

## Observation of the suppression of the flux of cosmic rays above $4 \times 10^{19}$ eV

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The energy spectrum of cosmic rays above  $2.5 \times 10^{18}$  eV, derived from 20,000 events recorded at the Pierre Auger Observatory, is described. The spectral index  $\gamma$  of the flux,  $J \propto E^{-\gamma}$ , at energies between  $4 \times 10^{18}$  eV and  $4 \times 10^{19}$  eV is  $2.69 \pm 0.02$  (stat)  $\pm 0.06$  (syst), steepening to  $4.2 \pm 0.4$  (stat)  $\pm 0.06$  (syst) at higher energies, consistent with the prediction by Greisen and by Zatsepin and Kuz'min.

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We report a measurement of the energy spectrum of cosmic rays showing that the flux is strongly suppressed above  $4 \times 10^{19}$  eV. This is in accord with the 1966 prediction of Greisen [1] and of Zatsepin and Kuz'min [2] (GZK) that the spectrum should steepen around  $5 \times 10^{19}$  eV as cosmic rays from cosmologically distant sources suffer energy losses when propagating through the cosmic microwave radiation. With an exposure twice that of HiRes [3] and 4 times that of AGASA [4], our evidence supports the recent report of the former.

The Pierre Auger Observatory, located near Malargüe (Argentina) at 1400 m a.s.l., is used to measure the properties of extensive air showers (EAS) produced by the highest-energy cosmic rays. At ground level the electrons, photons and muons of

EAS can be detected using instruments deployed in a large surface array. Additionally, as EAS move through the atmosphere, ultra-violet light is emitted from nitrogen excited by charged particles. This *fluorescence light* is proportional to the energy deposited by the shower along its path [5]. The Observatory uses 1600 water-Cherenkov detectors, each containing 12 tonnes of water, viewed by three 9" photomultipliers, to detect the photons and charged particles. The surface detectors are laid out over 3,000 km<sup>2</sup> on a triangular grid of 1.5 km spacing and is overlooked by 4 fluorescence detectors. Each fluorescence detector (FD), located on the perimeter of the area, houses 6 telescopes. EAS detected by both types of detector are *hybrid events* and play a key role in the analysis. The field of view of each telescope is 30° in azimuth, and 1.5° – 30° in elevation. Light is focused on a camera containing 440 hexagonal pixels, of 18 cm<sup>2</sup>, at the focus of a 11 m<sup>2</sup> mirror. The design and status of the Observatory are described in [6, 7]. Between 1 Jan 2004 to 31 Aug 2007 the numbers of telescopes increased from 6 to 24 and of surface detectors from 154 to 1388. The analysis of data from this period is described.

A cosmic ray of 10<sup>19</sup> eV arriving vertically typically produces signals in 8 surface detectors. Using relative timing, the direction of such an event is reconstructed with an angular accuracy of about 1° [8]. Signals are quantified in terms of the response of a surface detector (SD) to a muon travelling vertically and centrally through it (a *vertical equivalent muon* or VEM). Calibration of each SD is carried out continuously with 2% accuracy [9]. The signals are fitted in each event to find the VEM size at 1000 m,  $S(1000)$  [10]. The uncertainty in every  $S(1000)$  is found, accounting for statistical fluctuations of the signals, systematic uncertainties in the assumption of the fall-off of signal with distance and the shower-to-shower fluctuations [8]. Above 10<sup>19</sup> eV the uncertainty in  $S(1000)$  is about 10%.

