

# Low-amplitude rotational modulation rather than pulsations in the CoRoT B-type supergiant HD 46769<sup>★</sup>

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## ABSTRACT

**Aims.** We aim to detect and interpret photometric and spectroscopic variability of the bright CoRoT B-type supergiant target HD 46769 ( $V = 5.79$ ). We also attempt to detect a magnetic field in the target.

**Methods.** We analyse a 23-day oversampled CoRoT light curve after detrending and spectroscopic follow-up data using standard Fourier analysis and phase dispersion minimization methods. We determine the fundamental parameters of the star, as well as its abundances from the most prominent spectral lines. We perform a Monte Carlo analysis of spectropolarimetric data to obtain an upper limit of the polar magnetic field, assuming a dipole field.

**Results.** In the CoRoT data, we detect a dominant period of 4.84 d with an amplitude of 87 ppm and some of its (sub-)multiples. Given the shape of the phase-folded light curve and the absence of binary motion, we interpret the dominant variability in terms of rotational modulation, with a rotation period of 9.69 d. Subtraction of the rotational modulation signal does not reveal any sign of pulsations. Our results are consistent with the absence of variability in the Hipparcos light curve. The spectroscopy leads to a projected rotational velocity of  $72 \pm 2 \text{ km s}^{-1}$  and does not reveal periodic variability or the need to invoke macroturbulent line broadening. No signature of a magnetic field is detected in our data. A field stronger than  $\sim 500 \text{ G}$  at the poles can be excluded, unless the possible non-detected field were more complex than dipolar.

**Conclusions.** The absence of pulsations and macroturbulence of this evolved B-type supergiant is placed into the context of instability computations and of observed variability of evolved B-type stars.

**Key words.** Stars: individual: HD 46769 – Techniques: spectroscopic – Stars: fundamental parameters – Stars: rotation – Stars: supergiants

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## 1. Introduction

High-precision space asteroseismic data delivered by the MOST (Walker et al. 2003), CoRoT (Auvergne et al. 2009), and *Kepler* (Gilliland et al. 2010) missions have revolutionised our capacity to probe the interior physics and evolutionary status of various types of stars. For a few selected highlights concerning main-sequence stars with masses ranging from solar to  $25 M_{\odot}$ , as well as low-mass evolved stars in various nuclear burning stages, see, e.g. Michel et al. (2008), Huat et al. (2009), Degroote et al. (2010), Briquet et al. (2011), Bedding et al. (2011), Charpinet et al. (2011), Michel & Baglin (2012), Beck et al. (2012), Chaplin & Miglio (2013). Despite the immense and fast progress, the evolutionary models for some types of stars remain uncalibrated in terms of high-precision asteroseismic constraints. This

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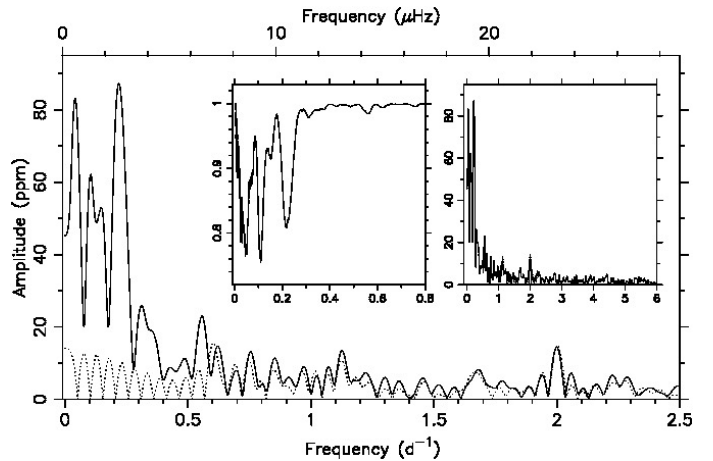
is particularly the case for evolved massive stars, while models of such objects have a large impact on various research areas of astrophysics, such as the computation of models with time-variable mass loss for core-collapse supernovae progenitors or the estimation of the chemical enrichment of galaxies.

Only three massive supergiants have been studied in detail from uninterrupted, time-resolved high-precision space photometry so far. The first case concerns the B2Ib/II star HD 163899, for which Saio et al. (2006) reported the discovery of tens of pressure and gravity modes with frequencies below  $2.8 \text{ d}^{-1}$  ( $32.4 \mu\text{Hz}$ ), and amplitudes below a few mmag from a 37 d MOST light curve. Although hot supergiant models for masses between 20 and  $24 M_{\odot}$  and their pulsation predictions showed qualitative agreement with the observations, the interior physics of those models could not be improved due to the lack of high-precision spectroscopy. Such spectroscopy is necessary for good fundamental parameter and abundance estimation as well as for mode identification of the detected pulsation frequencies. The second supergiant monitored from space is the B6Ia star HD 50064, which was observed by CoRoT for 137 d and followed up with high-resolution spectroscopy covering a time base of 169 d (Aerts et al. 2010b). This supergiant was found to have a variable-amplitude monoperoiodic radial strange-mode pulsation with a period of 37 d, which is connected with the variable mass loss of the star. While its monoperoiodic character did not allow asteroseismic tuning of its interior structure, the interpretation of the observations was recently confirmed by evolutionary model computations for He-core burning models re-entering a blue loop in the Hertzsprung-Russell (HR) diagram after having passed the red-supergiant phase (Saio et al. 2013). In the third case of HD 34085 (Rigel, B8Ia), the MOST light curve of 28 d was too short to uncover the oscillations found in long-term spectroscopy (Moravveji et al. 2012).

Clearly, there is a need to study supergiant pulsations in much more detail to fine-tune evolutionary models in the upper HR diagram. The *Kepler* mission cannot help in this respect as there are no hot massive supergiants in its fixed field of view. It is in this context that we studied the bright massive B8Ib supergiant HD 46769 as secondary asteroseismology target with CoRoT. The fundamental parameters of this star were previously derived by Lefever et al. (2007), based on high-precision spectroscopy containing H and He lines only. They obtained  $T_{\text{eff}} = 12\,000 \text{ K}$ ,  $\log g = 2.55$ ,  $v \sin i = 68 \pm 5 \text{ km s}^{-1}$ , and  $\log(L/L_{\odot}) = 4.22$ . These authors also examined the Hipparcos data and did not find any significant variability. Zorec et al. (2009), on the other hand, applied the Barbier-Chalonge-Divan spectrophotometric classification method and found the effective temperature to be some 2 000 K higher. They assigned a spectral type of B7III/II to the star without providing an estimate for  $\log g$ . In any case, irrespective of the precise effective temperature and luminosity class, the star is situated outside any known instability strip. This was also the case for HD 163899, which was originally measured by MOST as a standard star. Thus, in view of the lack of supergiant asteroseismology so far and because it was the only hot massive supergiant in the pointed fields of CoRoT, HD 46769 was selected as secondary target for the mission.

