts (1996). Since the width σ of the intrinsic profile is not constant in time, in contrast to the assumption in Aerts (1996), we defined a slightly different discriminant, in which the amplitudes involving σ are not taken into account. We also omitted the constant term of the second moment because it contains the contribution of all the modes. Therefore, the discriminant used is the one of Aerts (1996) but only with the amplitudes denoted by C, D, F, and G (see Aerts 1996) being taken into account.

We varied the free parameter of the projected rotation velocity, v sin i, from 1 to 40 km s⁻¹ with a step 0.1 km s⁻¹, and the inclination angle of the star, i, from 3 to 90° with a step 0.5°. For each combination (ℓ, m, i), the amplitude of the mode A_p (given by definition of Briquet & Aerts 2003) was not a free parameter but was determined to fit the observed amplitude of the first moment, within its error. For the low-frequency mode, we fixed ℓ to be 3, as derived from the photometry. Finally, in the criterion in Sect. 2.5 in Briquet & Aerts (2003) excluded ℓ ≥ 4, except (ℓ, m) = (4,0), for the mode with frequency 8.4079 d⁻¹. For the other modes, this criterion limited neither ℓ nor m.

We obtained a safe mode identification for the mode with frequency 8.4079 d⁻¹ only. The discriminant values for the other modes turned out to be too high for us to be confident about their outcome. The discriminant values for the identified mode are given in Table 7. It does not allow us to determine the sign of m.

However, prograde and retrograde modes can be distinguished by means of a blue-to-red and red-to-blue respectively descent of the phase in the profile. From Fig. 7, we deduce that the mode is prograde and denote it with a positive m-value. From Table 7, we conclude that the mode is unambiguously identified as (ℓ, m) = (3,2).

By using the discriminant Σ as a weight, we constructed histograms for ℓ, v sin i, and the equatorial rotational velocity v_eq, as in Mazumdar et al. (2006). We computed them by considering only solutions with (ℓ, m) = (3,2). For ℓ and v sin i, the distributions are almost flat. However, v_eq can be constrained (see Fig. 8). By calculating a weighted mean and standard deviation, we obtained v_eq = 38±15 km s⁻¹.

4. Summary

The β Cephei star HD 180642 was observed by the CoRoT satellite during a run of 156 days in 2007. The space white-light photometry revealed the rich frequency spectrum of the star (Degroote et al. 2009). In the present study, we have provided additional information about the target, based on both ground-based multi-colour photometry and high-resolution spectroscopy. We placed our object in the (T_eff, log g) diagram. In addition, we derived the chemical abundances of several elements as well as the metallicity of HD 180642. Finally, we placed constraints on the identification of some modes.

From the spectroscopic data, we determined T_eff = 24 500±1000 K and log g = 3.45±0.15 dex. In addition, a detailed NLTE abundance analysis showed that the considered abundance values are compatible with values found for B stars in the solar neighbourhood, except for a mild nitrogen excess. This nitrogen overabundance was also discovered for several other β Cephei stars (Morel et al. 2006, 2008). The deduced metallicity for HD 180642 is Z = 0.0099±0.0016 (or 0.0126±0.0016 depending on the choice of microturbulence).

In the photometry, three pulsation frequencies were found to be significant: 5.48694 d⁻¹, 0.30818 d⁻¹, and 7.36673 d⁻¹. The first mode is highly dominant with an amplitude in the U-filter about 15 times larger than that of the other modes. By means of the method of photometric amplitude ratios, we unambiguously identified the degree of the first two modes as ℓ = 0 and ℓ = 3. For the frequency 7.36673 d⁻¹, two possibilities remained, namely ℓ = 0 or 3.
In the radial velocity measurements, the dominant radial mode is highly non-linear, presents a so-called stillstand and has a peak-to-peak amplitude of ~30 km s⁻¹. This behaviour is typical of the presence of shocks. However, our dataset did not allow us to prove (or definitely refute) their existence in HD 180642.

We attempted mode identification for several low-amplitude modes found in the spectroscopy. In our first trials, we used multi-periodic versions of the moment method and FFP method. Both techniques were adapted and implemented to adequately model the dominant radial mode but we failed for the low-amplitude modes. However, a successful outcome for one of the modes was achieved by means of the discriminant of Aerts (1996), which we adapted slightly to our case. This discriminant identifies the modes independently. In the presence of a highly dominant mode, it is thus more robust than the multi-periodic discriminant of Briquet & Aerts (2003). We found the frequency 8.079 d⁻¹ to correspond to a (ℓ, m) = (3, 2). Finally, we derived a value for the equatorial rotational velocity of $V_{eq} = 38 ± 15$ km s⁻¹.

All these observational constraints, together with the CoRoT results, will be used to compute stellar models of HD 180642. This will be presented in a forthcoming paper.

Acknowledgements. We thank our colleagues from the Institute of Astronomy of Leuven University who contributed to the gathering of the photometric data. The FEROS data were obtained as part of the ESO programme 077.D-0311 (PI: K. Uytterhoeven) as well as part of the ESO Large Programme 178.D-0361 (PI: E. Poretti). Degroote, P., Briquet, M., Catala, C., et al. 2009, A&A, submitted.


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