The longitudinal development of EAS in the atmosphere is measured using the fluorescence detectors. The light produced is detected as a line of illuminated pixels in one or more FT cameras. The positions of these pixels and the arrival time of the light determine the shower direction. The signal, after correcting for attenuation due to Rayleigh and aerosol scattering, is proportional to the number of fluorescence photons emitted in the field of view of the pixel. Cherenkov light produced at angles close to the shower axis can be scattered towards the pixels: this contamination is accounted for [11]. A Gaisser-Hillas function [12] is used to reconstruct the shower profile which provides a measurement of the energy of the EAS deposited in the atmosphere. To derive the primary energy, an estimate of the missing energy carried into the ground by muons and neutrinos must be made based on assumptions about the mass of cosmic rays and of the appropriate hadronic model. For a primary beam that is a 50/50 mixture of protons and iron, simulations of showers with the QGSJET01 model indicate a correction of 10% [13]. The systematic uncertainty is 4% [14].

Detailed understanding of the fluorescence emission is needed for accurate energy determination. The absolute fluorescence yield in air at 293 K and 1013 hPa from the 337 nm band is  $5.05 \pm 0.71$  photons/MeV of energy deposited [15]. The wavelength and pressure dependence of the yield adopted follow [16]. Systematic uncertainties in the FD energy measurement have been estimated. Measurements, made in combination with the fluorescence detectors, are used to measure the quality and transmission properties of the atmosphere. In particular, the vertical aerosol optical depth (VAOD) profile [17] is found every 15 min by observing the light scattered from a centrally-located laser of an energy equivalent to a few 10<sup>19</sup> eV at 355 nm [18] yielding an hourly average. The average correction to  $E_{FD}$  from the VAOD measurement is +5% at  $3 \times 10^{18}$  eV rising to +18% at  $5 \times 10^{19}$  eV, reflecting the increase of the average distance of such events from an FD. The absolute calibration of the telescopes is measured every few months and is constantly adjusted using relative calibrations [19]. The largest uncertainties are in the absolute fluorescence yield (14%), the absolute calibration of the telescopes (10%) and the reconstruction method (10%). Systematic uncertainties from atmospheric aerosols, the dependence of the fluorescence spectrum on temperature and on humidity are each at the 5% level [7, 20]. These uncertainties are independent and added in quadrature give 22% for  $E_{FD}$ .

The fluorescence detectors are operated on clear, moonless nights limiting the duty cycle to 13%. Showers detected by both the surface array and the FD (hybrid events) are more precisely reconstructed than surface array- or FD-only events [7] and are essential to the evaluation of systematic uncertainties. The hybrid events have an angular accuracy that improves from 0.8° at  $3 \times 10^{18}$  eV to 0.5° above 10<sup>19</sup> eV. The surface array, with its near 100% duty cycle, gives the large sample used here. The comparison of the shower energy, measured using fluorescence, with the  $S(1000)$  for a subset of hybrid events is used to calibrate the energy scale for the array.

Only events with zenith angles less than 60° are used here. Candidate showers are selected on the basis of the topology and time compatibility of the triggered detectors [21]. The SD with the highest signal must be enclosed within an *active hexagon*, in which all six surrounding detectors were operational at the time of the event. Thus it is guaranteed that the intersection of the axis of the shower with the ground is within the array, and that the shower is sampled sufficiently to make reliable measurements of  $S(1000)$  and of the shower axis. From the analysis of hybrid events, using only the fall of the signal size with distance, these criteria result in a combined trigger and reconstruction efficiency greater than 99% for energies above about  $3 \times 10^{18}$  eV; at  $2.5 \times 10^{18}$  eV it is 90% [22]. The sensitive area has been calculated from the total area of the hexagons active every second.

The decrease of  $S(1000)$  with zenith angle arising from the attenuation of the shower and from geometrical effects is quantified by applying the *constant integral intensity cut method* [23], justified by the approximately isotropic flux of primaries. An energy estimator for each event, independent of  $\theta$ , is  $S_{38^\circ}$ , the  $S(1000)$  that EAS would have produced had it arrived at the median zenith angle, 38° [24]. Using information from the fluorescence detectors the energy corresponding to each  $S_{38^\circ}$  can be estimated almost entirely from data except for assumptions about the missing energy. The energy calibration is obtained from a subset of high-quality hybrid events, where the geometry of an event is determined from the times recorded at an FD, supplemented by