## 2. CoRoT light curve analysis

The supergiant HD 46769 was observed by CoRoT during a short run in the anticentre direction of the Milky Way (SRa03) in March 2010. The asteroseismology runs deliver data with a cadence of 32 s, and SRa03 lasted for 26.4 d, from HJD 2455258.426769 until HJD 2455284.802757. A strong dis-

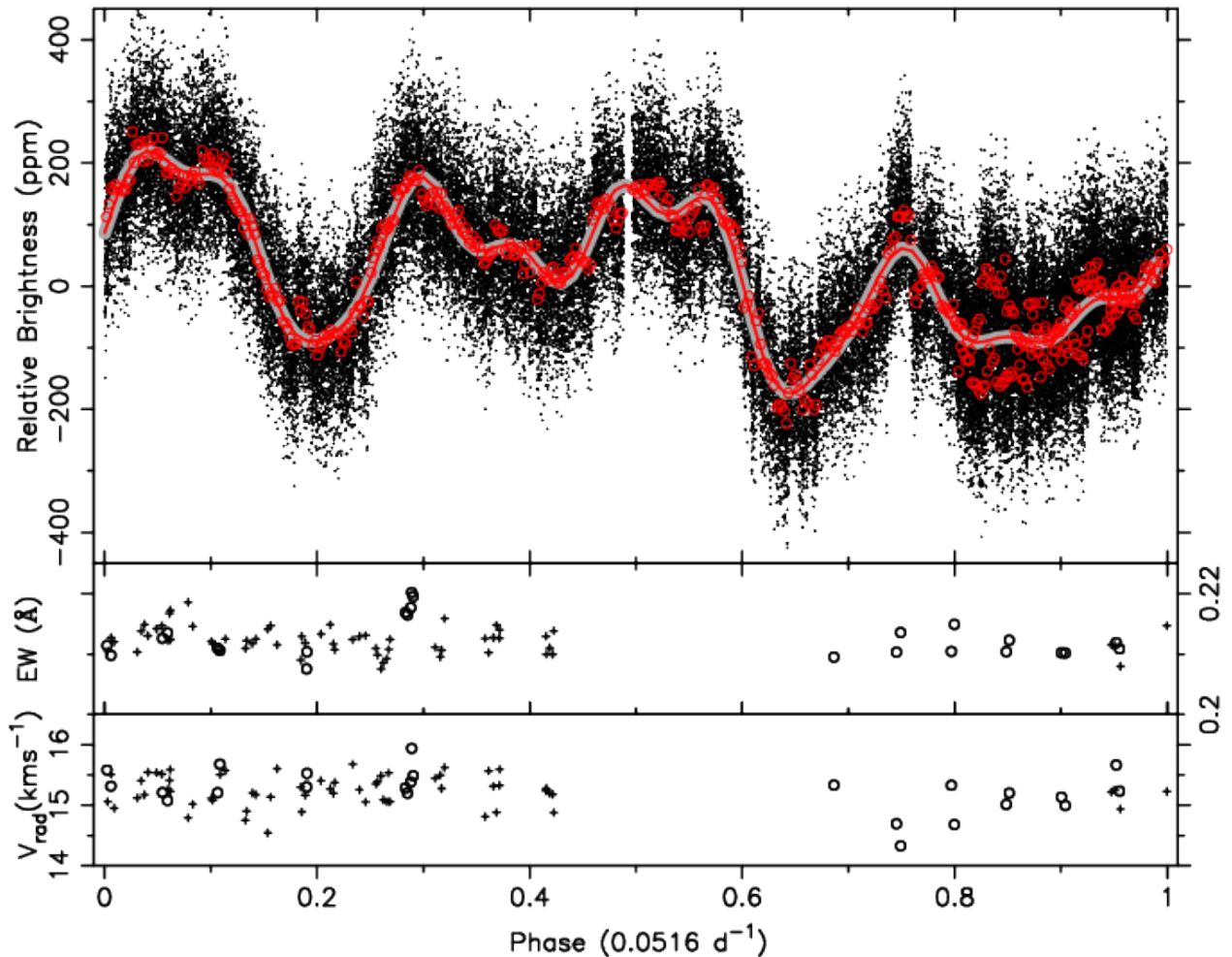


**Fig. 1.** Outcome of the frequency analysis for the CoRoT light curve of HD 46769. The main panel shows the Scargle periodogram over  $[0, 2.5] \text{ d}^{-1}$  (full line) and the version after prewhitening with  $f = 0.0516 \text{ d}^{-1}$  and seven of its harmonics as listed in Table 1 (dotted line). The left inset shows the  $\Theta$  statistic and the right inset shows the Scargle periodogram over  $[0, 6] \text{ d}^{-1}$ . No significant power occurs at frequencies above  $0.6 \text{ d}^{-1}$ , unless they are of instrumental origin (such as the peak near  $2 \text{ d}^{-1}$ ).

continuity connected with instrumental effects (Auvergne et al. 2009) occurs in the light curve near day HJD 2455262, after which both the mean flux level and the nature of the trend in the flux changed. As such, we restricted our analysis to the last 22.7 d of the light curve, which contains 53927 data points of good quality and has a Rayleigh limit of  $1/\Delta T = 0.044 \text{ d}^{-1}$ . The raw light curve shows an instrumental downward trend as present in most of the CoRoT light curves and interpreted as CCD ageing (Auvergne et al. 2009). Given that intrinsic variability occurs with a time scale similar to the duration of the data, which we do not want to affect, we only removed a linear trend before starting the frequency analysis. For the latter, we adopted the usual zero-point for the time stamps applied by the CoRoT consortium, i.e. 1 January 2000 at 12h UT, which corresponds with  $\text{HJD}_0 = 2451545.0$ .

We computed both the Scargle periodogram (Scargle 1982) and the  $\Theta$  statistic in the framework of phase dispersion minimization (PDM, Stellingwerf 1978). Maximum power in the Scargle periodogram occurs at the frequency near  $0.21 \text{ d}^{-1}$ , while  $\Theta$  reaches a minimum near  $0.11 \text{ d}^{-1}$  (Fig. 1). Both the Scargle and PDM periodograms also have a peak near  $0.05 \text{ d}^{-1}$ , corresponding with a periodicity of 19.4 d. The Scargle frequency at maximum amplitude is equal to four times the one of the second highest amplitude, within the frequency resolution.

