Summary of the variability processing and analysis results

The processed variables (LPVs), rotation modulation (BY Dra-type) stars, and Cepheids and RR Lyrae stars detected in the Large Magellanic Cloud (LMC), as were various visualisation tools to interpret the data. Many variable sources were published in DR1 (Eyer et al. 2017) for the published DR2 results are discussed in full length in the following dedicated papers:

- DR2: Variable stars in the Gaia colour-magnitude diagram: Gaia Collaboration et al. (2018a)
- DR2: All-sky classification of high-amplitude pulsating stars: Rimoldini & et al. (2018)
- DR2: Specific characterisation and validation of all-sky Cepheids and RR Lyrae stars: Clementini et al. (2018)

Many variable sources as were published in DR1 (Eyer et al. 2017) Clementini et al. (2016). Moreover, it covers the whole sky. The variability analysis for this release focused primarily on large-amplitude variable stars and rotation modulated stars. Other classes of variable stars are aimed to be introduced in later releases. We note that eclipsing binaries were identified as well, but will be treated separately and only delivered in future releases. We emphasize that specific selection criteria were applied to each variability type to limit contamination. Completeness within each variability type was not aimed at for DR2.

The various variability type-specific methods that were used for the published DR2 results are discussed in full length in the following dedicated papers:

- DR2: Variable stars in the Gaia colour-magnitude diagram: Gaia Collaboration et al. (2018a)
- DR2: All-sky classification of high-amplitude pulsating stars: Rimoldini & et al. (2018)
- DR2: Specific characterisation and validation of all-sky Cepheids and RR Lyrae stars: Clementini et al. (2018)
2. Data

The results in DR2 are based on the first 22 months of Gaia data, taken between 25 July 2014 and 23 May 2016. For the variability analysis we made use of the photometric time series in the G, G\(_{\text{P}}\), and G\(_{\text{BP}}\) bands, using field-of-view (FoV) averaged transit photometry. An exception to this was the analysis of short-timescale variables, for which some per-CCD G-band photometry was also analysed, but not published. The photometry is described in detail in Evans et al. (2018) and Riello et al. (2018). We emphasise that owing to the tight DR2 data-processing schedule, we only had access to a preliminary version of the DR2 astrometric solution (Lindgren et al. 2018), on which our results are based. No radial velocity data or astrophysical parameters were available during our processing.

2.1. Relative completeness

As a consequence of the Gaia scanning law, the sky-coverage is rather non-uniform for the 22 months of DR2 data, resulting in sky-coverage gaps when selecting a certain minimum number of FoV transits. In our processing we used cuts of \(\geq 2\), \(\geq 12\), and \(\geq 20\) G-band FoV transits, which we discuss in Sect. 3. Figure 1 shows the effect of these cuts on the sample of all available time series at the time of processing. Assuming that all sources with \(\geq 2\) FoV transits are bona fide sources, these figures show that a cut of \(\geq 12\) FoV transits results in a relative completeness of 80%, and a cut of \(\geq 20\) FoV transits results in a relative completeness of only 51%. It can generally be expected that these relative completenesses are an upper limit to the absolute (sky) completeness, which is assessed in detail in the follow-up variability papers mentioned in the introduction. Arenou et al. (2018) conclude that the catalogue is mostly complete, in the absolute sense, roughly between magnitudes 7 and 17, and incomplete beyond. For a first-order relative completeness assessment in certain general Galactic directions, we introduce the following regions and compute the Gaia source count in each of them:

- **Galactic plane** \(30^\circ < l < 330^\circ\) & \(|b| < 15^\circ\)
- **High Galactic latitudes** \(|b| > 45^\circ\), excluding a 10\(^\circ\) radius around the SMC
- **LMC region** within a radius of 7.5\(^\circ\) from Galactic coordinates \(l = 280.47^\circ, b = -32.89^\circ\)
- **SMC region** within a radius of 2.5\(^\circ\) from Galactic coordinates \(l = 302.81^\circ, b = -44.33^\circ\)

We repeat this for the (approximate) sky footprints of various external catalogues, list the results in Table 1 and display the regions in Fig. 1. It shows that a (maximum) completeness of between 70–85\% for \(\geq 12\) G-band FoV transits and a relative completeness of between 50–70\% for \(\geq 20\) G-band FoV transits can be expected. The bottom panel of Fig. 1 shows that various sky regions are not or only partially sampled for \(\geq 20\) FoV: the Galactic centre region, for example, has a relative completeness of only 14\%.

2.2. Outliers and calibration errors

Various issues are visible in the current photometric data. Some of them are caused by instrumental effects, while others are sky related. Improved calibration strategies and careful flagging of these events are being designed and will be implemented in the chain of processing systems that convert the raw data into calibrated epoch photometry for the next Gaia Data Release. Some of these issues are also described in Evans et al. (2018):

- hot pixel columns are not yet treated,
- poor background estimates exist for some observations,
- observations close to bright sources can have biased background estimates or might even be spurious detections due to diffraction spikes,
- spatially close sources can have scan-angle direction dependent outliers,
- isolated sources can still be affected through overlap of the two FoVs.

It is expected that a certain fraction of observation affected by these issues has not been flagged and hence users should be aware of such unflagged outliers in the published data. We discuss variability flagging in more detail in Sect. 4.1.

In addition to the flagging of observations, the Gaia data processing consortium has applied various filters that remove sources that are affected by specific calibration issues; see Gaia Collaboration et al. (2018) for more details.

3. DR2 variability data processing and analysis

The general variability data processing has been described in detail in our DR1 paper (Eyer et al. 2017), to which we refer here. We here only summarise the main aspects that are specific to the DR2 release. More exhaustive information is provided in the DR2 documentation.

The variability processing is summarised in Fig. 2 and the number of sources involved at various levels of the processing is indicated in the diagram. The processing followed two main paths: one starting from 1607 million sources with \(\geq 2\) G-band FoV transits (hereafter geq2), and the other from 826 million
sufficient quality. The chain of consecutive operators was used to remove FoV transits by flagging and filtering them when they did not meet the required criteria, that is, when their associated SOS modules: SOS Cep and RRL (Clementini et al. 2016) and SOS LPV (Mowlavi et al. 2018).

The second path leads to the majority of the pipelines and SOS results, as described in Sect. 4.3. It had two different variability detection methods. The first is a ‘general’ variability detection, defined by a classifier trained on known cross-matched constant objects1 from OGLE4 (Soszyński et al. 2012), Hipparcos (ESA 1997), the Sloan Digital Sky Survey (SDSS) (Ivezic et al. 2007), and classified constants, as well as many additional catalogues for variables (Rimoldini & et al. 2018). The second consists of multiple dedicated variability detection methods: one tuned to detect short-timescale variables and another to rotation modulation effects, which were subsequently followed by their associated SOS modules: SOS short timescale (Roelens et al. 2018) and rotation modulation (Lanzafame et al. 2018), analysing these sources in more detail. The general variability detection was followed by a generic multi-harmonic Fourier modelling based on the periodogram peak-frequency of an unweighted least-squares period search. Various classification attributes were derived from this model and the time-series statistical parameters, on which the classifier for the geq20 path was trained. The classified RR Lyrae stars, Cepheids, and LPVs were subsequently fed into the corresponding modules SOS Cep&RRL and SOS LPV, being their second input source. All modules of the variability pipeline include validation rules to ensure that the output values are within acceptable ranges, which allowed us to identify issues early on in the processing.

