Efficient surface formation route of interstellar hydroxylamine through NO hydrogenation. II. The multilayer regime in interstellar relevant ices

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Hydroxylamine (NH2OH) is one of the potential precursors of complex pre-biotic species in space. Here, we present a detailed experimental study of hydroxylamine formation through nitric oxide (NO) surface hydrogenation for astronomically relevant conditions. The aim of this work is to investigate hydroxylamine formation efficiencies in polar (water-rich) and non-polar (carbon monoxide-rich) interstellar ice analogues. A complex reaction network involving both final (N2O, NH2OH) and intermediate (HNO, NH2O−, etc.) products is discussed. The main conclusion is that hydroxylamine formation takes place via a fast and barrierless mechanism and it is found to be even more abundantly formed in a water-rich environment at lower temperatures. In parallel, we experimentally verify the non-formation of hydroxylamine upon UV photolysis of NO ice at cryogenic temperatures as well as the non-detection of NC- and NCO-bond bearing species after UV processing of NO in carbon monoxide-rich ices. Our results are implemented into an astrochemical reaction model, which shows that NH2OH is abundant in the solid phase under dark molecular cloud conditions. Once NH2OH desorbs from the ice grains, it becomes available to form more complex species (e.g., glycine and β-alanine) in gas phase reaction schemes. © 2012 American Institute of Physics. [http://dx.doi.org/10.1063/1.4738893]

I. INTRODUCTION

The observed chemical complexity in space results from the cumulative outcome of gas, grain, and gas-grain interactions. Recent studies indicate that a substantial part of the stable and complex species identified so far, form on icy dust grains.1,2 Interstellar grains, indeed, act as micrometer-sized cryo-pumps providing surfaces onto which gas-phase species can accrete, meet, and react. During the last decade several independent laboratory studies showed the relevance of such astrochemical solid-state reactions. Meanwhile, the surface formation of the bulk of the identified interstellar ices (i.e., water, methanol, carbon dioxide, formaldehyde, and formic acid) has been confirmed through subsequent H-atom addition reactions to simple species, like CO- and/or O2-ices.3–10

In space, non-energetic atom addition induced solid-state chemistry occurs mostly at low temperatures (~10 K), i.e., in the innermost part of circumstellar clouds where newly formed molecules are, to a great extent, shielded from radiation by dust particles. These regions are part of collapsing envelopes feeding new stars—young stellar objects (YSOs)—and provide the original material from which comets and ultimately planets are made.11 Since solid-state neutral-neutral reactions often do not require activation energy,1,2 they play a key role in the chemistry of dark interstellar clouds. In the final stages of star formation energetic processing—like UV photolysis, cosmic ray irradiation, and thermal processing—also contributes to surface reaction schemes by adding external energy into the ice. It has indeed been demonstrated that complex organic molecules and even amino acids are formed after energetic processing of interstellar ice analogues.12,13 Therefore, laboratory based studies of the solid-state formation of organic material is of considerable interest, since efficient reaction routes provide a recipe to form pre-biotic species in star and planet forming regions.14

Hydroxylamine (NH2OH) is an important precursor species in the formation of amino acids.15 Several formation schemes of hydroxylamine have been proposed in the past. Blagoevic et al.15 describe NH2OH formation through UV photolysis of NH3 + H2O and/or “NO + 3H” ices. Charnley et al.14 suggested that hydroxylamine is formed through non-energetic hydrogenation of NO ice under dark cloud conditions. Nitric oxide has been detected in the gas phase towards many dark and warm clouds with relative abundances of 1 × 10−8–1 × 10−7 with respect to H2.16–19 Detailed astrochemical models14,19,20 indicate that the main formation route for NO in the gas phase is a neutral-neutral reaction N + OH → NO + H, while the main gas-phase destruction channel is NO + N → N2 + O. Under dense cloud conditions, gas-phase NO can accrete, like CO, on the surface of dust grains and is, therefore, expected to participate in the solid-state chemical network leading to the formation of N- and NO-bearing species. Moreover, solid NO may also be formed...
through surface reactions in the early stages of quiescent dark clouds. Although NO has the potential to be one of the main precursors of complex molecules in space, solid-state reaction schemes involving NO-bearing ices have not yet been studied in detailed experiments. Recently, we presented for the first time hydroxylamine formation via a non-energetic solid-phase route—pure NO + H on crystalline H$_2$O ice as well as on other different substrates$^{21}$—and we showed the astrochemical impact of that scheme. In the present paper, the focus is on the investigation of the underlying physical-chemical processes that lead to surface hydroxylamine formation in the multilayer regime. Thus, experiments are extended to more realistic and interstellar relevant polar (NO:H$_2$O + H/D) and non-polar (NO:CO + H/D) environments. In Secs. III and IV, we investigate the effect of temperature, ice composition, and H-atom flux on the NH$_2$OH and HNO formation efficiencies. We also discuss experiments in which NO, NO:H$_2$O, and NO:CO ices are UV irradiated with Ly-$\alpha$ light. The astrophysical implications of this work are briefly discussed. In an accompanying and complementary paper$^{22}$ (published back-to-back with the present one), NO hydrogenation reactions are studied with the focus on the submonolayer regime using interstellar relevant substrates, i.e., amorphous silicate and crystalline H$_2$O ice.