In the next section, we add the spectroscopic time series to the photometry to end up with the most likely frequency  $4f$ , with value  $f = 0.0516 \text{ d}^{-1}$ . We thus made a sinusoidal fit with  $4f$  to the light curve. Subsequent prewhitening leads, within the frequency resolution, to  $f$ ,  $2f$ ,  $5f$ ,  $3f$ , and  $11f$  respectively, reflecting the strongly non-sinusoidal shape of the light curve (Fig. 2). The frequencies in the data with a significance above  $4\sigma$  are listed in Table 1. After prewhitening with those frequencies, we find a residual frequency at  $2 \text{ d}^{-1}$ . This frequency, which is well visible in Fig. 1, is a known artefact related to day-versus-night variations connected with CoRoT's low-Earth orbit (Auvergne et al. 2009). The satellite orbital frequency itself is  $13.97 \text{ d}^{-1}$ , and



**Fig. 2.** Upper: the linearly detrended CoRoT light curve of HD 46769 phased with respect to the frequency  $f = 0.0516 \text{ d}^{-1}$ . The grey thick full line is a harmonic fit with  $f$  and seven of its harmonics, of which  $4f = 0.2064 \text{ d}^{-1}$  is the dominant frequency in the light curve. The red symbols and red line denote averages over 101 consecutive data points and a fit to that binned light curve respectively. Middle and bottom: the equivalent width and radial velocity of the  $\text{Mg II } 4481 \text{ \AA}$  line as observed in the HERMES (circles) and HARPS (plus signs) spectra.

**Table 1.** The frequencies of HD 46769 in the CoRoT light curve along with their amplitudes and significance level.

Frequency ( $\text{d}^{-1}$ )	Amplitude (ppm)	Significance ( $\times$ noise level)
$f = 0.0516$	$67.3 \pm 0.6$	14.0
$2f$	$63.2 \pm 0.6$	13.2
$3f$	$42.9 \pm 0.6$	8.9
$4f$	$86.9 \pm 0.7$	18.1
$5f$	$44.9 \pm 0.7$	9.4
$6f$	$23.2 \pm 0.6$	4.8
$7f$	$21.4 \pm 0.8$	4.5
$11f$	$25.1 \pm 0.7$	5.2

**Notes.** The significance level is expressed in terms of the average noise level in the Scargle periodogram, which was computed as the average amplitude in the range  $[0, 2] \text{ d}^{-1}$  after prewhitening of all the frequencies listed here and amounts to 4.8 ppm.

its harmonics also occur in CoRoT light curves. However, this regime in frequency is not relevant for our study of HD 46769.

Given that the data is oversampled compared with the dominant periodicities, we repeated the frequency analysis after binning the light curve over 101, respectively 51, data points. For both cases, these binned light curves gave the same frequency results as those listed in Table 1 within the errors. The 101-binned phased light curve and its fit are shown in red in Fig. 2.

We subsequently re-analysed the Hipparcos data (Perryman 1997) of the star and did not find any significant variability. This is in agreement with the results by Lefever et al. (2007), who used HD 46769 as a constant comparison star in their study of 28 periodically variable B supergiants, whose variability they interpreted as being caused by gravity-mode pulsations. The highest peak in the Hipparcos Scargle periodogram for HD 46769 occurs near  $9 \text{ d}^{-1}$ ; however it reaches an amplitude of 4.6 mmag and is not significant in that data set. We explicitly checked that it does not occur in the CoRoT light curve. Given that the Hipparcos data can only reveal variability at a level of mmag (i.e. ppt) or higher, it is not surprising that we cannot detect the dominant frequency found in the CoRoT light curve in that data set.

We conclude that the CoRoT light curve of HD 46769 reveals it to be a monophasic variable with a period of  $4.84 \pm 0.2 \text{ d}$  having an amplitude of 87 ppm and that (sub-)multiples of this period also occur, representing non-sinusoidal variability unde-

tectable by ground-based photometry. The shape of the light curve, with deep and less deep minima that re-occur, leads us to suggest that  $2f$  corresponds with the rotation frequency of the star. This translates into a rotation period of 9.69 d, although we cannot firmly exclude  $4f$  to be the rotation frequency. The stellar parameters and  $v \sin i$  derived by Lefever et al. (2007) do exclude  $f$  to be the rotation frequency.

### 3. Ground-based follow-up data

In order to reach a unique interpretation of the variability found in the CoRoT light curve, with rotation and binarity as most likely scenarios from the CoRoT photometry, we set up a follow-up spectroscopic campaign. We first made use of the HERMES spectrograph (Raskin et al. 2011) attached to the 1.2m Mercator telescope at La Palma and assembled 24 spectra spread over one year, in the period February 2010 until January 2011. The spectra have a high resolving power of 85 000 and a typical average signal-to-noise ratio (S/N) between 100 and 200 in the wavelength range [4000, 5000]Å. At first sight, these data excluded binarity and did not reveal any long-term periodic spectroscopic variability. To enhance the detection capabilities and search for short-period variability as found in the CoRoT light curve, the star was subsequently included in the HARPS Large Programme for the ground-based preparatory and follow-up observations for the CoRoT space mission for more intensive monitoring of several spectra per night during a 25-day campaign. This led to an additional data set of 62 spectra taken in the HARPS HAM mode (resolving power of 115 000), assembled in December 2012 – January 2013. These spectra are of similar S/N as the HERMES spectra. For both campaigns, the exposure time was of the order of a few minutes and was adapted to the atmospheric conditions. We refer to Table 2 for a summary of the spectroscopic observations.

The pipeline-extracted spectral orders of the HARPS measurements were merged and treated together with the pipeline-reduced HERMES exposures. All spectra were normalised following the procedures developed and outlined in Pápics et al. (2011, 2012, 2013). The HERMES and HARPS spectra were found to be fully compatible with each other and could be merged into a single time series data set with a total time span of 1057 d, leading to a Rayleigh limit of  $0.00095 \text{ d}^{-1}$ . From these data, we can exclude that HD 46769 is a binary with an orbital period below several thousands of days.

#### 3.1. Fundamental parameters and abundances

The average HERMES spectrum was used for spectral classification and for the determination of the fundamental parameters and abundances of the star in a comparison with the two spectroscopic standards: HD 184930 (B5 III, Lesh et al. 1968) and HD 191243 (B5 Ib, Morgan et al. 1955). For each of these standards we took one HERMES spectrum. Although Lesh (1968) classified HD 46769 as B8Ib star, our spectra lead us to suggest a spectral type of B5II. This is illustrated in Fig. 3, where we show one HERMES spectrum of HD 46769 downgraded to a resolution of 4000, which is often used for spectral classifications, along with the single HERMES spectrum that we obtained for each of the two spectroscopic standards. The wings of the Balmer lines, as well as the Si II/He I and Mg II/He I line ratios, lead to B5II, although the underabundance of Si and Mg compared to the solar value as discussed below is a caveat for a clear

**Table 3.** Diagnostic lines considered for the quantitative spectroscopic analysis of HD 46769.

	Spectral lines (Å)
H	H $\alpha$ , H $\beta$ , H $\gamma$ , H $\delta$
He I	4026, 4387, 4471, 4713, 4922, 5015, 5048, 5875, 6678
Si II	3856/62, 4128/30, 5041/56, 6347/71
Si III	4552, 4567, 4574
C II	4267, 6578/82
Mg II	4481
O II	4414/16/52, 4317/19/66, 4641/50/61
N II	3995

unique assignment of spectral type (hence the question mark in Fig. 3).