An accounting of the number of sources published in the classification and SOS modules is given in Fig. 3. It shows that of the 550,737 unique variable sources, 363,969 (66%) appear in the classification results (Sect. 4.3) and 390,529 (71%) in SOS (Sect. 4.4). For several variability types, 203,681 sources (37%) appear in both, while the rest of the sources appear only in either classification (160,288, i.e. 29%) or SOS results (186,768, i.e. 34%). Note that the number appearing in both classification and SOS tables (203,681) is 721 more than the sum of overlapping types shown in Fig. 3 because these are cases of different assigned types in classification and SOS tables, which are omitted from the figure.

Depending on the type of object, there might only exist either a classification or SOS result. SOS results have generally lower contamination than the classification results and contain various type-specific atmospheric parameters that can be used to obtain more detailed information (such as the period), although this information might be available for a less sky-uniform sample due to the minimum number of G-band FoV observations of 12 or 20, depending on the SOS type. The classification contains only a label and ‘score’ and has generally higher contamination rates, but it contains larger samples that are (more) uniformly distributed over the sky due to the minimum number of G-band FoV observations, which are 5 for RR Lyrae, 6 for δ Scuti and SX Phoenicis, 9 for Cepheids, and 12 for LPVs. These differences between classification and SOS in contamination rate and sample size are part of the pipeline design in which classification results feed into most of the SOS modules, as shown in Fig. 2, and discussed in more detail in the DR1 paper Eyer et al. (2017).

4. Results

4.1. Observation filtering and operators

The variability processing makes use of photometric data provided in units of fluxes (e/s), and then converts them into magnitudes using the magnitude zero-points defined in Evans et al. (2018). In our pipeline, this transformation is done by our GaiaFluxToMagOperator operator. Several additional operators were used to remove FoV transits by flagging and filtering them when they did not meet the required criteria, that is, when they were of insufficient quality. The chain of consecutive operators is described briefly here, for more details see the DR2 documentation:

Table 1: Relative source counts for sources with ≥2 G-band FoV transits, illustrating the decrease in relative completeness for source sets with ≥12 and ≥20 G-band FoV transits, as applied in our processing. The top part specifies the numbers for some generic directions in the Galaxy, as defined in Sect. 2.1 and illustrated in Fig. 1. The bottom part specifies the numbers for several external catalogues, listing the Gaia source counts in their approximate sky footprints (not their source cross-match counts). In parenthesis, we provide the actual Gaia source counts in millions. Estimates of the absolute completeness of the published variable samples are presented in Tables 2 and 3.

<table>
<thead>
<tr>
<th>Direction</th>
<th>≥ 2 FoV</th>
<th>≥ 12 FoV</th>
<th>≥ 20 FoV</th>
</tr>
</thead>
<tbody>
<tr>
<td>ALL</td>
<td>100% (1607)</td>
<td>80% (1283)</td>
<td>51% (826)</td>
</tr>
<tr>
<td>Galactic plane</td>
<td>100% (778)</td>
<td>85% (665)</td>
<td>68% (532)</td>
</tr>
<tr>
<td>Galactic centre</td>
<td>100% (351)</td>
<td>67% (236)</td>
<td>14% (50.0)</td>
</tr>
<tr>
<td>High Galactic latitudes</td>
<td>100% (50.4)</td>
<td>76% (38.4)</td>
<td>48% (24.4)</td>
</tr>
<tr>
<td>LMC region</td>
<td>100% (23.5)</td>
<td>76% (17.9)</td>
<td>59% (13.9)</td>
</tr>
<tr>
<td>SMC region</td>
<td>100% (3.4)</td>
<td>78% (2.6)</td>
<td>66% (2.2)</td>
</tr>
<tr>
<td>Catalina DR2</td>
<td>100% (400)</td>
<td>80% (320)</td>
<td>53% (210)</td>
</tr>
<tr>
<td>Catalina DR1</td>
<td>100% (200)</td>
<td>80% (160)</td>
<td>50% (100)</td>
</tr>
<tr>
<td>OGLE-IV Bulge</td>
<td>100% (305)</td>
<td>66% (200)</td>
<td>10% (29)</td>
</tr>
<tr>
<td>OGLE-IV MCs</td>
<td>100% (33)</td>
<td>77% (25)</td>
<td>62% (20)</td>
</tr>
<tr>
<td>Kepler</td>
<td>100% (8.4)</td>
<td>87% (7.4)</td>
<td>78% (6.6)</td>
</tr>
<tr>
<td>Rebull et al. (2016) (K2 Pleiades)</td>
<td>100% (0.38)</td>
<td>85% (0.32)</td>
<td>66% (0.25)</td>
</tr>
<tr>
<td>Hartman et al. (2010) (Pleiades)</td>
<td>100% (0.87)</td>
<td>84% (0.73)</td>
<td>51% (0.44)</td>
</tr>
</tbody>
</table>

1 The least variable sources found in ftp://ftp.astrouw.edu.pl/ogle/ogle4/GSEP/maps/
– **RemoveNaNNegativeAndZeroValuesOperator** flagged and removed transits with NaN, negative, or zero flux values.
– **RemoveDuplicateObservationsOperator** flagged and removed pairs of transits with too close observation times (essentially one transit from the pair was incorrectly assigned to the source). This operator was responsible for the majority of the transit removals for G magnitudes brighter than ~8, see Mowlavi et al. (2018) for more details.
– **GaiaFluxToMagOperator** converted the Gaia fluxes into magnitude using the Gaia zero-point magnitudes.
– **ExtremeValueCleaning** flagged and removed transits with unrealistically faint magnitudes. The cut values were chosen to be $G \geq 25$, $G_{BP} \geq 24$, and $G_{RP} \geq 22$.
– **ExtremeErrorCleaningMagnitudeDependent** flagged and removed transits with extreme magnitude errors. For $G$, we used both a lower (0.01%) and upper (99.7%) threshold based on the observed distribution of transit magnitude errors (as a function of magnitude), while for $G_{BP}$ and $G_{RP}$, we used only an upper threshold (99.9%).
– **RemoveOutliersFaintAndBrightOperator** flagged and removed the time-series outliers of the source. Different criteria were used to account for the distribution of magnitudes or of the magnitude errors in the time series.

The time series following the above last operator was then the input for the different variability modules, and all transits that were flagged and filtered in the operator chain can be identified in the published time series by the `rejected_by_variability` flag (see Sect. 4.2). The resulting number of selected (i.e. not-rejected) transits can be found in the `num_selected_g_fov/bp/rp` fields of the `vari_time_series_statistics` table. In Figs. 6 and 8 the rejected transits are plotted as crosses. However, some SOS modules (e.g. SOS Cep and RRL, SOS rotation modulation) applied individual stricter conditions that were tuned for their analysis. We note that the SOS short timescale module also used another operator, **RemoveOutlierPerTransitOperator**, to remove CCD outliers per transit.

The sky distribution of the selected number of transits in the $G$, $G_{BP}$, and $G_{RP}$ time series are shown in Fig. 9 for all 550,737 published sources. The distribution of the mean values per sky-pixel is dominated by the scanning law, which is also visible in the large masked regions of the $\geq 25$ FoV panel of Fig. 1. Figure 10 shows a histogram of the time-series lengths. The minimum number of transits in $G$ is 5.

### 4.2. Time series, statistics, and variability flag

The full list of variable stars published in DR2 can be easily identified in the `gaia_source` table by the field `phot_variable_flag`, which is set to VARIABLE, while for all other sources it is set to NOT_AVAILBLE. The time series of all variable stars are published in $G$, $G_{BP}$, and $G_{RP}$. In contrast to DR1, the time series are no longer provided in a separate archive table, but instead are accessible in a Virtual Observatory Table (VOTable) linked via the field `epoch_photometry_url` in the `gaia_source` table. For various reasons, a fraction of the transit observations was not photometrically processed in DR2, resulting in null values. The variability processing flags per FoV transit (see previous section) are provided in the column `rejected_by_variability` of the VOTable epoch photometry. Additional flags from the photometric processing are available for transits in the column `rejected_by_photometry`. Unchanged from DR1 is the expression of observation time in units of barycentric JD (in TCB) in days $\leq 245$ 197.5, see DR2 documentation or Eyer et al. (2017) for more details.