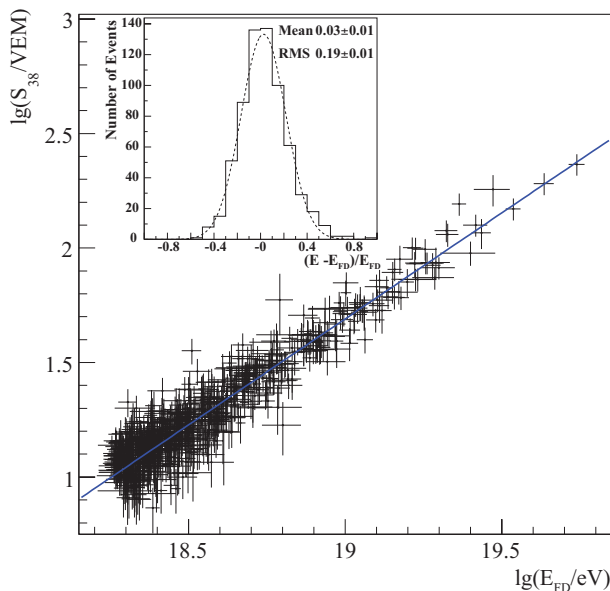


FIG. 1: Correlation between  $\lg S_{38^\circ}$  and  $\lg E_{\text{FD}}$  for the 661 hybrid events used in the fit. The full line is the best fit to the data. The fractional differences between the two energy estimators are inset.

the time at the SD with the highest signal, if it is within 750 m from the shower axis [25, 26]. It is also required that a reduced  $\chi^2$  is less than 2.5 for the fit of the longitudinal profile and that the depth of shower maximum be within the field of view of the telescopes. The fraction of the signal attributed to Cherenkov light must be less than 50%. Statistical uncertainties in  $S_{38^\circ}$  and  $E_{\text{FD}}$  were assigned to each event: averaged over the sample these were 16% and 8%, respectively.

The correlation of  $S_{38^\circ}$  with  $E_{\text{FD}}$  is shown in Fig. 1, together with the least-squares fit of the data to a power-law,  $E_{\text{FD}} = a \cdot S_{38^\circ}^b$ . The best fit yields  $a = (1.49 \pm 0.06 \text{ (stat)} \pm 0.12 \text{ (syst)}) \times 10^{17} \text{ eV}$  and  $b = 1.08 \pm 0.01 \text{ (stat)} \pm 0.04 \text{ (syst)}$  with a reduced  $\chi^2$  of 1.1.  $S_{38^\circ}$  grows approximately linearly with energy. The energy resolution, estimated from the fractional difference between  $E_{\text{FD}}$  and the derived SD energy,  $E = a \cdot S_{38^\circ}^b$ , is shown inset. The root-mean-square deviation of the distribution is 19%, in good agreement with the quadratic sum of the  $S_{38^\circ}$  and  $E_{\text{FD}}$  statistical uncertainties of 18%. The calibration accuracy at the highest energies is limited by the number of events: the most energetic is  $\sim 6 \times 10^{19} \text{ eV}$ . The calibration at low energies extends below the range of interest.

The energy spectrum based on  $\sim 20,000$  events is shown in Fig. 2. Statistical uncertainties and 84% confidence-level limits are calculated according to [27]. Systematic uncertainties on the energy scale due to the calibration procedure are 7% at  $10^{19} \text{ eV}$  and 15% at  $10^{20} \text{ eV}$ , while a 22% systematic uncertainty in the absolute energy scale comes from the FD energy measurement. The possibility of a change in hadronic interactions or in the mean primary mass above  $6 \times 10^{19} \text{ eV}$  will be addressed with more data. In photon-initiated showers the value of  $S(1000)$  is 2-3 times smaller than for nuclear primaries, so that a large photon flux would change the spectrum. However, a limit to the photon-flux of 2% above  $10^{19} \text{ eV}$  exists [29].