II. EXPERIMENTAL DETAILS AND DATA REDUCTION

A. Experimental procedure

The experiments are performed, using two similar ultra-high vacuum (UHV) setups: a surface reaction simulation device (SURFRESIDE), optimized to study H-atom addition reactions, and a cryogenic photoproduct analysis device (CRYOPAD), designed for UV photolysis experiments of interstellar ice analogues.

1. SURFRESIDE

Thermal H/D-atom addition reactions are investigated using SURFRESIDE which consists of a stainless steel UHV main chamber and an atomic line. Experimental details are available in Refs. 4, 6, and 9. A rotatable gold-coated copper substrate is placed in the centre of an UHV chamber. The room temperature base pressure in the main chamber is $<3.5 \times 10^{-10}$ mbar. The substrate temperature is controlled from 12 to 300 K using a He closed-cycle cryostat. Ice deposition proceeds under an angle of 45$^\circ$, with controllable rates from 0.5 to 2.5 Langmuir per minute (L/min, where 1 L = 1.3 $\times 10^{-6}$ mbar s$^{-1}$). A quadrupole mass spectrometer (QMS) placed behind the substrate is used to monitor the main-chamber gas composition. The atomic line faces the sample and comprises a well-characterized hydrogen thermal cracking source$^{23}$ used to produce non-energetic H/D atoms from H$_2$/D$_2$ molecules. Hydrogen/deuterium molecules are cracked by passing through a capillary pipe surrounded by a heated tungsten filament. Dissociation of molecules occurs through collisions with the hot (1850 °C) walls of the capillary pipe with a degree of dissociation $>35%$. A nose-shaped quartz pipe is placed along the path of the beam and is used to thermalize both H/D atoms and non-dissociated H$_2$/D$_2$ molecules to room temperature before they reach the surface of the ice sample. H/D-atom fluxes are measured quantitatively by placing the QMS at the substrate position and are chosen between $7 \times 10^{12}$ and $3 \times 10^{13}$ atoms cm$^{-2}$ s$^{-1}$. This procedure has been described in detail in Ref. 24.

A high vacuum glass manifold with a base pressure $<5 \times 10^{-5}$ mbar is used for gas mixture preparation. To avoid water contamination, the glass line is connected to a liquid nitrogen trap and is flushed well by mixture components. A new mixture is prepared for each experiment. After preparation, the mixture is introduced to a pre-pumped stainless steel dosing line ($<1 \times 10^{-5}$ mbar). The metal line is typically filled with a gas pressure of 30 mbar and is kept isolated for the duration of the experiment. The deposition rate is controlled by a precise all-metal leak valve. The gases used in this work are: NO (Linde 99.5%), CO (Linde 99.997%), N$_2$O (Praxair 99.5%), NO$_2$ (Linde 99%), and a Milli-Q water sample degassed under high vacuum conditions.

The ice composition is monitored in situ by means of reflection absorption infrared spectroscopy (RAIRS) in the range: 700–4000 cm$^{-1}$ (14–2.5 $\mu$m) with a spectral resolution of 0.5 cm$^{-1}$ using a Fourier transform infrared spectrometer (FTIR). Two different experimental procedures are applied here. In “pre-deposition” experiments, ices are first deposited on a gold surface and are subsequently exposed to a thermal H/D-atom beam. In this case, RAIR difference spectra are acquired during hydrogenation/deuteration of the sample with respect to a background spectrum of the initial deposited ice at low temperature. In the case of “co-deposition” experiments, nitrogen oxide bearing ices are continuously deposited simultaneously with H/D atoms. The formation of intermediate species and final products is monitored in the ice by changing ice deposition rates and H/D-atom fluxes, as discussed in Ref. 25. In this case, RAIR difference spectra are acquired during co-deposition with respect to a background spectrum of the bare gold substrate at a low temperature.

At the end of every H/D-atom exposure, a temperature programmed desorption (TPD) experiment is performed. Because the QMS is situated behind the substrate, RAIR spectra and QMS data cannot be taken simultaneously during the heating phase, and crucial experiments are performed twice in order to interpret TPD results using both RAIRS and QMS techniques. In the first experiment, the sample is rotated 180$^\circ$ to face the QMS and it is heated linearly at a rate of 1 K/min. Desorbed species are subsequently detected as a function of sample temperature using the QMS. In the second experiment, the sample is not rotated and RAIR spectra of the ice are acquired every 5 K while heating at a rate of 1 K/min. In this way infrared spectroscopic and mass spectrometric information can be combined. It should be noted that a TPD experiment implicitly leads to the morphological modification and eventual destruction of the ice. Therefore, TPD data are mainly used in our experiments as an additional tool to constrain the low temperature RAIR results.