All time series are provided in both flux and magnitude, where the flux is provided with the associated standard error. The magnitude-transformed error value is omitted given its non-symmetric nature. The Vega-system flux to magnitude zero-points are defined in Evans et al. (2018).

Several basic statistical parameters of the cleaned time series are listed in the `vari_time_series_statistics` table. The variable stars will furthermore have an entry in at least one of the `vari_*` tables, see Fig. 3 for their numbers. An overview of time-series statistics per Gaia band and types is provided in Tables 4 and 5. The mean magnitudes provided in `gaia_source` (`phot_g/bp/rp_mean_mag`) differ from the mean magnitudes in `vari_time_series_statistics` (`mean_mag_g_fov/bp/rp`) because they are calculated differently (see Evans et al. 2018) and because the filtering of the light curves was different; furthermore, the `gaia_source` mean photometry can be absent, but present in the `vari_time_series_statistics` table (together with the related time series in `epoch_photometry_url`) because various filters were applied by the photometry pipeline. The sky and median magnitude distributions of these sources are shown in more detail in Figs. 14 and 15.

#### 4.3. Classification

In DR2 we introduce the output of the semi-supervised classifier targeted to identify high-amplitude variable stars over the whole sky, resulting from the geq2 path described in Sect. 5. This is the aim of the `nTransits:2+` classification. Owing to the limited 22 months of data and rejected observations, the number of FoV transits per source can be very small, prohibiting the use of Fourier modelling parameters for a large portion of the sources. Previous studies such as Ivezic et al. (2000) and Herrmann et al. (2016) have shown that it is still possible to identify good RR Lyrae star candidates with just a few observations. We classifier was trained to identify high-amplitude variables of the type of RR Lyrae stars (anomalous RRd, see Sozzi et al. (2016); RRd, RRab, RRc), LPV (Mira-type and semi-regulars), Cepheids (anomalous Cepheids, δ Cepheids, type-II Cepheids), δ Scuti and SX Phoenicis stars, amongst the many other (non-)variable source types that exist. These classes can be found in the `Gaia` archive `vari_classifier_result` table with `best_class_name` = ARRd, RRd, RRAB, RRC, MIRA_SR, ACCEP, CEPP, T2CEP, and DSCT_SXPH, respectively. Each entry has an associated `best_class_score` between 0 and 1. Additional (non-published) classes were included in the prediction types, see Rimoldini et al. (2018) for all details. Descriptions of the classifier and the published classes can be found in the `Gaia` archive `vari_classifier_definition` and `vari_classifier_class_definition` table, respectively.

The sky and magnitude distribution of the published classes, grouped by main type, is shown in the left panels of Figs. 5 and 7 and some time-series examples are shown in Figs. 6 and 7.

2 See [http://www.ivoa.net/documents/VOTable/](http://www.ivoa.net/documents/VOTable/)
3 For tutorials on using this data link, see the Gaia archive help webpage [http://gea.esac.esa.int/archive-help](http://gea.esac.esa.int/archive-help) Time-series bulk download is possible from [http://cdn.gea.esac.esa.int/Gaia/](http://cdn.gea.esac.esa.int/Gaia/)

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4 Practical approximation: $\delta_{flux} \sim (2.5/\ln 10) \cdot \frac{flux_{mean}}{flux_{value}}$

5 Selecting FoV transits with `rejected_by_variability=false`
The published source samples were cleaned to reduce contamination levels caused by data and processing artefacts, as well as by genuine variability of other types. The results and details of this procedure are fully detailed in Rimoldini et al. (2018).

The output of this classifier for sources containing at least 12 measurements in their G-band time series was used as one of the two inputs to the SOS of RR Lyrae stars, Cepheids, and LPV, as shown in Fig. 3 hence accounting for the overlap between classification and SOS results in Fig. 3. Some SOS modules may reclassify the type in the SOS tables, see Sect. 4.6. An overview of the inferred completeness and contamination with respect to several cross-match catalogues of the different types is provided in Table 2.

4.4. Specific Object Studies

In the geq20 path described in Sect. 3, potential variable type candidates are identified in the supervised classification and special variability detection (SVD) parts of the pipeline were analysed in SOS modules, which often derived a model and/or additional astrophysical properties. The SOS modules can also decide whether an object is not of the expected type and as a result not produce any output for it. We recall that a sub-sample of the sources processed in the geq2 path was processed by some SOS modules (see Fig. 4). We present short descriptions of the SOS modules used for DR2 below.

The sky and magnitude distribution of each published type is shown in the right panels of Figs. 3 and 4, and example light curves are presented in Figs. 5, 7, and 8. Because SOS modules can be fed by (partially) independent pipeline-paths, it is not excluded that a source is listed in multiple SOS tables, as discussed in Sect. 4.5 or has a different classification and SOS type, as discussed in Sect. 4.6. Despite extensive validation efforts, the nature of our processing is an automated and statistical characterisation of the sources, hence identified sources that are not yet known in the literature should be considered to be candidates, unless they are specifically validated in one of the SOS articles referenced below.

An overview of the inferred (absolute) sky completeness and contamination with respect to existing catalogues of the different types is provided in Table 3, see sections below and the references therein for more details. Many of the listed reference catalogues have a (very) limited sky coverage, and hence the contamination might be locally biased. To assess contamination in the whole SOS published samples, we performed the following additional visual inspection for each SOS output separately: we took a random sample of 500 sources from the published results, as well as a sky-uniform sample of 500 sources (non-overlapping with the random sample). The former sample follows the general sky distribution as shown in Fig. 4, while the latter achieves an almost sky-uniform sampling and hence also draws samples from very low density regions. The results are summarised below for each SOS result:

- **Cepheids**: 5% of the random sample was visually rejected, of which 3% might be due to binaries or ellipsoidals. Most of this sample is located in the Magellanic Clouds. Of the sky-uniform sample, 15% was visually rejected, of which 11% might be due to binaries or ellipsoidals.

- **RR Lyrae stars**: 9% of the random sample was visually rejected, of which 4% are faint (20.0 < G < 20.7 mag). Of the sky-uniform sample, 12% was visually rejected, of which 5% are faint (19.9 < G < 20.7 mag).

- **Long-period variables**: 0.1% (1) was visually rejected; it looks like a young stellar object, confirmed by its position in the Hertzsprung-Russel (HR) diagram.

- **Short-timescale variability**: 0.3% was visually rejected, and about 24% seem to exhibit short-timescale periodic variability, but this can currently not be confirmed with a high confidence.

- **Rotation modulation**: 0.3% was visually rejected.

These identified contamination rates are compatible with the cross-match sample contamination rates found in Table 3.

4.4.1. Cepheids and RR Lyrae stars

The SOS Cepheids and RR Lyrae module (SOS Cep and RRL, Clementini et al. 2016, 2018) confirmed, characterised, and reclassified in type 140 784 RR Lyrae stars and 9 575 Cepheids of the candidates provided by the classifiers. The module can reclassify the type provided by the classifiers from Cepheid to RR Lyrae star, and vice versa. Amongst the RR Lyrae stars, the following number of sub-types were identified: 98 026 fundamental mode pulsators (RRab), 40 380 first-overtone pulsators (RRc), and 2378 double-mode pulsators (RRd). Amongst the Cepheids, there were 8890 δ Cepheids, 100 anomalous Cepheids, and 585 type II Cepheids, of which 223 were classified as BL Her, 253 as W Vir, and 109 as RV Tau.