A zoom on the Mg II line profile in a spectrum of these three stars (Fig. 4) reveals that the line profile of HD 46769 is very symmetric and mainly broadened by rotation, while the line profiles of HD 184930 and HD 191243 are slightly asymmetric and have broader wings. The latter phenomenon is often ascribed to the so-called macroturbulent broadening of hot stars (e.g. Howarth et al. 1997, Ryans et al. 2002, Lefever et al. 2007, Markova & Puls 2008, Simón-Díaz et al. 2010, Simón-Díaz 2011). We derived the rotational broadening,  $v \sin i$ , and the macroturbulent broadening assuming an anisotropic radial-tangential model described by  $\Theta_{RT}$ , following the definition by Gray (2005), as implemented by Simón-Díaz & Herrero (2007) and further developed by one of us (SSD). This most recent procedure is based on a combined Fourier transform and profile fitting analysis, which consider rotational and macroturbulent broadening in addition to the intrinsic line broadening and the instrumental broadening due to the characteristics of the spectrograph. The derived values are indicated in Fig. 4 for the three stars. Ignoring macroturbulent broadening for HD 46769 still gives a good fit to the Mg II line if the  $v \sin i$  is increased by  $2 \text{ km s}^{-1}$ . However, it leads to far too narrow wings for the other two stars.

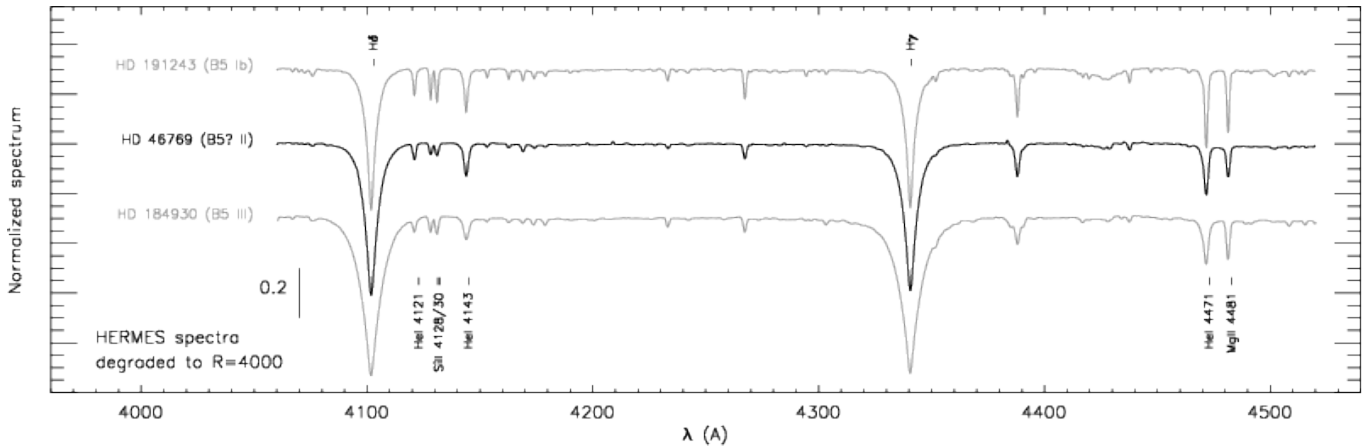
We used the stellar atmosphere code FASTWIND (Santolaya-Rey et al. 1997, Puls et al. 2005) to perform a quantitative spectroscopic analysis of HD 46769 and the two standards, as well as to determine their stellar parameters and abundances. The atomic models presently available for FASTWIND computations are Si, Mg, C, N, and O. The stellar parameter and abundance determinations were done in a two-step process, following an automatic spectroscopic synthesis strategy similar to the one described in Castro et al. (2012). We used a precomputed grid of FASTWIND synthetic spectra, which included lines of H, He I-II, C II-III, N II, O II-III, Si II-IV, and Mg II. More information about the grid characteristics, coverage, and model atoms incorporated in our FASTWIND computations can be found in Castro et al. (2012). Table 3 summarizes the main diagnostic lines considered here. The optical spectrum of HD 46769 includes two ionization stages for the element Si only, but the Si III triplet near 4567Å is very weak. Hence, for the determination of the effective temperature we relied on the dependencies of the lines of the available ions of this stellar parameter. In particular, we have found the He I lines, in combination with H, to be the best  $T_{\text{eff}}$  diagnostics due to their strong dependence on  $T_{\text{eff}}$  and weak dependence on gravity for mid-B stars.

In a first step, we derived the stellar parameters of the star, giving more weight to the H and He I lines and scaling the abun-

**Table 2.** Logbook of the spectroscopy gathered for HD 46769.

Telescope	Instrument	Wavelength Range	HJD start	HJD end	N	$\lambda/\Delta\lambda$	S/N range
1.2m Mercator, La Palma	HERMES	[3700, 9000] Å	5240.472	5581.636	24	85 000	100 – 200
3.6m ESO, La Silla	HARPS	[3700, 6800] Å	6272.664	6297.825	62	115 000	120 – 250

**Notes.** The HJD was subtracted with 2450000, N denotes the number of spectra,  $\lambda/\Delta\lambda$  stands for the resolving power; the listed range in S/N was computed in the wavelength range [4482.6, 4483]Å.



**Fig. 3.** Optical spectrum of HD 46769 compared with the spectrum of two spectroscopic standards, all brought to a common resolution of 4000. The question mark indicates the underabundance of Mg and Si with respect to the solar values, as discussed in the text.

dances of the other elements to the metallicity of the models. The gravity was determined from the H $\gamma$  line. In a second step, we performed the abundance analysis. The stellar parameters were fixed to the values derived in step 1, and the abundances of C, N, O, Si, and Mg were varied independently to create a set of synthetic spectra used for comparison with the observed spectra. In order to have consistency in the comparison for the three stars, we used only one HERMES spectrum for each of them. The results of this analysis are summarised in Table 4. The derived stellar parameters for HD 46769 are in very good agreement with those determined by Lefever et al. (2007). This is as expected, since we used the same stellar atmosphere code and spectroscopic analysis techniques but our work is based on more extended and higher quality spectra. Some line fits for our target star are shown in Fig. 5.