2. CRYOPAD

UV photolysis experiments are performed using CRYOPAD which is described in detail in Ref. 26 and references
that is used for SURFRESIDE, covering 700–4000 cm
10
tion are monitored by means of the same FTIR spectrometer
1. The room temperature base pressure of the system is
9 mbar and the temperature of the substrate is controlled
15 and 200 K. Ice deposition proceeds at an angle of 90° and a rate of 1.5 L/min. After deposition, the ice
film is irradiated by UV light from a broadband hydrogen mi-
crowave discharge, which peaks at 121 nm (Ly-α) and cov-
ers the range from 115 to 170 nm (7–10.5 eV) with an in-
cidence angle of 45°. Photolysis products upon UV irradia-
tion are monitored by means of the same FTIR spectrometer
10. QMS is incorporated into the setup in order to monitor molecules in the gas phase and specifi-
cally photodesorption products. At the end of a UV irradiation
experiment, a TPD experiment is routinely performed. The
setup is constructed such that the QMS already faces the sam-
ple, and consequently, during a TPD it is possible to monitor
ice constituents spectroscopically and mass spectrometrically.
The heating rate used for these experiments is the same as for
SURFRESIDE (1 K/min). The UV flux is measured indirectly
by calibration via the previously studied photodesorption rate
of a pure CO ice: carbon monoxide ice is exposed to Ly-α
light and the number of photodesorbed CO molecules is de-
termined using the RAIR difference spectra. The UV flux is
then derived, assuming that the CO photodesorption rate ob-
tained here and previously in Ref. 27 are identical. Typical
UV fluxes amount to $1.5 \times 10^{14}$ photons cm$^{-2}$ s$^{-1}$.

Gas mixtures used during experiments with CRYOPAD
are prepared separately in the same glass manifold as de-
scribed above. The mixtures are then introduced into a pre-
pumped stainless steel dosing line (<1 × 10$^{-4}$ mbar) that is
always filled twice in order to avoid the decomposition of gas
mixture components on the metal walls during the preparation
hours: a first time (at the beginning of the experiment) to se-
lect at room temperature the desired deposition rate by adjust-
ing a precise all-metal leak valve and monitoring the increase
in pressure in the main chamber; and a second time (right be-
fore deposition) using the same pressure used the first time
to fill the dosing line. When not in use, the deposition line
is kept under high vacuum. In this way highly reproducible
deposition rate can be guaranteed.

The experiments at SURFRESIDE and CRYOPAD fo-
cus on selected surface reactions involving the hydrogena-
tion and UV irradiation of pure NO ices and mixtures of NO
with CO, H$_2$O and N$_2$ over a wide range of laboratory condi-
tions including different atomic fluxes, ice temperatures, co-
deposition rates, and mixture ratios. A number of control ex-
periments are performed to prove that newly detected species
are formed upon H/D-atom exposure or UV photolysis on top of
or in the ice sample, and that they are not the result of con-
amination or background gas-phase reactions. All exper-
iments are summarized in Table I. Moreover, in order to un-
ambiguously identify absorption bands of possible reaction
products, we have performed separate experiments to acquire
spectra of pure N$_2$O as well as NO:N$_2$O$_3$ and NO$_2$N$_2$O$_4$
ice mixtures (see Table II) to make direct comparison
possible.

B. Data analysis

Straight baseline segments are subtracted from all ac-
quired spectra. Subsequently, the bands areas corresponding
to species present in the ice are integrated. As previously
discussed in Refs. 4 and 6, transmission band strengths can-
not be used directly to derive column densities of species in
reflection experiments. Moreover, isothermal desorption ex-
periments, which provide us with absorbance per monolayer,
cannot be performed for the unstable intermediates that are
detected during co-deposition experiments. Therefore, we re-
frain from deriving column densities and instead follow the
formation trends of the detected species only by integrat-
ing the corresponding band areas as a function of time (see Sec.
III). As a consequence, a full quantitative character-
ization is not possible. However, we can compare forma-
tion trends of the same species from different experiments
and thereby derive information on temperature dependence,
ice composition effects, and the reaction network. All the
assigned absorption features that correspond to nitrogen-
bearing species and present in our experiments are summa-
rized in Table II.

All absorbance features chosen for data analysis have a
good signal-to-noise ratio and do not overlap with other ab-
sorbance features. The range of integration is set manually
for each species and kept the same for all the spectra acquired
during a single experiment. The error bars that are indicated
in the figures listed in Sec. III are estimated as follows: several
blank spectra, i.e., without deposited ice but with a cold sub-
strate, are acquired before each experiment. The correspond-
unding baseline segments are subtracted from the blank data. The
noise area in the blank spectra is integrated over the same fre-
quency range that is used for band integration of the selected
species. The uncertainties are then derived by averaging the
values of the measured blank area for each selected range. Al-
though these error bars do not include deviations in the base-
line subtraction procedure, they provide lower limits for the
detectable signal.

III. RESULTS AND DISCUSSION

A. Hydrogenation of pure NO ice

Figure 1(a) shows RAIR difference spectra recorded with
SURFRESIDE of a pre-deposited NO ice after exposure to a
thermal H-atom beam at 15 K. The negative peaks around
1776 and 1865 cm$^{-1}$ correspond to cis-(NO)$_2$ (Refs. 28 and
29) used up during hydrogenation of the ice. Due to the small
exothermicity of NO dimerization (14 kJ per mol) (Ref. 35)
and a fast radical-radical recombination reaction rate, nitric
oxide forms solid cis-(NO)$_2$ immediately after deposition.
Therefore, NO monomers are detected in the ice by means of
RAIRS only when NO is trapped in an ice matrix of other species (see below). Thus, hydrogenation/deuteration of a pure NO ice involves as a first step the hydrogena-
tion/deuteration of (NO)$_2$ dimers. The positive peaks in
Fig. 1(a) indicate the presence of newly formed species in the
ice. However, their identification is not clear a priori because of their low signal-to-noise ratio. A possible reason for this
low coverage of processed ice can be found in the structure
of the ice. Crystalline (NO)2 has a compact monoclinic unit cell with a space group P21/a,36 which is more comparable to a CO than an O2 ice structure.37,38 Although our deposited ice is not necessarily crystalline, the local structure of (NO)2 agglomerates is probably close to crystalline. As shown in Refs. 4 and 9 a compact ice structure prevents H/D atoms from penetrating deep into the bulk of the ice. As a consequence, in a pre-deposition experiment only the first few monolayers are actively involved in surface reactions, the final yield of H/D atom addition products is low and no intermediate products are detected.