A rough estimate of completeness and contamination was obtained by comparing our samples with OGLE catalogues of RR Lyrae stars and Cepheids in the Large Magellanic Cloud (LMC). We selected an area of the LMC centred at α = 83° and δ = −67.5° and extending from 67.5° to 97.5° in right ascension and from −63.5° to −73° in declination. Completeness of our RR Lyrae sample in this area is 64%, and contamination is less than 0.1%. In the same area, the Cepheid sample is 74% complete with 100% purity. A more in-depth description and analyses of the results, as well as the origin of the Bulge completeness rates quoted in Table 3 can be found in Clementini et al. (2018). Of the 599 Cepheids and 2595 RR Lyrae stars that were identified in DR1 (Clementini et al. 2016), Eyer et al. (2017), 533 Cepheids and 2517 RR Lyrae stars are also available in this DR2 release.

4.4.2. Long-period variables

Long-period variables are red giants that are characterised by long periods and variability amplitudes of up to a few magnitudes. They are identified by classification as MIXED_SR. On the basis of data available for DR2, CU7 selected 151 761 LPV candidates that fulfil the selection criteria for LPVs, which included a minimum of 12 points in G, a GBP − GRP colour larger than 0.5 mag, a variability amplitude (quantified with the 5%-95% trimmed range in G) larger than 0.2 mag, and a correlation between G and GBP − GRP higher than 0.5. Of these, 150 757 are published in the classification table, and 89 617 have SOS results. The detailed distribution into these two tables is indicated in Fig. 4. The SOS LPV module processes LPV candidates when their period is longer than 60 days. The steps applied
Table 2: Classification (absolute) completeness and contamination estimates with respect to available cross-matched reference catalogues. For more details on the validation of the classification results, we refer to [Rimoldini & et al. 2018].

<table>
<thead>
<tr>
<th>Group</th>
<th>Class (best_class_name)</th>
<th>Catalogue (and region)</th>
<th>Completeness</th>
<th>Contamination</th>
</tr>
</thead>
<tbody>
<tr>
<td>Cepheids</td>
<td>CEP, ACEP, T2CEP</td>
<td>OGLE-IV (LMC)</td>
<td>72%</td>
<td>&gt;5%</td>
</tr>
<tr>
<td>Cepheids</td>
<td>CEP, ACEP, T2CEP</td>
<td>OGLE-IV (SMC)</td>
<td>68%</td>
<td>&gt;5%</td>
</tr>
<tr>
<td>δ Scuti and SX Phoenicis</td>
<td>DSCT_SXPHE</td>
<td>Catalina (Drake et al. 2014)</td>
<td>99%</td>
<td>&gt;13%</td>
</tr>
<tr>
<td>RR Lyrae stars</td>
<td>RRAB, RRC, RRD, ARRD</td>
<td>OGLE-IV (LMC)</td>
<td>60%</td>
<td>&gt;9%</td>
</tr>
<tr>
<td>RR Lyrae stars</td>
<td>RRAB, RRC, RRD, ARRD</td>
<td>OGLE-IV (SMC)</td>
<td>64%</td>
<td>15%</td>
</tr>
<tr>
<td>RR Lyrae stars</td>
<td>RRAB, RRC, RRD, ARRD</td>
<td>OGLE-IV (Bulge)</td>
<td>49%</td>
<td>44%</td>
</tr>
<tr>
<td>RR Lyrae stars</td>
<td>RRAB, RRC, RRD, ARRD</td>
<td>Catalina (Drake et al. 2014)</td>
<td>63%</td>
<td>&gt;5%</td>
</tr>
<tr>
<td>LPV</td>
<td>MIRA_SR</td>
<td>OGLE-III (LMC+SMC)</td>
<td>40–50%</td>
<td>&lt; 5%</td>
</tr>
<tr>
<td>LPV</td>
<td>MIRA_SR</td>
<td>Gaia DR2 with $\sigma_\omega/\sigma_\alpha &gt; 10$</td>
<td>–</td>
<td>~7%</td>
</tr>
</tbody>
</table>

Notes. (a) Estimates are restricted to sources with peak-to-peak amplitude >0.2 mag. (b) $\sigma_\omega$ and $\sigma_\alpha$ refer to the DR2 parallax and associated error.

Table 3: SOS (absolute) completeness and contamination estimates with respect to cross-matched reference catalogues. Processing paths are based on the number of selected FoV transits, as shown in Fig. 2 and discussed in Sect. 3. See Sect. 4.4 for additional contamination rate estimates for random and sky-uniform samples. For more details on the validation of the SOS type results, see Sect. 4.4 and the references therein.

<table>
<thead>
<tr>
<th>Processing path</th>
<th>SOS Type</th>
<th>Catalogue (and region)</th>
<th>Completeness</th>
<th>Contamination</th>
</tr>
</thead>
<tbody>
<tr>
<td>geq12 &amp; geq20 Cepheids</td>
<td>OGLE-IV (LMC, limited region)</td>
<td>74%</td>
<td>~5%</td>
<td></td>
</tr>
<tr>
<td>geq12 &amp; geq20 Cepheids</td>
<td>OGLE-IV (Bulge)</td>
<td>3%</td>
<td>b</td>
<td></td>
</tr>
<tr>
<td>geq12 &amp; geq20 RR Lyrae stars</td>
<td>OGLE-IV (LMC, limited region)</td>
<td>64%</td>
<td>~9%</td>
<td></td>
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<tr>
<td>geq12 &amp; geq20 RR Lyrae stars</td>
<td>OGLE-IV (Bulge)</td>
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</tr>
<tr>
<td>geq12 &amp; geq20 LPV</td>
<td>OGLE-III (SMC+LMC)</td>
<td>~30%</td>
<td>&lt; few%</td>
<td></td>
</tr>
<tr>
<td>geq12 &amp; geq20 LPV</td>
<td>OGLE-III (SMC+LMC)</td>
<td>~30%</td>
<td>&lt; few%</td>
<td></td>
</tr>
<tr>
<td>geq20 Short timescale</td>
<td>OGLE-II, III, IV (MCs)</td>
<td>0.05%</td>
<td>10–20%</td>
<td></td>
</tr>
<tr>
<td>geq20 Rotational modulation</td>
<td>The Pleiades (Hartman et al. 2010)</td>
<td>0.7 – 5.0%</td>
<td>&lt; 50%</td>
<td></td>
</tr>
</tbody>
</table>

Notes. (a) Estimated from non-cross-match random samples inspection discussed in Sect. 4.4, extending partially beyond OGLE-IV footprint. (b) No reliable contamination rate for the Bulge has been derived, but in the Galactic disk footprint (see Udalski 2017), contamination of the SOS Cepheids and RR Lyrae stars might be of the order of several tens of percent (A. Udalski, priv. comm. based on unpublished OGLE-IV data). (c) Estimates are restricted to sources with a peak-to-peak amplitude >0.2 mag and periods longer than 60 days. (d) The 50% ‘contamination’ quantifies the non-occurrence in the reference catalogue; the actual contamination of grazing binaries is estimated to be <1%, see Sect. 4.4.

include period search, the determination of bolometric correction from $G_{BP}$ and $G_{RP}$ fluxes, the identification of red supergiants, and the computation of absolute magnitude. It must be noted that all LPV specific attributes published in DR2 that depend on the parallax must be recomputed based on the published parallaxes, because these LPV-specific attributes were computed based on an older version of the parallaxes that was available at the time of variability processing within the consortium. We refer to Mowlavi et al. [2018] for additional information. The light curves of all LPV candidates, whether identified in the classification table or in the SOS table, are published in DR2.