We find an underabundance of Mg and Si with 0.9 dex compared to the solar values (Asplund et al. 2009) with  $\sim 0.2$  dex compared to the Si and Mg abundances of B-type stars in the Orion nebula (Simón-Díaz 2010; Nieva & Simón-Díaz 2011) and, more generally, in the solar neighbourhood (Nieva & Przybilla 2012). While HD 191243 has essentially solar abundances, HD 184930 is also Mg and Si depleted, but less so than HD 46769. We also investigated the impact of the uncertainties in  $T_{\text{eff}}$  and  $\log g$  on the derived abundances and found that a variation of 1000 K and 0.1 dex in both parameters leads to a maximum variation of  $\sim 0.2$  dex in the Si and Mg abundances. This is well below the variation needed to bring the measured Mg and Si abundances derived for HD 46769 in agreement with the solar values.

We come to the conclusion that HD 46769 is an evolved, moderately rotating ( $v \sin i \approx 70 \text{ km s}^{-1}$ ) B5II star with a normal He abundance but with an underabundance of 0.9 dex in Mg and Si. On the main sequence, such Si and Mg deficiencies occur in the so-called HgMn stars (e.g. Smith 1993), which

are slow rotators with accompanying strong overabundances in elements such as Hg, Mn, Cr, Ti, and Y. They are, however, underabundant in He. These properties are interpreted in terms of atomic diffusion processes, more particularly the competition between gravitational settling and radiative levitation. It remains unsure whether binarity and/or magnetic fields play a role in the creation and time-evolution of the chemical spots on the surface (e.g. Briquet et al. 2007, 2010; Kochukhov et al. 2011, Hubrig et al. 2012). We checked explicitly for Mn and Hg lines in the average spectrum of HD 46769, but could not detect any. Other than the HgMn stars, also Bp stars exhibit surface spots of elements such as He and Si, giving rise to periodic mmag-amplitude photometric variations and clear line-profile variability with periodicities of days (e.g. Briquet et al. 2004). Below, we investigate if HD 46769 is a line-profile variable.

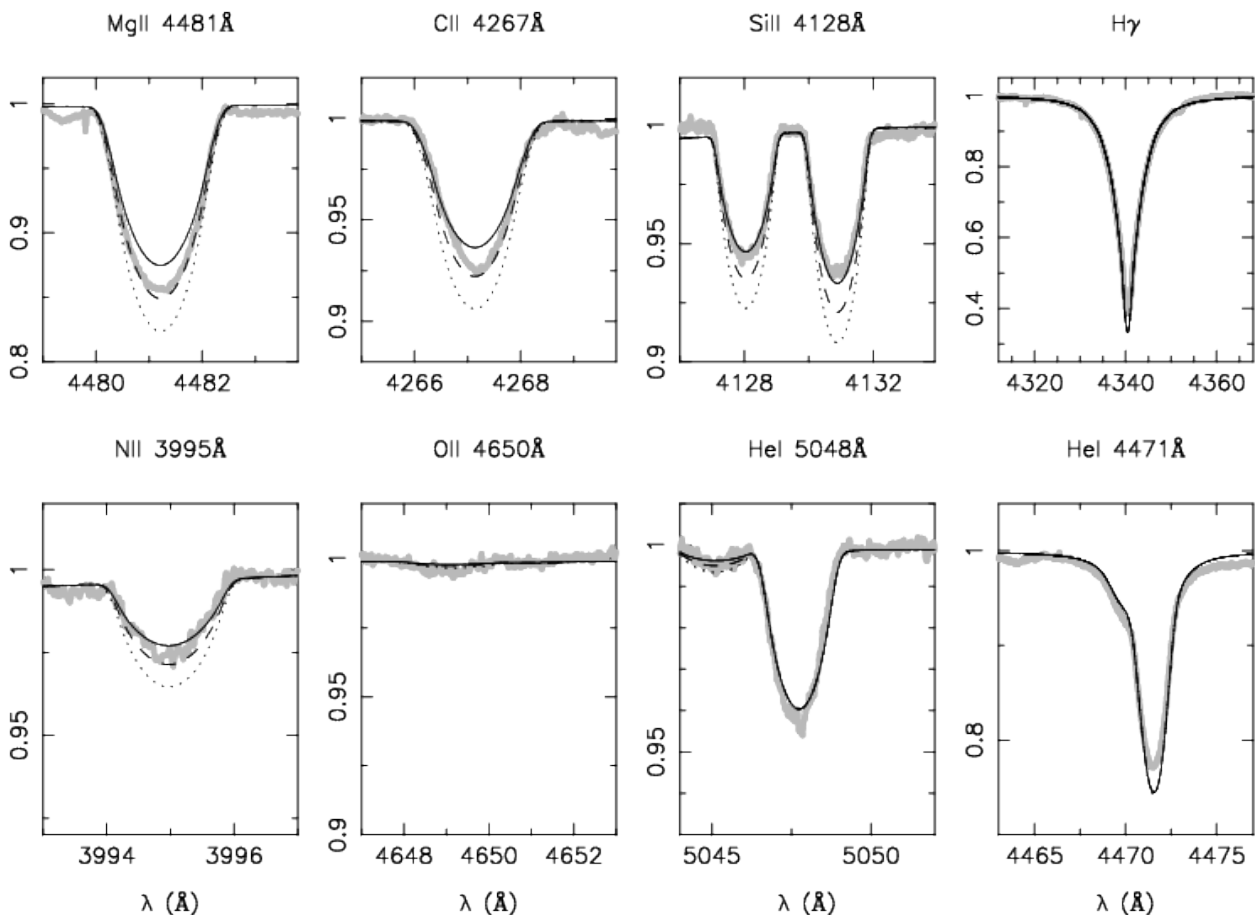
### 3.2. Time series analysis

In order to search for short-term periodicity in the spectroscopy and compare it with the detected photometric variability, we first did a visual inspection of the line profiles of various chemical elements in the observed spectra for the lines listed in Table 3, as well as some additional metal lines. This did not reveal any obvious variability in the line profile shapes. In order to quantify the line-profile variations, we performed a line-profile analysis by means of line moments (e.g. Aerts et al. 1992 and Briquet & Aerts 2003) and the pixel-by-pixel method (Zima 2006). We refer to Aerts et al. (2010a, Chapter 6) for full details of such analyses. An important aspect of such applications is the selection criterion to identify the best spectral lines, particularly when line variability is not obvious by visual inspection. In practice, one looks for the least blended lines with considerable strength and least affected by pressure broadening and effective temperature variations; see De Ridder et al. (2002) and Aerts & De Cat