Figures 1(b) and 1(c) show NO/H-atom co-deposition experiments at 15 K. The main advantage of co-deposition experiments is the ability to change the NO/H(D) ratio in order to: (1) trap non-fully hydrogenated products inside the growing ice by using a high NO flow with a low H(D)-atom flux (Figs. 1(c) and 1(e)) or (2) deposit a fully hydrogenated growing ice by using a high NO flow with a low H(D)-atom flux (Figs. 1(b) and 1(d)). Growing a thick hydrogenated ice during co-deposition experiments increases the column densities (Figs. 1(b) and 1(d)). Growing a thick hydrogenated ice during co-deposition experiments increases the column densities.

<table>
<thead>
<tr>
<th>Depositing ice composition</th>
<th>Type</th>
<th>Temperature (K)</th>
<th>Deposition rate (L/min)</th>
<th>H/D flux (atom cm⁻² s⁻¹)</th>
<th>H₂/D₂ flux (mol cm⁻² s⁻¹)</th>
<th>UV flux (ph cm⁻² s⁻¹)</th>
<th>Time (min)</th>
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<td>15</td>
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<td>3 × 10¹³</td>
<td>3.5 × 10¹⁵</td>
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<td>0.5</td>
<td>3 × 10¹³</td>
<td>3.5 × 10¹⁵</td>
<td>...</td>
<td>120</td>
</tr>
<tr>
<td>NO</td>
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<td>25</td>
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<td>7 × 10¹²</td>
<td>2.5 × 10¹⁴</td>
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<td>120</td>
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<td>2.5</td>
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<tr>
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<td>15</td>
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<td>1.5 × 10¹⁴</td>
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aCo-deposition is an experiment in which NO containing ices are deposited during H/D flux exposure.
bPre-deposition is an experiment in which NO containing ices are first deposited and subsequently exposed to H/D atom or UV fluxes.
cThis is a set of experiments in which the spectra of different nitrogen oxides were obtained and then monitored for stability upon H-atoms flux exposure.
dNO₂ gas contained considerable NO admixture; formation of N₂O₃ (ON–NO₂) along with N₂O₄ (O₂N–NO₂) is observed upon deposition.
eMixture of N₂O₃ (ON–NO₂) and cis(NO₂) is formed upon deposition.

References:
4, 9 a compact ice structure prevents H/D atoms from penetrating deep into the bulk of the ice. As a consequence, in a pre-deposition experiment only the first few monolayers are actively involved in surface reactions, the final yield of H/D atom addition products is low and no intermediate products are detected. Figures 1(b) and 1(c) show NO/H-atom co-deposition experiments at 15 K. The main advantage of co-deposition experiments is the ability to change the NO/H(D) ratio in order to: (1) trap non-fully hydrogenated products inside the growing ice by using a high NO flow with a low H(D)-atom flux (Figs. 1(c) and 1(e)) or (2) deposit a fully hydrogenated thick ice by using a low NO flow with a high H(D)-atom flux (Figs. 1(b) and 1(d)). Growing a thick hydrogenated ice during co-deposition experiments increases the column densities of the reaction products and allows their unambiguous identification in situ and at low temperatures by means of RAIRS. Although QMS is more sensitive than RAIRS, it allows for the detection of species only above their desorption temperatures. For the species studied here, this is higher than 15 K, i.e., > 75 K for N₂O and > 140 K for NH₃OH.

Several positive absorption features, e.g., at 1608 cm⁻¹ (NH₂ scissoring), 1514 cm⁻¹ (HON bending), 1203 cm⁻¹ (NH2 bending), and at low temperatures by means of RAIRS.
TABLE II. List of assigned nitrogen-bearing species.

<table>
<thead>
<tr>
<th>Mode</th>
<th>Regular species</th>
<th>Frequency (cm(^{-1}))</th>
<th>Deuterated species</th>
<th>Frequency (cm(^{-1}))</th>
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<tr>
<td></td>
<td>NO (monomer)</td>
<td>1875</td>
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<tr>
<td>(\nu_6)</td>
<td>NH(_2)OH</td>
<td>889</td>
<td>ND</td>
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<td>919</td>
<td>ND(_2)OD</td>
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<tr>
<td>(\nu_4)</td>
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<td>ND(_2)OD</td>
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<td></td>
<td>X–NO</td>
<td>1829</td>
<td>X–NO</td>
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<tr>
<td>(\nu_1)</td>
<td>N(_2)O(_3)</td>
<td>1851</td>
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<td>(\nu_4)</td>
<td>N(_2)O(_3)</td>
<td>782</td>
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<td>28, 35</td>
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\(^a\)Spectral features used for integration and relative quantification.