The completeness and contamination numbers for the LPV candidates towards the Magellanic Clouds are computed relative to a sub-sample of the OGLE-III catalogue of LPVs [Soszyński et al. 2009, 2011]. The sub-sample consists of all OGLE-III LPVs with 7-band amplitude larger than 0.2 mag for the classification completeness and contamination estimates, and is further restricted to the OGLE-III LPVs with periods longer than 60 d for the SOS completeness and contamination estimates. Additionally, an all-sky contamination estimate is provided for Gaia LPV candidates that have small relative parallax uncertainties based on their position in the HR diagram. The main contaminants in this sample consist of young stellar objects. Details of SOS LPV processing and estimates of completeness and contamination can be found in Mowlavi et al. [2018], where the DR2 data set of LPV candidates is presented and further analysed.

4.4.3. Short-timescale variability

Short-timescale variable sources are astrophysical objects showing any variability in their optical light-curve with a characteristic timescale of the variation shorter than 1 d (Roelens et al. 2017). A variety of astronomical objects are known to exhibit such fast photometric variations, including both periodic and transient variability, involving amplitudes from a few millimag-nitudes to a few magnitudes, with different variability characteristics and phenomena at the origin of variation, from pulsations to flares to eclipsing systems. The diverse variability types targeted by the short-timescale module range from short-period binary stars to pulsating white dwarfs (e.g., ZZ Ceti stars or V777 Herculis stars) and cataclysmic variables such as AM CVn stars. However, for the Gaia DR 2, the short-timescale variability processing is oriented towards periodic variability, and the exact variable type of the detected short-timescale candidates is not yet determined. Both the investigation of transient variability (Wevers et al. 2018) and the further classification of the candidates are foreseen for the future Gaia data releases.

As mentioned previously, the identification of the short-timescale variable candidates published in Gaia DR2 involves particular variability detection methods that are specifically tuned to detect fast photometric variability, complementarily to the general variability detection approach (which may miss some of these events because they are so fast). These methods are applied to a preselected subset of the Gaia data set (Eyer et al. [2018]) and the short-timescale variability is expected to be recovered for the full Gaia DR2 data set.
for Gaia DR 2, the short-timescale variability analysis is restricted to relatively faint sources ($G \sim 16.5 – 20$ mag) whose $G$ per-CCD time series indicates that they are likely to show rapid variations (Roelens et al. 2018). Consequently, the candidates resulting from the short-timescale module are expected to overlap with other short-period SOS variables such as RR Lyrae and δ Scuti stars with periods shorter than 1 d, even though not there is no complete overlap with these SOS variable candidates.

The SOS short-timescale module identified 3018 short-timescale, suspected periodic candidates. These bona fide, short-timescale, suspected periodic candidates were identified based on the combination of the variogram analysis (Eyer & Genoni 1999; Roelens et al. 2017), least-squares frequency search (Zechmeister & Kürster 2009), analysis of the environment of candidate sources over the sky, and a series of selection criteria from various statistics, such as the Abbe or inter-quartile range (IQR) values in the three bands ($G, G_{BP}$ and $G_{RP}$).

Completeness and contamination numbers for the short-timescale variable candidate sample, estimated by comparing OGLE-II, III, IV, and Gaia data in the Magellanic Clouds, are listed in Table 3. Furthermore, about 50% of the known OGLE-III and IV variables belonging to the short-timescale sample in these regions are longer period variables (i.e. with period > 1d). However, most of these contaminant sources have periods of a few days and quite high amplitudes, hence although they are not true short-period variables, their presence in the short-timescale published sample is justified. A more detailed description and analysis of the results can be found in Roelens et al. (2018).

4.4.4. Rotational modulation

Stellar flux modulation induced by surface inhomogeneities and rotation is searched for in a region of the HR diagram, broadly embracing stars of spectral type later than F on the main sequence. This type of variability, often indicated in the literature as BY Draconis stars (BY Dra) regardless of possible binarity, is an indication of active regions (dark spots and bright faculae on the stellar photosphere) that are produced by stellar magnetic activity. The evolution of stellar magnetic fields, similar to solar fields, produces variability phenomena on a wide range of timescales (see e.g. Lanza et al. 2004; Distefano et al. 2012; Lanza et al. 2006 and references therein), one of the most prominent phenomena is the rotational modulation itself, from which the stellar rotation period can be inferred.

The variability pipeline comprises two packages dedicated to the detection and characterisation of solar-like stars with rotational modulation. First, solar-like variable candidates are selected, and if they are confirmed, they are studied in more detail to determine stellar rotation periods and other properties. Fig. 8 shows one of the published light-curves together with the segments in which a significant (similar) period was detected.

We detected some $7 \times 10^5$ periodic variables in the pre-selected HR diagram region, which excludes areas populated by pulsating stars. The remaining expected main contaminants are eclipsing binaries and spurious detections derived from an incomplete sampling. To filter out these cases as well as possible, we applied filters that exclude sources whose period-folded light curves have a significantly uneven distribution in phase, with significant gaps, and are far from being sinusoidal (see Lanzafame et al. 2018 for details). The final DR2 clean sample contains 147 535 rotational modulation candidates and fills 38% of the whole sky when divided into bins of $\approx 0.84$ deg$^2$ (level 6 HEALPix), see for example Fig. 8.

An estimate of the final completeness and contamination of the DR2 rotational modulation variables is hampered by the fact that the occurrence rate of the BY Dra phenomenon is largely unknown so far. It is expected that all low-mass dwarfs are magnetically active to some extent, and the rotational modulation detectability essentially depends on the instrumental sensitivity, and also on the active region distribution and the phase of the magnetic cycle at the epoch of observations. New low-mass dwarfs displaying rotational modulation are continuously detected with the ever-wider span and increasing sensitivity of modern surveys, none of which has the full sky coverage capabilities of Gaia, however. At this stage, it is possible to perform some meaningful comparison only with the Hartman et al. (2010) and Rebull et al. (2016) observations of the Pleiades, with which, nevertheless, the DR2 geq20 sky coverage still overlaps only marginally. Assuming that the Hartman et al. (2010) catalogue lists all the BY Dra in its FoV down to $G = 14.5$, we estimate that the completeness of the BY Dra sample is 14% down to $G = 16.5$ in the overlapping field. Assuming that this value is uniform over the whole sky, we estimate a completeness upper limit of 5%. At the other extreme, we may assume that all low-mass dwarfs are BY Dra variables. Then, comparing with all stars observed by Gaia in the same sky region and magnitude range, we estimate a lower completeness limit of 0.7%.

We may roughly estimate an upper limit for contamination by assuming that the Gaia BY Dra detections that are not present in the Hartman et al. (2010) or in the Rebull et al. (2016) sample in the common region on the sky are variables of other types. This upper limit is obviously largely overestimated given that no BY Dra catalogue can be deemed complete to date, as also testified by the fact that there are sources in the Hartman et al. (2010) sample that are not contained in the Rebull et al. (2016) sample, and vice versa.

At the other extreme, we can make an educated guess regarding the contamination of close grazing binaries that would be incorrectly classified as BY Dra because of the similarity of the light curves. If the orbital period is long, then the grazing eclipse occurs only in a small part of the orbit, and it does not look like rotation modulation. We assume an upper limit period of 10 days for this contaminating effect. The fraction of stars with such short-period companions is only about 2–5% (Raghavan et al. 2010), of which only 2–5% are grazing binaries (at most), which means 0.04–0.25 % of the stars, while we did not yet take into account that for most binaries, the secondaries are substantially smaller. Even though we cannot consider the rotation modulation sample a random subset of G and later-type stars (which is the assumption in above estimate), it seems unlikely that this type of contamination would exceed 1%. The same conclusion is reached from our internal validation against known and newly identified (but not published) eclipsing binaries, based on which, we estimate an upper limit of 0.5% grazing contamination. From the estimated grazing binary contamination and non-occurrence in the reference catalogue discussed above, we estimate a contamination level ranging from $< 1\%$ to $\approx 50\%$, as listed in Table 3. More details on the SVD solar-like and SOS rotational modulation packages and results can be found in Lanzafame et al. (2018).