**Table 4.** Derived abundances of HD 46769, HD 191243, and HD 184930.

Parameter	HD 46769	HD 191243	HD 184930	Solar	B stars in solar neighbourhood
$T_{\text{eff}}$ (K)	13 000±1 000	14 000±1 000	13 000±1 000	—	—
log $g$	2.7±0.1	2.5±0.1	3.5±0.1	—	—
Si	-5.49	-4.30	-5.10	-4.45	-4.50
Mg	-5.47	-4.60	-5.10	-4.36	-4.43
C	-4.11	-3.30	-3.80	-4.53	-3.65
N	-3.92	-3.80	—	-4.13	-4.18
O	—	-3.30	—	-3.27	-3.23

**Notes.** Metal abundances are listed as  $\log(N_{\text{ion}}/N_{\text{tot}})$  and have an uncertainty of 0.2 dex. In each of the cases,  $n(\text{He})/n(\text{H})=0.1$  was kept fixed and the mass loss parameter  $\log Q$  (see Kudritzki & Puls 2000 for a definition) was below -13.5. The microturbulent velocity was  $10 \text{ km s}^{-1}$  for HD 46769 and HD 191243, while it was  $5 \text{ km s}^{-1}$  for HD 184930. The solar values were taken from Asplund et al. (2009) and those of B stars in the solar neighbourhood from Simón-Díaz (2010), Nieva & Simón-Díaz (2011), and Nieva & Przybilla (2012). It was not possible to deduce an O abundance for HD 46769 and HD 184930 because their spectra do not show O II lines, while FASTWIND does not yet have the option to predict O I lines.

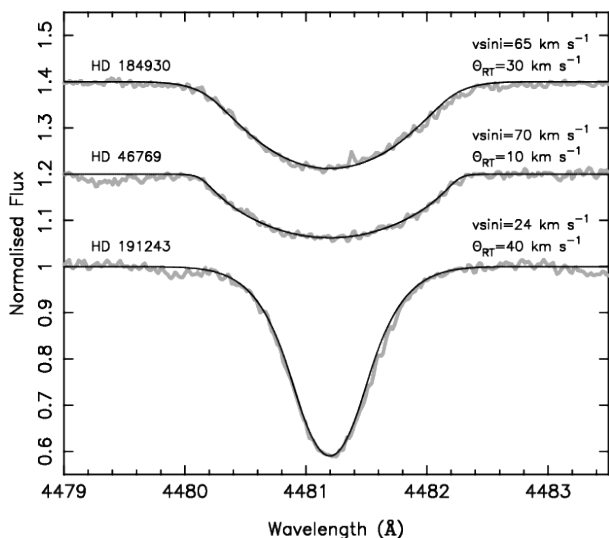


**Fig. 5.** Comparison of the line predictions for HD 46769, based on FASTWIND atmosphere models for the stellar and abundance parameters listed in Table 4 (dashed). The metal abundances were lowered with 0.2 (full) and increased with 0.2 (dotted). In each case, the He abundance was fixed according to  $n(\text{He})/n(\text{H})=0.1$ .

(2003) for thorough discussions on these issues for line-profile studies of confirmed pulsating B stars. For the effective temperature of HD 46769 and from the gathered spectroscopy, we identified the Mg II doublet near  $4481 \text{ \AA}$  as best suited in terms of S/N level, followed by the C II doublet near  $4267 \text{ \AA}$  and the Si II doublet near  $4130 \text{ \AA}$ . We show the results obtained for the Mg II, which are completely confirmed by the other metal lines.

Fig. 6 shows the time series of the Mg II line deduced from the HARPS spectra, along with the time-average profile of that

line and the residuals after subtraction of the average represented in a colour-scale plot phased according to the most likely rotation period of the star. It can be seen that no line patterns stand out. This was also the conclusion from a 2D frequency analysis with the pixel-by-pixel method done with the software package FAMIAS (Zima 2008). Subsequently, we resorted to dedicated quantities to study line-profile variability, according to the definition by Aerts et al. (1992). The time series of the zeroth-order moment (which is the equivalent width of the line, further ab-



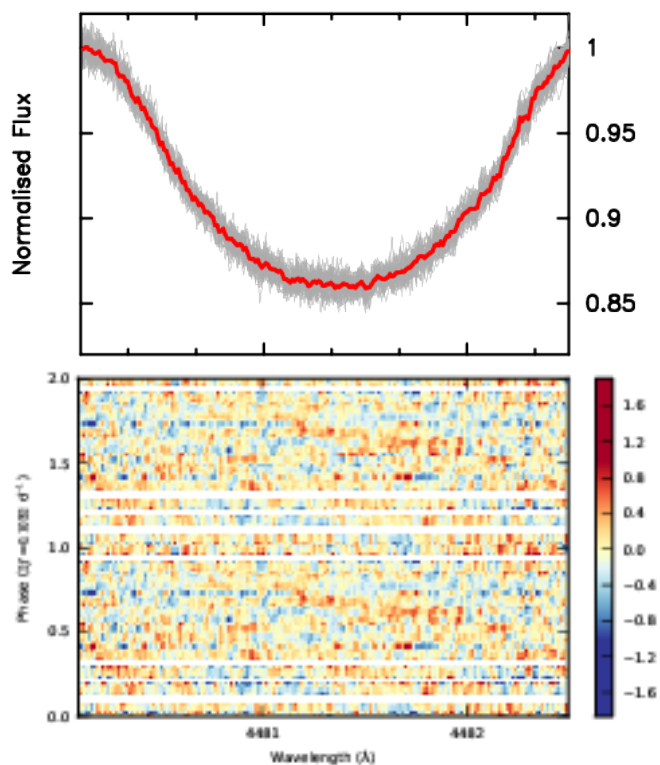
**Fig. 4.** Zoom in on the Mg II line at 4481 Å of the three stars whose entire optical spectrum was shown in Fig. 3. The values for the projected rotation velocity  $v \sin i$  and for the macroturbulent line broadening in the approximation of a radial-tangential model ( $\Theta_{RT}$ ) are indicated in  $\text{km s}^{-1}$ .

breveated as EW) and the first moment (indicated as  $V_{rad}$  since it is connected with the radial velocity) derived from the Mg II line deduced for all 86 HERMES and HARPS spectra are shown in Fig. 2, where they were phased with the dominant frequency found in the CoRoT light curve. The outcome of the Scargle periodograms for the EW and  $V_{rad}$  are shown as insets in Fig. 7. Both these figures illustrate again that HD 46769 does not reveal obvious periodic spectroscopic variability over a total time base of 1057 d. Noteworthy is that the maximum frequency peak in  $V_{rad}$  corresponds with  $1 - 4f$  within the frequency resolution, having an amplitude of  $0.15 \pm 0.05 \text{ km s}^{-1}$  at  $2.3\sigma$ . We cannot relate any of the highest peaks in the Scargle periodogram of the EW to the photometric frequencies. Enforcing a fit with one of the photometric frequencies leads to the highest amplitude for  $4f$ , but the least-squares fit gives an insignificant amplitude estimate at only  $1.9\sigma$ .