(NH\(_2\) wagging), and 919 cm\(^{-1}\) (ON stretching),\(^{31}\) presented in Fig. 1(b) clearly show that the main hydrogenation product of a pure NO ice at low temperature is solid hydroxylamine. The NH\(_2\)OH formed during the co-deposition experiment shown in Fig. 1(b) is bulk hydroxylamine, bound with four different types of hydrogen bonds, such as OH···H, OH···N, NH···O, NH···N, and with different strengths. The hydroxylamine OH and NH stretching modes are in the range from 2600 to 3400 cm\(^{-1}\) and overlap with each other. There are two groups of broad bands in this region, at 3317, 3261, and 3194 cm\(^{-1}\) and at 2899 and 2716 cm\(^{-1}\). The first group can be assigned to the symmetric and asymmetric N–H stretches, and the O–H···O stretch, while the second group is due to a hydrogen bonded OH···N stretch.\(^{31, 39}\) The formation of hydroxylamine in our pre-deposition experiments is confirmed by comparing all the hydroxylamine peaks presented in Fig. 1(b) with those shown in Fig. 1(a).

Figure 1(c) presents several features due to the (NO)\(_2\) dimer, N\(_2\)O, and isolated NH\(_2\)OH, and reveals several new features: two narrow bands at 1561 and 1507 cm\(^{-1}\) that can be assigned to HNO, and another two at 1829 cm\(^{-1}\) and 1680 cm\(^{-1}\) with an unclear origin. The feature at 1829 cm\(^{-1}\) overlaps with the NO stretching mode which may indicate that this species contains weakly bound NO molecules. Hereafter, we will refer to this unknown component as X–NO. A broad peak at 1680 cm\(^{-1}\) is present in almost every spectrum and demonstrates similar behaviour under various experimental conditions, thus interpretation of this feature is uncertain. To verify our assignments, we repeated the same co-deposition experiments discussed above using D atoms instead of H atoms (Figs. 1(d) and 1(e)).

TPD experiments further constrain the formation of NH\(_2\)OH and N\(_2\)O. Under our experimental conditions, the desorption of NH\(_2\)OH starts at 140 K and peaks at 175 K (TPD rate 1 K/min). The NH\(_2\)OH desorption peak is reported at 188 K in Ref. 22. There a desorption energy of 54.2 kJ mol\(^{-1}\) is calculated for pure NH\(_2\)OH from a surface of amorphous silicate. The difference in desorption temperature between the present work and Ref. 22 is not surprising, as we study ices in the multilayer regime. Here, desorption of NH\(_2\)OH takes place from the surface of the ice itself. Therefore, under this regime molecules from a layer of NH\(_2\)OH are weakly bound to other molecules (like NH\(_2\)OH, or NH\(_2\)OH mixed with CO and/or H\(_2\)O, see Secs. III B and III C). This leads to a low temperature desorption peak. In Ref. 22, at most one monolayer of NH\(_2\)OH is deposited onto an amorphous silicate surface. In that case NH\(_2\)OH molecules are bound to the silicate surface and not to another layer of ice. The resulting binding energy is higher, leading to a higher temperature desorption peak.

B. Hydrogenation of NO in non-polar CO surroundings

We performed a series of experiments aimed at studying the effect of NO dilution in CO and N\(_2\) (non-polar) matrices on the hydrogenation pathways. These experiments simulate the formation of hydroxylamine under dense and cold...
interstellar cloud conditions, when gas-phase CO freezes out. Figure 2 shows several examples of RAIR spectra of NO:CO co-deposition experiments with different ice mixtures (1:1, 1:6) and H/D-atom ratios.

All bands observed in a pure NO co-deposition experiment (Figs. 1(b) and 1(c)) are present here, and in addition, new features appear. In Figure 2(a), the NO monomer feature peaks at 1875 cm$^{-1}$, while the trans-(NO)$_2$ absorbance band is centred around 1745 cm$^{-1}$. The presence of the NO monomer in the ice mixture indicates that the NO mobility is limited in a CO lattice at 15 K. The remaining absorption features belong to the products of H/D-atom addition to CO, such as H$_2$CO (1720, 1499 cm$^{-1}$) and CH$_3$OH (1060 cm$^{-1}$). These reactions have been examined extensively in previous studies. In the case of D-atom exposure, traces of D$_2$CO are detected at 1674 and 2092 cm$^{-1}$ but deuterated methanol is not found. It should be noted that formaldehyde is detected only in the experiments where all NO is converted into its final hydrogenation products (Figs. 2(b) and 2(d)). In the experiments where nitric oxide is still present in the ice, no products of H/D-atom addition to CO ice are found. This observation will be addressed in Secs. III E–III G. Furthermore, no absorption features that can be assigned to N–C bond bearing species, like NCO radicals, HNCO or NH$_2$CHO, are found by RAIRS or QMS. Unstable intermediates like HCO and HNOH are not detected either. Note that due to the abundance of the $^{13}$CO isotope in the CO gas cylinder (1.1%), a clear $^{13}$CO absorbance feature at 2092 cm$^{-1}$ is also observed.

Unlike CO ice, N$_2$ ice is inert to H-atom addition. The hydrogenation of nitric oxide in a nitrogen matrix with a ratio of NO:N$_2$ = 1:5 shows qualitatively that the final products from NO + H are the same as in a pure NO hydrogenation co-deposition experiment (Fig. 1(c)) with the difference that this time a small feature around 1875 cm$^{-1}$ can be assigned to the NO monomer (as in Fig. 2). Here, N$_2$O seems to be formed more efficiently compared to the NO hydrogenation experiment in a CO matrix.