4.5. Sources with multiple SOS types

There are 80 sources in DR2 with an entry in more than one SOS table. These cases are all overlaps between short-timescale objects and RR Lyrae stars (72), Cepheids (5), and rotation mod-
ulation objects (3), all of which are justified by their overlapping type definition. No overlap is found with long-period variables.

4.6. Sources with different classification and SOS type

Differences are found between classification types (as obtained by supervised classification in the geq2 path) and the types that are eventually derived in SOS results (either from the geq12 or geq20 paths). In particular, one source was classified as DSCT_SXPHE, but also appears in SOS rotational modulation. Similarly, one source was classified as MIRA_SR, was not confirmed by SOS LPVs, but appears in SOS rotational modulation. Of the sources classified as a Cepheid (any type), 6 appear in more than one SOS table. One is a confirmed SOS Cep and also appears in SOS short timescale, while 3 appear as SOS rotational modulation and 2 as SOS short timescale. Of the sources classified as an RR Lyrae star (any type), 70 appear in the SOS short timescale table, one was instead confirmed as a Cepheid by SOS (shown in the top panel of Fig.2), and many were confirmed as RR Lyrae star by SOS. Overall, a reassignment of the type from RRL into Cep by SOS occurred in 618 cases, but only in 77 cases from Cepheid to RR Lyrae star by SOS. We have decided to retain the original classification type despite the possible reassignment by the SOS Cep&RRL module or in cases where they were confirmed as another type by SOS modules.

5. Conclusions

In this second Gaia DR2 data release, a sample of 550 737 variable sources with their three-band time series and analysis results has been released to the astronomical community, showcasing samples of high-amplitude variable candidates distributed over (a large portion of) the whole sky. It demonstrates the immense potential of Gaia data to provide an unbiased all-sky photometric survey together with astrometric and spectroscopic data. Although the variability analyses in DR2 did only mildly rely on the astrometric data and on none of the spectroscopic data, this dependence will become much heavier in future data releases.

Several caveats have been outlined in this paper:

- Most provided stellar samples are rather incomplete as a result of the current state of data calibration (e.g. even well-known variables might be missing) or the limited scanning law coverage for the minimum number of selected observations,
- known incomplete flagging of outliers in the time series,
- existence of uncalibrated systematics or spurious calibration error signals in some of the time series,
- the probabilistic and automated nature of this work implies certain completeness and contamination rates (Tables 2 and 3), hence even well-known literature sources could be misidentified.

Overall, we estimate to have identified the following numbers of new variables: a few hundred Cepheids, a few tens of thousand RR Lyrae stars, about one hundred thousand stars with rotation modulation, several tens of thousand long-period variables, a few thousand δ Scuti and SX Phoenicis stars, and a few thousand short-timescale variables. In total, this is about half of the variable sources released in this DR2. Compared to the DR1 example, we increased the sample size about 100 times, which is expected to result in a large number of novel studies, regardless of whether they are combined with existing data sets.

The next variability release will be based on data that are better calibrated yet again and on a longer time baseline. It will also include BP/RP spectrophotometry and radial velocities from RVS, which will allow us to improve the quality and quantity of the Gaia DR3 release.

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We dedicate this paper to our dear friend and colleague Jan Cuypers (1956-2017), whose passing has left a great void in our variability group. Thank you, Jan, for your friendship and lifelong contribution to science. We would like to thank the referee, Andrej Udalski, for various suggestions and providing so far unpublishedOGLE-IV estimates that improved the contents of this paper. We thank Xavier Luri, Antonella Valenari, and Annie Robin for valuable feedback during the preparation of this paper. We also acknowledge Mark Taylor for creating and implementing new requested features in the astronomy-oriented data handling and visualization software TOPCAT (Taylor 2005). This work has made use of data from the ESA space mission Gaia, processed by the Gaia Data Processing and Analysis Consortium (DPAC).

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Lanza, A. F., Messina, S., Pagano, I., & Rodonò, M. 2006, Astronomische Nachrichten, 327, 21

8 From occurrence in a great variety of cross-match catalogues (most of which are documented in the DR2 documentation), and taking into account that our results have a varying degree of contamination.

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Table 4: Statistical parameters for the nTransits: 2+ classification types published in DR2. The 1%, 50% (median), and 99% percentile are provided for each type.

<table>
<thead>
<tr>
<th>G band</th>
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<th>Max. magnitude</th>
<th>Mean magnitude</th>
<th>Median magnitude</th>
<th>Magnitude range</th>
<th>Standard deviation (mag)</th>
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**Notes:**
- Cepheid groups: best_class_name = CEP, ACEP, T2CEP
- RR Lyrae star groups: best_class_name = RRAB, RRC, RRD, ARRD
- Long-period var. is best_class_name = MIRA_SR
- δ Sct and SX Phe is: best_class_name = DSCT_SXPHE
- For the skewness and kurtosis, we included only stars with > 2 and > 3 FoV transits, respectively.
Table 5: Statistical parameters for the SOS types published in DR2. The 1%, 50% (median), and 99% percentile are provided for each type.

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For RR Lyrae stars, the percentiles for $G_{BP}$ and $G_{RP}$ were computed for a number of FOV transits $> 1$. Although there are stars with 0 or 1 transits in the band, the following are the results.

Notes: For RR Lyrae stars, the percentiles for $G_{BP}$ and $G_{RP}$ were computed for a number of FOV transits $> 1$. Although there are stars with 0 or 1 transits in the band, the following are the results.

For the skewness and kurtosis, we included only RR Lyrae stars with $G_{BP} < 0.1$.
Fig. 1: Relative source counts with respect to sources having ≥2 $G$-band FoV transits, illustrating the drop in relative completeness for source-sets with ≥12 and ≥20 $G$-band FoV transits, as applied in our processing (see Fig. 2). Regional values are listed in Table 1 and discussed in Sect. 2.1. Values for external catalogs are Gaia source counts in the approximate sky footprint, not source cross-match counts. Note that actual time series are published only for the 550K variable sources in DR2 (see Fig. 3).
Fig. 2: DR2 variability processing overview. Data were published from the highlighted yellow boxes for the validation filtered sources. There were two main tracks: one starting from ≥ 2 G-band FoV transits (left), and the other from ≥ 20 G-band FoV transits (right). The former resulted in the published nTransits:2+ classifier in the vari_classifier_result table, and the latter resulted in the published SOS tables of vari_short_timescale and vari_rotation_modulation. The published SOS tables of vari_rrlyrae, vari_cepheid and vari_long_period_variable result from a mixed feed of the published nTransits:2+ classifier (for sources with at least ≥ 12 G-band FoV transits) and from the unpublished classifier of the ≥ 20 track.

Fig. 3: Accounting of the 550 737 unique variable sources published in DR2. All have an entry in the Gaia archive gaia_source and vari_time_series_statistics tables. Their distributions in the additional variability tables is shown. The classification output types are mutually exclusive so that their sum matches the total number of entries in the published vari_classifier_result table. Eighty sources appear in more than one SOS table (not illustrated), which is detailed in Sect. 4.5. The overlap between similar-type classification and SOS is detailed in a Venn-like fashion showing the counts in the three possible subsets. A few hundred sources have different assigned types in classification and SOS tables (not illustrated), as explained in Sect. 4.6, hence the difference on the left between the total type count and the unique source_id count. No SOS module treated δ Scuti & SX Phoenicis stars.
Fig. 4: Sky source densities [count deg$^{-2}$] in Galactic coordinates of the published sources in the classification table (left column, see Sect. 4.3) and SOS tables (right column, see Sect. 4.4). In the classification plots the best_class_name entries were grouped by main type, as listed in parentheses. Galactic longitude increases to the left side.