Even though the frequency  $f$  or its multiples is not significantly present in the spectroscopy alone, we can still check if its behaviour is consistent with the CoRoT data or is only connected with noise. We thus applied the same analysis as already done by Aerts et al. (2006) in their search for low-amplitude variability in the  $\beta$  Cep star  $\delta$  Ceti by combining MOST and Hipparcos space photometry with ground-based spectroscopy. Since we expect flux variations to occur in both the photometry and in the EW of spectral lines, we multiplied the EW Scargle periodogram with the one of the CoRoT and of the Hipparcos data, after dividing each of those three Scargle periodograms by their maximum, i.e. we computed

$$P(f) = \frac{P_{Scargle}^{EW}(f)}{P_{Scargle}^{EW}(f_{max})} \cdot \frac{P_{Scargle}^{Hipparcos}(f)}{P_{Scargle}^{Hipparcos}(f_{max}^{Hipparcos})} \cdot \frac{P_{Scargle}^{CoRoT}(f)}{P_{Scargle}^{CoRoT}(f_{max}^{CoRoT})}.$$

In this way, we give equal weight to the CoRoT, Hipparcos, and HERMES+HARPS data in the search for consistent frequencies in those three completely independent data sets. The idea is that all the peaks due to noise will keep on producing a noisy periodogram after multiplication of the normalized periodograms. In contrast, peaks that are consistently present in the data, even

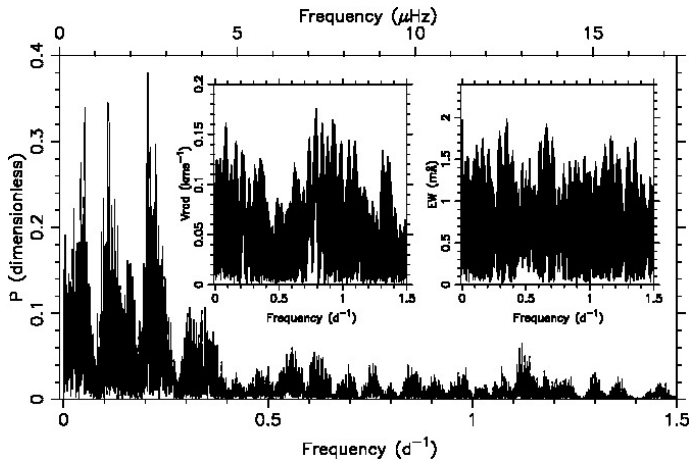


**Fig. 6.** Top: time series of the Mg II line at 4481 Å distilled from the HARPS spectra, with a time base of 25.2 d. The red thick line is the average Mg II line. The bottom panel shows the residual Mg II line profiles with respect to the average line profile as a function of the rotational phase and represented in a colour scale, where we applied a sliding boxcar smoothing with a bin width of 0.05 Å and a step size of 0.005 cycles; for better visibility we show two cycles.

though not significantly above the noise level in the individual Scargle periodograms, will still keep a value that is clearly above the noise level when full compatibility between the three Scargle periodograms occurs. Prior to normalization, the Scargle periodograms were computed with a frequency step according to the best Rayleigh limit, which occurs for the Hipparcos data set. The outcome of this product of the three normalized Scargle periodograms, which is a dimensionless quantity with a value between zero (bad frequency) and one (frequency dominant in all three independent data sets), is shown in the main panel of Fig. 7. It can be seen that three sets of peaks stand out, i.e. that the frequencies deduced from the CoRoT light curve “survive” this consistency test. These occur at  $0.2063$ ,  $0.1083$ , and  $0.0522 \text{ d}^{-1}$ , where the first one coincides with  $4f$  and all three are equal to the frequencies in the CoRoT data listed in Table 1 within the frequency resolution. Hence, the spectroscopy and Hipparcos data by themselves do not reveal significant frequencies but are consistent with the low-amplitude periodic variability detected in the CoRoT light curve.

### 3.3. Spectropolarimetric measurements

Given the photometric variability and the abundance pattern of HD 46769, as well as the debate about the role of magnetic fields in chemically anomalous rotationally variable main-sequence stars, we made use of the Narval spectropolarimeter installed at the 2 m TBL telescope at the Pic du Midi in France to search for the presence of a magnetic field in HD 46769. On December



**Fig. 7.** Product of the Scargle periodograms of the CoRoT and Hipparcos light curves and of the EW (right inset) of the Mg II 4481 Å line, after normalizing those three individual periodograms through division by the maximal amplitude (unitless, value between 0 and 1). The Scargle periodogram of the radial velocity of the Mg II 4481 Å line is shown in the left inset for comparison.

2012 we obtained two Stokes V spectra computed by constructively combining four subexposures of 900 seconds, each taken in different configurations of the polarimeter wave plates. The data reduction was performed using LIBRE-ESPRI (Échelle Spectra Reduction: an Interactive Tool; Donati et al. 1997).

For each of the two measurements, we applied the least-squares deconvolution (LSD) technique (Donati et al. 1997) to the photospheric spectral lines in the whole échelle spectrum ( $\lambda \in [3750, 10500] \text{ \AA}$ ) in order to construct single averaged Stokes I and V profiles with an increased S/N ratio. To this end, we used a mask including photospheric He and metal lines of various chemical elements, which we created from the Kurucz atomic database and ATLAS 9 atmospheric models of solar abundance (Kurucz 1993) for  $T_{\text{eff}}=13000 \text{ K}$  and  $\log g=3.0$ , with intrinsic line depths larger than 0.1. Two adjustments were applied to the original mask. First, we removed lines from the mask that are not suited for the LSD analysis, i.e. those that are blended with the Balmer lines or with atmospheric telluric lines and weak lines that are indistinguishable from the noise. Second, the relative depths of all spectral lines used in the mask were modified to correspond to those of the observed spectral lines. The final mask contains 311 lines, and the two resulting LSD Stokes V profiles have S/N values of 5864 and 5309 per  $2.6 \text{ km s}^{-1}$  pixels. A visual inspection of each of them shows no signature of a magnetic field.

We used the LSD Stokes I and V profiles to compute the line-of-sight component of the magnetic field integrated over the visible stellar surface, i.e. the longitudinal magnetic field in Gauss, given e.g. by Eq. (1) of Wade et al. (2000). In this equation we used a mean wavelength  $\lambda=500 \text{ nm}$  and a mean Landé factor  $g=1.13$ . For the integration limits, a range of  $70 \text{ km s}^{-1}$  around the line centre was adopted. We obtained longitudinal field values of  $B_{l1}=-9.4 \pm 21.0 \text{ G}$  and  $B_{l2}=22.1 \pm 22.9 \text{ G}$  at  $\text{Mid-HJD}_1=2456271.657354$  and  $\text{Mid-HJD}_2=2456272.627450$ , respectively, i.e. the longitudinal field values are consistent with the absence of a magnetic field.