C. Polar H$_2$O surroundings

Water is the most abundant species detected in interstellar ices towards dark clouds, low-mass and high-mass YSOs. It is generally accepted that water ice is formed mainly during the early stages of a translucent and quiescent cloud. Gas-grain chemical models predict that the NO abundance peaks later during CO freeze out. Solid hydroxylamine can be formed during the same early stages in star forming regions from accreted NO or from NO formed on the surface of grains. Therefore, we studied the hydrogenation of NO in a water (polar) matrix (NO:H$_2$O = 1:6). H$_2$O molecules are inert to H-atom addition, but provide hydrogen bonds that may interact with NO molecules as well as possible reaction products upon NO hydrogenation, i.e., HNO and NH$_2$OH. Here, a NO:H$_2$O (1:6) ratio is used for comparison with the aforementioned non-polar mixtures.

Water ice deposited at 15 K forms an amorphous bulk with several very broad bands: the OH stretching mode between 3000 and 3700 cm$^{-1}$, the HOH bending mode in the 1400–1700 cm$^{-1}$ range and the libration mode between 700 and 1000 cm$^{-1}$. All these bands are present in our spectra (see
Fig. 3). As opposed to NO:CO ices, no nitric oxide monomer and trans-(NO)$_2$ are detected in the ice, which can be explained by a different mobility of NO in water ice compared to NO in a CO matrix. Although most of the NH$_2$OH absorption features overlap with the water OH stretching and HOH bending modes, all of them can be found in Fig. 3(b), where the H-atom flux is high enough to efficiently convert all NO to hydroxylamine. Figure 3(a) shows a RAIR spectrum after co-deposition of NO:H$_2$O with a low H-atom flux. Here, unlike Fig. 3(b), HNO, X-NO, and N$_2$O features can be observed in the ice.

D. UV photolysis of NO containing ices

Several pre-deposition experiments were performed with CRYOPAD to study the UV photon induced chemistry in pure NO ice and in ice mixtures containing nitric oxide in water or carbon monoxide. In addition, tertiary ice mixtures NO:CO:H$_2$O with ratios 1:1:1 and 1:6:6 were exposed to UV photons.

As discussed in Sec. III A, pure NO converts efficiently into cis-(NO)$_2$ ice upon deposition (Fig. 4(b)). Several positive bands appear after UV processing of cis-(NO)$_2$ ice: the features at 2235 and 1286 cm$^{-1}$ already observed in the NO hydrogenation experiments are assigned to N$_2$O. New bands with maxima at 1602, 1299, 779 cm$^{-1}$ and, although not well resolved, 1850 cm$^{-1}$ appear only in the UV photolysis experiments. These can be assigned to N$_2$O$_3$ (see Fig. 4(a)).$^{28,35,40}$ The dilution of NO molecules in a CO matrix does not seem to influence this outcome substantially. As for Fig. 4(a), all the absorption features assigned to cis-(NO)$_2$, N$_2$O, and N$_2$O$_3$ are present in the RAIIR spectra after exposure of NO:CO = 1:6 ice to UV photons (see Fig. 4(c)). In addition, the NO monomer is clearly detected because of the low mobility of NO in a CO lattice before and after UV photolysis of the ice sample. Moreover, the NO monomer abundance increases upon UV irradiation.

Two more carbon bearing species, apart from $^{12}$CO and $^{13}$CO are detected in a NO:CO ice upon UV photolysis: carbon dioxide and its $^{13}$C isotope analogue. Since the solid carbon dioxide IR absorption feature overlaps with atmospheric gas-phase CO$_2$ present outside the UHV chamber and along the line of the FTIR beam, the formation of carbon dioxide in the ice upon UV photolysis is constrained by QMS data.

Figures 4(e) and 4(f) show the resulting spectra upon UV photolysis of a NO:H$_2$O (1:6) ice mixture. Also in this case, two species are present in the ice upon deposition: cis-(NO)$_2$ and H$_2$O. UV processing of the sample induces formation of several products: N$_2$O and monomeric NO, hydrogen peroxide (H$_2$O$_2$), and traces of HNO (right shoulder at 1565 cm$^{-1}$). Absorption features at 1607 and 1303 cm$^{-1}$ may be due to N$_2$O$_3$ or NO$_2$. The absence of an absorption feature at 1850 cm$^{-1}$ makes the NO$_2$ assignment more likely, but due to the low final yield this is not conclusive.

The UV processing of the tertiary NO:CO:H$_2$O ice mixtures reveals the formation of the same species formed in the aforementioned NO:CO and NO:H$_2$O ice photolysis experiments. Here, the same features are assigned to H$_2$CO with peaks at 1715 and 1500 cm$^{-1}$ and tentatively assigned to the NO stretching mode of HNO$_2$ peaking at 1640 cm$^{-1}$. It should be noted that the final yield of solid CO$_2$ is higher than in the case of UV irradiated NO:CO experiments. The addition of water ice to the NO:CO ice mixture increases the carbon dioxide formation. This can be explained by the presence of free OH radicals in the ice after UV irradiation that can react with CO to form CO$_2$ (more in Sec. III G 2).

As can be concluded from Fig. 4, NCO radicals, NC$^-$ and NCO$^-$ anions as well as their hydrogenated analogues...
HNC and HNCO (Refs. 41–44) are not present in any of our RAIR spectra. Another main result here is that NH$_2$OH is not formed upon UV photolysis of water rich ices. The straightforward conclusion is that surface H-atom addition reactions are needed to provide a pathway for the formation of NH$_2$OH in simple interstellar ice analogues. We will address these points in Sec. III G.