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Fig. 5: Histogram counts of the median G-band magnitude distribution of the published classification classes (left, see Sect. 4.3) and SOS tables (right, see Sect. 4.4). In the classification plots the best_class_name entries were grouped by main type, as listed in parentheses. The filled bars correspond to the linear scale on the left side of the plots (which varies in scale), while the dotted line corresponds to the logarithmic scale on the right side of the plots (always in the same scale). The bin size is fixed to 0.5 mag.
Fig. 6: Example light-curves of RR Lyrae stars and LPV in the data. For each band we show the valid FoV transits, as well as variability-rejected FoV transits (number in parentheses, data plotted with crosses). In most cases, two FoV-transits occur within the 6 h spin period, causing indistinguishable overlapping points in the left diagrams. Some outliers are not flagged in the exported data (see e.g. the G$_{RP}$ light curve of the RRc and RRd example); these transits were removed in the more strict outlier rejection used in the SOS Cep and RRL module, as mentioned in Sect. 4.1.
therefore the period for the DSCT_SXPHE rejection used in the SOS Cep and RRL module, as mentioned in Sect. 4.1. No period is published for the classification results.

Fig. 7: Example light-curves of the Cepheid, δ Scuti, and short-timescale variables. For each band we show the valid FoV transits, as well as variability-rejected FoV transits (number in parentheses, data plotted with crosses). In most cases, two FoV-transits occur within the 6 h spin period, causing indistinguishable overlapping points in the left diagrams. Some outliers are not flagged in the exported data (see e.g. the $G_{BP}$ and $G_{RP}$ light curve of the ACEP example); these transits were removed in the more strict outlier rejection used in the SOS Cep and RRL module, as mentioned in Sect. 4.1. No period is published for the classification results, therefore the period for the DSCT_SXPHE example is taken from the LINEAR survey (Palaversa et al. 2013).
Fig. 8: Example light-curves of a solar-like star with rotation modulation. Period search is performed over multiple segments, as shown in the right panels; see Sect. 4.4.4 for more details. The bottom left panel expands a high-cadence time range of the time series. The top left panel shows for each band the valid FoV transits and variability-rejected FoV transits (number in parentheses, here 0).
Fig. 9: Selected (i.e. not-rejected) number of transits in the time series of the $G$, $G_{\text{BP}}$, and $G_{\text{RP}}$ photometric bands in Galactic coordinates for all 550,737 published sources. These numbers have been aggregated in sky bins of about 0.84 deg$^2$ (level 6 HEALPix) and their mean value is shown. The colour scale is clipped at 100, although the number of transits reaches up to 242, as shown in Fig. 10. The choice of the rainbow colour scale has been motivated by the convenient colour divisions between 5, 10, 20, 30, 40, 50, and 100 transits. Galactic longitude increases to the left side.

Fig. 10: Histogram of selected (i.e. not-rejected) number of transits in time series of the $G$, $G_{\text{BP}}$, and $G_{\text{RP}}$ photometric bands for the 550,737 published variable sources. The filled bars correspond to the linear scale on the left side, while the solid lines corresponds to the logarithmic scale on the right side of the plot. The bin size is 1 FoV transit.
Appendix A: Examples of Gaia archive queries

This section describes Gaia archive queries in the ADQL format that return the various selections presented in this paper. These queries can be made online in the Gaia archive at [http://gea.esac.esa.int/archive/](http://gea.esac.esa.int/archive/)

Appendix A.1: Total counts and table merging

Figure [total number of unique sources with variability results (bottom left unique source_id count):

```
select count(*) from gaiadr2.gaia_source where phot_variable_flag='VARIABLE'
```

Retrieve various astrometric and general statistical variability fields for all published variable sources in Gaia DR2:

```
select gs.source_id, ra, dec, parallax, parallax_error, pmra, pmdec, epoch_photometry_url,
num_selected_g_fov, mean_mag_g_fov, range_mag_g_fov, std_dev_mag_g_fov, time_duration_g_fov,
num_selected_bp, mean_mag_bp, range_mag_bp, std_dev_mag_bp, time_duration_bp,
num_selected_rp, mean_mag_rp, range_mag_rp, std_dev_mag_rp, time_duration_rp
from gaiadr2.gaia_source as gs join gaiadr2.vari_time_series_statistics as varistats using(source_id)
where phot_variable_flag='VARIABLE'
```

Appendix A.2: Creating a variable summary table

It is very desirable to have an overview of the occurrence of a given source_id in any of the vari_* tables. The following query generates such a table in which each row contains a unique variable source source_id, and each column represents one of the vari-tables. A column contains integer 1 when it contains the source_id and NULL when not. After logging into the Gaia archive a user can be created from the query result (suggested name: gaiadr2_vari_summary), after which it appears in the user tables and can be queried like any other DR2 table.

```
select
  gs.source_id,
  vrrl.incl as vari_rrlyrae -- 1 if RR Lyrae SOS table entry
  vcep.incl as vari_cepheid -- 1 if Cepheid SOS table entry
  vrm.incl as vari_rotation_modulation -- 1 if rotation modulation SOS table entry
  vsts.incl as vari_short_timescale -- 1 if short timescale SOS table entry
  vlpv.incl as vari_long_period_variable -- 1 if long period variable SOS table entry
  vcr_rrl.incl as vari_classifier_result_rrl -- 1 if classif. entry of any RR Lyrae type
  vcr_cpe.incl as vari_classifier_result_cpe -- 1 if classif. entry of any Cepheid type
  vcr_dscsxp.incl as vari_classifier_result_dsct_sxphe -- 1 if classif. entry of Delta Scuti
  vcr_mirasr.incl as vari_classifier_result_mira -- 1 if classif. entry of Mira or semi-regular type
  1 as vari_time_series_statistics -- 1 if time series statistics entry (1 for all in table)
  1 as gaia_source_variable_flag -- 1 if gaia_source phot_variable_flag='VARIABLE' (1 for all in table)
from gaiadr2.gaia_source as gs
LEFT OUTER join (select source_id, 1 as incl from gaiadr2.vari_rrlyrae) as vrrl using(source_id)
LEFT OUTER join (select source_id, 1 as incl from gaiadr2.vari_cepheid) as vcep using(source_id)
LEFT OUTER join (select source_id, 1 as incl from gaiadr2.vari_rotation_modulation) as vrm using(source_id)
LEFT OUTER join (select source_id, 1 as incl from gaiadr2.vari_short_timescale) as vsts using(source_id)
LEFT OUTER join (select source_id, 1 as incl from gaiadr2.vari_long_period_variable) as vlpv using(source_id)
LEFT OUTER join (select source_id, 1 as incl from gaiadr2.vari_classifier_result where best_class_name='RRAB' or best_class_name='RRC' or best_class_name='RRD' or best_class_name='ARRD') as vcr_rrl using(source_id)
LEFT OUTER join (select source_id, 1 as incl from gaiadr2.vari_classifier_result where best_class_name='CEP' or best_class_name='ACEP' or best_class_name='T2CEP') as vcr_cpe using(source_id)
order by gs.source_id
```