The spectrum is fitted by a smooth transition function with the suppression energy of  $4 \times 10^{19} \text{ eV}$  defined as that at which the flux falls below an extrapolated power law by 50%. To examine the spectral shape at the highest energies, we fit a power-law function between  $4 \times 10^{18} \text{ eV}$  and  $4 \times 10^{19} \text{ eV}$ ,  $J \propto E^{-\gamma}$ , using a binned likelihood method [30]. A power-law is a good parameterization: the spectral index obtained is  $2.69 \pm 0.02 \text{ (stat)} \pm 0.06 \text{ (syst)}$  (reduced  $\chi^2 = 1.2$ ), the systematic uncertainty coming from the calibration curve. The numbers expected if this power-law were to hold above  $4 \times 10^{19} \text{ eV}$  or  $10^{20} \text{ eV}$ , would be  $167 \pm 3$  and  $35 \pm 1$  while 69 events and 1 event are observed. The spectral index above  $4 \times 10^{19} \text{ eV}$  is  $4.2 \pm 0.4 \text{ (stat)} \pm 0.06 \text{ (syst)}$ . A method which is independent of the slope of the energy spectrum is used to reject a single power-law hypothesis above  $4 \times 10^{18} \text{ eV}$  with a significance of more than 6 standard deviations [30], a conclusion independent of the systematic uncertainties currently associated with the energy scale.

In Fig. 2 the fractional differences with respect to an assumed flux  $\propto E^{-2.69}$  are shown. HiRes I data [3] show a softer spectrum where our index is 2.69 while the position of suppression agrees within the quoted systematic uncertainties. The AGASA data are not displayed as they are being revised [31]. The change of spectral index indicated below  $4 \times 10^{18} \text{ eV}$  will be discussed elsewhere.

To summarize, we reject the hypothesis that the cosmic-ray spectrum continues with a constant slope above  $4 \times 10^{19} \text{ eV}$ , with a significance of 6 standard deviations. In a previous paper [32], we reported that sources of cosmic rays above  $5.7 \times 10^{19} \text{ eV}$  are extragalactic and lie within 75 Mpc. Taken together, the results suggest that the GZK prediction of spectral steepening may have been verified. A full identification of the reasons for the suppression will come from knowledge of the mass spectrum in

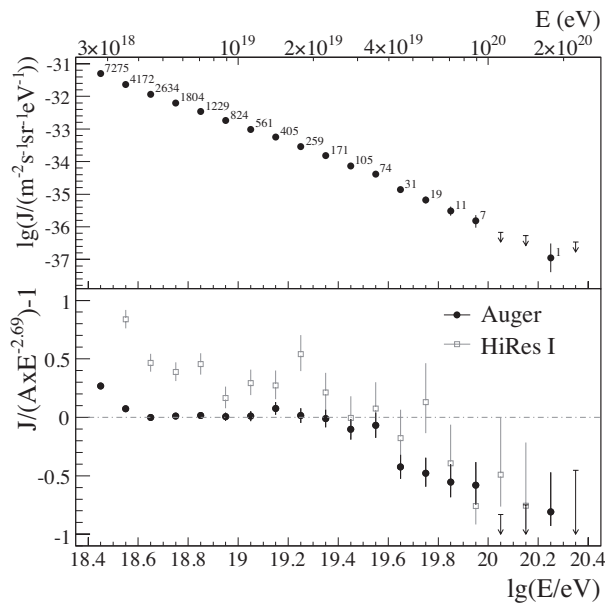


FIG. 2: Upper panel: The differential flux  $J$  as a function of energy, with statistical uncertainties. Data are listed at [28]. Lower Panel: The fractional differences between Auger and HiRes I data [3] compared with a spectrum with an index of 2.69.

the highest-energy region and from reductions of the systematic uncertainties in the energy scale which will allow the derivation of a deconvolved spectrum.

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