### E. Temperature dependence

Temperature effects on reaction mechanisms can be rather complex and reflect the cumulative outcome of different temperature dependent processes. First of all, a higher ice temperature means more thermal energy that can help to overcome activation barriers. This has a direct effect on the final yields of the formed species and, therefore, on the formation trends. Second, a higher temperature can have an effect on the structure of the ice leading to molecular reorganization. For instance, it has been shown that different deposition temperatures result in different ice structures, e.g., amorphous water ice versus crystalline water ice.\(^{45, 46}\) This affects the penetration depth of H atoms in the ice and more generally diffusion processes. Then, the hopping and swapping rates of molecules in the ice increase with temperature, while the residence lifetime of a species on the ice surface decreases.

The temperature dependence of selected surface formation reaction pathways is studied here by repeating the same experiments for two different and astronomically relevant temperatures, i.e., 15 and 25 K. Since the thermal desorption temperatures of CO and N$_2$ are around this value, higher temperatures have not been studied. The top panel of Fig. 5 shows the formation trends of HNO and NH$_2$OH at 15 and 25 K as a function of the H-atom fluence for a co-deposition experiment of pure NO and H atoms. Here, the integrated absorption of the formed species presents a linear trend because of the constant co-deposition of the parent species. The bottom left-side panel of Fig. 5 shows NH$_2$OH, HNO, and N$_2$O final yields from the aforementioned pure NO + H experiments, while the bottom right-side panel shows NH$_2$OH, HNO, and N$_2$O final yields for two NO:CO:H co-deposition experiments at 15 and 25 K. All these experiments present lower final product yields at higher temperatures. Moreover, as shown in both bottom panels, the ratios between HNO and NH$_2$OH final yields at 15 and 25 K are roughly constant and are 1.5 and 3.5 for the pure NO + H and the NO:CO + H experiments, respectively. This indicates that the formation of HNO and NH$_2$OH are linked and are most likely temperature independent, i.e., reactions take place without an activation barrier. The lower final yields at higher temperature can be explained by the lifetime of H atoms on the ice surface that, for instance, is $10^3$ times shorter at 25 K than at 15 K on water ice.\(^{47}\) In this way, H atoms have less time to react with NO molecules leading to a decrease in the hydrogenation final yields. All final yields in the bottom right panel are lower than their counterparts in the bottom left panel. This can be explained by a lower abundance of NO molecules in the ice and the presence of solid CO. For the pure NO + H experiments, the final yield of solid N$_2$O is the same at different temperatures within the experimental uncertainties, while it decreases for higher temperatures in the NO:CO + H experiments. This indicates that the formation of N$_2$O requires a temperature dependent reaction scheme and a more complex ice composition.

### F. Matrix effects

The dilution of NO in polar (water-rich ice) and non-polar (water-poor ice) matrices has a number of important effects on the hydrogenation pathways: (1) NO hydrogenation products can react with components of the mixture other than hydrogen atoms leading to a more complex reaction scheme; (2) the other molecules present in the deposited matrix (like CO) can participate in H-atom addition reactions, effectively influencing reaction rates through additional H-atom consumption. Their reaction products may interact with NO and NO hydrogenation reaction products; (3) other species in the ice can trap NO molecules in the bulk of the ice hiding them from the impinging H atoms; (4) H-bonding can improve coupling and heat dissipation through the ice, promoting the stabilization of intermediate products and changing reaction branching ratios.

In order to investigate non-polar ices, we have compared the results from the hydrogenation of pure NO ice with those from the hydrogenation of NO:CO ice mixtures with ratios of 1:1 and 1:6. The choice of these ratios is not random: if NO hydrogenation products interact with those from CO hydrogenation, then the 1:1 ratio should provide the highest yield of the final products; on the other hand, for a NO:CO = 1:6 ratio a single NO molecule is surrounded on average by an octahedron of six CO molecules (the NO dimer is surrounded by 12 CO molecules). Therefore, this is a good approach to see whether NO and CO interact with each other upon hydrogenation of the ice. Moreover, since CO is highly abundant in...
FIG. 6. RAIR integrated intensities of hydroxylamine (left-side panels), nitroxyl (centre panels), and nitrous oxide (right-side panels) versus H-atom fluence for several co-deposition experiments using high H-atom flux (top panels) and low H-atom flux (bottom panels). Every panel shows three curves: one for H atoms co-deposited with pure NO (black squares), another for H atoms co-deposited with NO:CO = 1:1 mixture (red circles), and one for co-deposition of H atoms co-deposited with NO:CO = 1:6 mixture (blue triangles). A missing curve indicates that the product is not detected during the experiment.

quiescent dark clouds compared to NO, a 1:6 ratio is astrochemically more relevant. We furthermore investigated the two NO:CO 1:1 and 1:6 mixtures for two different H-atom fluxes of $7 \times 10^{12}$ and $3 \times 10^{13}$ atom cm$^{-2}$ s$^{-1}$. During these co-deposition experiments, the ratio between H atoms and the sum of NO and CO molecules is kept constant. Figure 6 plots the integrated intensities of NO hydrogenation products as a function of the H-atom fluence for pure NO and NO:CO mixtures with ratios 1:1 and 1:6 co-deposited with H atoms.