As we only have two measurements available, we cannot definitively exclude the presence of a field in the star. Indeed, it is

possible that a magnetic field exists with a strength below our detection level. Therefore we performed a Monte Carlo analysis of the spectropolarimetric data to derive the upper limit of the polar magnetic field that could be present at the surface of HD 46769.

To perform our simulation we assumed a centred oblique dipole field model, for which the free parameters are the stellar inclination angle  $i$  (in degrees), the magnetic obliquity angle  $\beta$  (in degrees), the dipole magnetic intensity  $B_{\text{pol}}$  (in Gauss), and a phase shift  $\Delta\Phi$  compared to the adopted stellar ephemeris. In our model, we use Gaussian local intensity profiles with a width calculated using the resolving power of Narval and a line-broadening value of  $20 \text{ km s}^{-1}$ , representing the overall effect of the intrinsic line widths as well as micro- and macroturbulence. The depth is determined by fitting the observed LSD Stokes I profiles. The local Stokes V profiles are calculated assuming the weak-field case and using the weighted mean Landé factor and wavelength of the lines derived from the LSD mask (see above). The synthetic Stokes V profiles are obtained by integrating over the visible stellar surface by using a linear limb-darkening law with a parameter equal to 0.386 (Claret & Bloemen 2011). The Stokes V profiles are normalised to the continuum intensity.

We considered the two possible values for the rotation period  $P_{\text{rot}} = 4.84$  or  $9.69 \text{ d}$  and various values for the polar field strength. We then simulated 1000 dipole models with random values of  $i$ ,  $\beta$ , and  $\Delta\Phi$  for each  $B_{\text{pol}}$  value. We simulated the two Stokes V profiles simultaneously with a noise corresponding to the S/N ratio of the Narval data. We then computed among each set of 1000 models the chances of detecting the field in at least one of the two measurements. We find that there was a 50% (90%) chance of detection at  $3\sigma$  of a field with any oblique dipolar configuration and a polar strength above  $\sim 290 \text{ G}$  ( $\sim 515 \text{ G}$ ) for the case of  $P_{\text{rot}} = 9.69 \text{ d}$ . This is illustrated in Fig. 8, where we show the probability of detection of a field according to its strength in at least one of the two available measurements of HD 46769, assuming  $P_{\text{rot}} = 9.69 \text{ d}$ . Results are similar, whichever of the two possible rotation periods we consider.

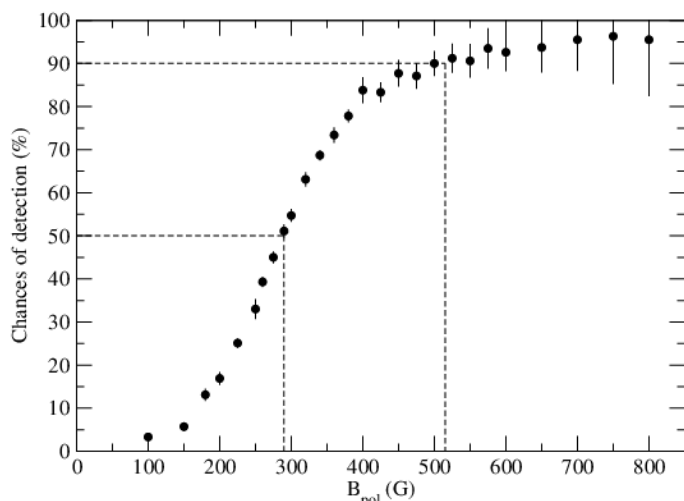
We conclude that an oblique magnetic dipole field of a few hundred Gauss at the poles could be present while it remained undetected in the two Narval measurements. A stronger magnetic field can be excluded, unless the hypothetical field were more complex than dipolar. However, for massive non-Bp stars, detected fields are usually observed to be dipolar (e.g. Neiner et al. 2003) and also theoretically predicted to be dipolar (Duez & Mathis 2010). Non-dipolar examples (e.g. Kochukhov et al. 2011) remain very rare so far. Our results are consistent with the scarcity of magnetic fields in OBA supergiants, despite intensive efforts to detect them (e.g. Verdugo et al. 2003, Hubrig et al. 2013, Shultz et al. 2013).

## 4. Discussion

All the data taken together lead us to conclude that HD 46769 is a low-amplitude, rotationally modulated, photometrically variable B5II supergiant. The photometric variability of some 100 ppm level is interpreted in terms of low-contrast chemical and/or temperature inhomogeneities. This gives rise to rotational modulation in the light curve, which remains undetectable in high-precision spectroscopy alone.

The fundamental parameters of the star place it just outside pulsational instability regions in the ( $T_{\text{eff}}, \log g$ ) diagram (e.g. Fig. 10 of Lefever et al. 2010), although the error bars cover the red edge of excited gravity modes computed by Saio et al. (2006). Despite the absence of pulsations with a detection





**Fig. 8.** Chances of detection of a magnetic field in at least one of the two Narval measurements of HD 46769, depending on the polar field strength, assuming  $P_{\text{rot}} = 9.69$  d.

threshold of some 10 ppm in photometric precision, the star is an interesting object for its underabundance in Mg and Si, and also in the context of the occurrence of macroturbulence in evolved massive stars whose physical cause remains unknown. Indeed, HD 46769 is a moderate rotator, and its spectral lines have very sharp wings in comparison with more luminous stars of similar temperature. We found an anisotropic radial-tangential macroturbulence of less than  $10 \text{ km s}^{-1}$ , i.e. far lower than the projected rotational velocity broadening. A good fit to the line wings would also be achieved without considering macroturbulent broadening, increasing the  $v \sin i$  by only a few  $\text{km s}^{-1}$ . In contrast, to be able to explain their spectral line wings, the majority of the evolved variable B supergiants require large macroturbulent velocities, several of which are supersonic and above the rotational velocity. Aerts et al. (2009) suggested that the detected high level of macroturbulence needed to explain the line wings of massive supergiants in high-precision spectroscopy could be due to the collective effect of numerous low-amplitude gravity-mode pulsations, which are unresolved individually and give rise to a global line broadening in single-snapshot spectra. This idea is currently being investigated in more detail from a systematic long-term observational study of a large sample of OB-type stars from high-precision time-series spectroscopy (e.g. Simón-Díaz et al. 2010, 2012 and its follow-up work, in preparation). Our results for the case of HD 46769 are consistent with this interpretation in the sense that it concerns an evolved non-pulsating B-type supergiant with sharp line wings, for which macroturbulence does not need to be invoked to get a good fit of the spectral line wings.

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