Appendix A.3: HEALPix grouped data

To reproduce sky-distribution data sets using HEALPix grouping as shown in Figures [4] and [10] the convenient gaia_healpix_index(norder, source_id) function in the archive that extracts the HEALPix pixel id that is encoded in the Gaia source_id can be used. Using norder=6 creates pixels of about 0.84 deg^2. A program like TOPCAT [Taylor 2005] can read and visualise such data set with ease.
Figure 4 top left panel ‘Classif.: RR Lyrae stars (RRAB, RRC, RRD, ARRD)’ number of sources per square degree, and also
the min, average, max, and standard deviation of the vari_classifier_result.best_class_score value per HEALPix pixel:

\[
\text{select} \\
\text{gaia Healpix index(6, source id) AS healpix_6,} \\
count(*) / 0.83929 as sources per sq deg, \\
\min\text{(best class score)} \text{AS min best class score,} \\
\text{avg\text{(best class score)} AS avg best class score,} \\
\max\text{(best class score)} \text{AS max best class score,} \\
\text{stdev\text{(best class score)} AS stddc_int average g} \\
\text{FROM gaiadr2.vari_classifier_result WHERE} \\
\text{best class name='RRAB' or best class name='RRC' or best class name='RRD' or best class name='ARRD'} \\
\text{GROUP BY healpix_6}
\]

Figure 4 top right panel ‘SOS: RR Lyrae stars’: number of sources per square degree, and also the min, average, max, and
standard deviation of the vari_rrlyrae.int_average_g value per HEALPix pixel:

\[
\text{select} \\
\text{gaia Healpix index(6, source id) AS healpix_6,} \\
count(*) / 0.83929 as sources per sq deg, \\
\min\text{(int average g)} \text{AS min int average g,} \\
\text{avg\text{(int average g)} AS avg int average g,} \\
\max\text{(int average g)} \text{AS max int average g,} \\
\text{stddc\text{(int average g)} AS stddc_int average g} \\
\text{FROM gaiadr2.vari_rrlyrae} \\
\text{GROUP BY healpix_6}
\]

Figure 10 all three panels (‘number of transits in time series of the G, G_BP, and G_RP photometric bands’) average, and also the
min, max, and standard deviation of the values per HEALPix pixel:

\[
\text{select} \\
\text{gaia Healpix index(6, source id) AS healpix_6,} \\
count(*) / 0.83929 as sources per sq deg, \\
-- G-band FoV \text{avg\text{(num selected g fov)}} \text{AS avg num selected g fov,} \\
\text{min\text{(num selected g fov)}} \text{AS min num selected g fov,} \\
\text{max\text{(num selected g fov)}} \text{AS max num selected g fov,} \\
\text{stddc\text{(num selected g fov)} AS stddc num selected g fov,} \\
-- BP-band \text{avg\text{(num selected bp)}} \text{AS avg num selected bp,} \\
\text{min\text{(num selected bp)}} \text{AS min num selected bp,} \\
\text{max\text{(num selected bp)}} \text{AS max num selected bp,} \\
\text{stddc\text{(num selected bp)} AS stddc num selected bp,} \\
-- RP-band \text{avg\text{(num selected rp)}} \text{AS avg num selected rp,} \\
\text{min\text{(num selected rp)}} \text{AS min num selected rp,} \\
\text{max\text{(num selected rp)}} \text{AS max num selected rp,} \\
\text{stddc\text{(num selected rp)} AS stddc num selected rp} \\
\text{FROM gaiadr2.vari_time_series_statistics} \\
\text{GROUP BY healpix_6}
\]

Appendix A.4: Detailed counting

This section shows how the counts presented in Figures 3 and 5 can be reproduced from the published DR2 tables. Alternatively, a
summary table can first be created as described in appendix A.2 and the relevant selections on the columns involved (not shown)
can be made.

Figures 3 and 5 total number of sources with classification: all / δ Scuti and SX Phoenicis / RR Lyrae stars / Cepheids / LPV (Mira and semi-regular):

\[
\text{select count(*) from gaiadr2.vari_classifier_result} \\
\text{select count(*) from gaiadr2.vari_classifier_result where best class name='DSCT_SXPHE'} \\
\text{select count(*) from gaiadr2.vari_classifier_result where best class name='RRAB'} \\
\text{or best class name='RRC' or best class name='RRD' or best class name='ARRD'} \\
\text{select count(*) from gaiadr2.vari_classifier_result where best class name='CEP'} \\
\text{or best class name='ACEP' or best class name='T2CEP'} \\
\text{select count(*) from gaiadr2.vari_classifier_result where best class name='MIRA_SR'}
\]

Figures 3 and 5 total number of sources with SOS: RR Lyrae / Cepheid / LPV / short timescale / rotation modulation results:
select count(*) from gaiadr2.vari_rrlyrae
select count(*) from gaiadr2.vari_cepheid
select count(*) from gaiadr2.vari_long_period_variable
select count(*) from gaiadr2.vari_short_timescale
select count(*) from gaiadr2.vari_rotation_modulation

Figure 3 number of sources in common between classification and SOS for: RR Lyrae / Cepheid / LPV in Gaia DR2, these are the numbers in the overlap blocks of the Venn-diagrams:

select count(*) from gaiadr2.vari_classifier_result as clasrrl
INNER join gaiadr2.vari_rrlyrae as sosrrl on clasrrl.source_id=sosrrl.source_id
where best_class_name='RRAB' or best_class_name='RRC' or best_class_name='RRD'
  or best_class_name='ARRD'
select count(*) from gaiadr2.vari_classifier_result as clascep
INNER join gaiadr2.vari_cepheid as soscep on clascep.source_id=soscep.source_id
where best_class_name='CEP' or best_class_name='ACEP' or best_class_name='T2CEP'
select count(*) from gaiadr2.vari_classifier_result as claslpv
INNER join gaiadr2.vari_long_period_variable as soslpv
on claslpv.source_id=soslpv.source_id
where best_class_name='MIRA_SR'

Figure 4 number of sources that are only found in classification and not in SOS for RR Lyrae / Cepheid / LPV in Gaia DR2, these are the numbers in the left blocks of the Venn-diagrams:

select count(*) from gaiadr2.vari_classifier_result as clasrrl
LEFT OUTER join gaiadr2.vari_rrlyrae as sosrrl on clasrrl.source_id=sosrrl.source_id
where (best_class_name='RRAB' or best_class_name='RRC' or best_class_name='RRD'
  or best_class_name='ARRD') and sosrrl.source_id is NULL
select count(*) from gaiadr2.vari_classifier_result as clascep
LEFT OUTER join gaiadr2.vari_cepheid as soscep on clascep.source_id=soscep.source_id
where (best_class_name='CEP' or best_class_name='ACEP' or best_class_name='T2CEP')
  and soscep.source_id is NULL
select count(*) from gaiadr2.vari_classifier_result as claslpv
LEFT OUTER join gaiadr2.vari_long_period_variable as soslpv
on claslpv.source_id=soslpv.source_id
where best_class_name='MIRA_SR' and soslpv.source_id is NULL

Figure 5 number of sources that are only found in SOS and not in classification for: RR Lyrae / Cepheid / LPV in Gaia DR2, these are the numbers in the right blocks of the Venn-diagrams:

select count(*) from gaiadr2.vari_classifier_result as clasrrl
RIGHT OUTER join gaiadr2.vari_rrlyrae as sosrrl on clasrrl.source_id=sosrrl.source_id
where (best_class_name='RRAB' and best_class_name='RRC' and best_class_name='RRD'
  and best_class_name='ARRD') or clasrrl.source_id is NULL
select count(*) from gaiadr2.vari_classifier_result as clascep
RIGHT OUTER join gaiadr2.vari_cepheid as soscep on clascep.source_id=soscep.source_id
where (best_class_name='CEP' and best_class_name='ACEP' and best_class_name='T2CEP')
  or clascep.source_id is NULL
select count(*) from gaiadr2.vari_classifier_result as claslpv
RIGHT OUTER join gaiadr2.vari_long_period_variable as soslpv
on claslpv.source_id=soslpv.source_id
where best_class_name='MIRA_SR' or claslpv.source_id is NULL