The upper panels of Fig. 6 show the integrated intensities of NH$_2$OH, HNO, and N$_2$O versus the H-atom fluence for three co-deposition experiments (NO/H = 1/10, NO:CO/H = 1:20, 1:6/70) where the H-atom flux is high in order to efficiently convert nitric and carbon oxides into their hydrogenation products. The abundance of NH$_2$OH in the ice is proportional to the abundance of NO in the mixture, while the HNO abundance is inversely proportional to that of NO. The N$_2$O ice abundance is strongly affected by the presence of CO in the ice, however, it does not seem to depend on the mixture ratios used here. In addition, and as mentioned before, NO absorption features (from both the dimer and monomer) are not present in the RAIR spectra of pure NO + H and NO:CO = 1:1 + H (see Figs. 1(b) and 2(b)) but only in the NO:CO = 1:6 hydrogenation experiment (Fig. 2(a)). The presence of NO in the ice after hydrogenation can be explained by the trapping effect of NO inside a CO lattice and the consequent blocking effect of H atoms in a CO ice. Consequently, the gradual decrease of the hydroxylamine yield with increasing CO content in the ice can be explained by a decrease in the absolute amount of NO available for hydrogenation reactions. This conclusion is further constrained by the nitroxyl final yield, which is inversely proportional to the hydroxylamine yield: HNO is only observed in the ice when some NO remains unconverted into NH$_2$OH. This confirms that the formation paths of nitroxyl and hydroxylamine are linked.

The lower panels of Fig. 6 show the RAIR integrated absorbances of NH$_2$OH, HNO, and N$_2$O versus the H-atom fluence for three co-deposition experiments (NO/H = 1/0.5, NO:CO/H = 1:1/1, 1:6/3.5) but now the H-atom flux is low in order to detect the not fully hydrogenated species in the ice. All plots show very similar abundances of N$_2$O and HNO and slightly higher yields for NH$_2$OH in a pure NO hydrogenation experiment compared to the hydrogenation of NO:CO mixtures. The products of H-atom addition to CO molecules (H$_2$CO, CH$_3$OH) are not detected in any of these experiments. This leads to an important conclusion: the hydrogenation of NO ice is more efficient than the hydrogenation of CO ice and is most likely barrierless as concluded here and in previous work. It was shown before that the intermediate CO + H and H$_2$CO + H reaction steps have an activation barrier, while HCO + H and H$_3$CO + H take place without an activation barrier. Thus, if the subsequent hydrogenation of nitric oxide (NO + H, HNO + H, H$_2$NO + H) proceeds with no activation barrier, then the NH$_2$OH final yield depends only on the amount of NO molecules and H atoms available for reaction. This is fully consistent with the observation that all infrared NH$_2$OH, HNO, and N$_2$O integrated band intensities shown in the lower panels of Fig. 6 present a similar final yield. Minor differences among the formation trends may be due to the presence of CO molecules.
FIG. 7. RAIR integrated intensities of hydroxylamine (left-side panel), nitroxyl (centre panel), and nitrous oxide (right-side panel) formed upon H-atom exposure versus the ice thickness in a co-deposition experiment of NO:CO (1:6) + H and NO:H2O (1:6) + H. Each panel shows four curves: low H-atom flux co-deposited with NO:H2O = 1:6 mixture (full black circles); low H-atom flux co-deposited with NO:CO = 1:6 mixture (empty black circles); high H-atom flux co-deposited with NO:H2O = 1:6 mixture (full red diamonds); low H-atom flux co-deposited with NO:CO = 1:6 mixture (empty red diamonds).

The dipole moments of CO, NO, and the cis-(NO)2 are all small and comparable,48–50 and the ice structures are also quite similar.36, 38 This explains why dilution of NO in a CO matrix does not substantially influence the reaction network and, consequently, the formation trends of the final products. In contrast, water molecules have a strong dipole moment and form strong hydrogen bonds in the ice with other species, such as NH2OH, NH2O·, and NHOH· radicals, and H2O itself.

We performed several experiments co-depositing NO:H2O binary mixtures and H atoms using similar experimental conditions to those adopted for the non-polar ice investigation. The results are summarized in Fig. 7. In this case, we compared results for a NO:H2O = 1:6 ratio co-deposition experiment with low and high H-atom fluxes. For a high H-atom flux, the water matrix significantly increases the hydroxylamine final yield, while it decreases both nitroxyl and nitrous oxide final yields. For low H-atom flux only hydroxylamine exhibits a higher yield in a water lattice compared to a similar experiment with NO:CO = 1:6.

The NHOH and NH2O intermediates, as well as final product NH2OH, form strong hydrogen bonds with a water lattice. HNO also forms this kind of bond, but less efficiently. Such hydrogen bonds help HNO and NH2OH molecules to dissipate excess energy quickly upon formation and make the hydrogenation of NO and HNO kinetically favourable compared to the case of NO + H in a CO lattice in which these kinds of bonds are not formed. It should be noted that a water matrix significantly decreases the N2O final yield for high H-atom flux, but increases it for low flux. Moreover, hydrogenation of NO in a N2 matrix (NO:N2 = 1:5) for high H-atom flux gives a N2O yield that is three times higher than that obtained in a NO:CO hydrogenation experiment. The interpretation of this result is not straightforward and the influence of different lattice types on the N2O formation route is therefore uncertain.

G. Possible reaction pathways

1. NO hydrogenation network

Figure 8 shows the NH2OH formation reaction scheme as investigated in this work. The solid lines indicate reaction pathways that take place without an activation barrier or with a very small activation barrier and are therefore efficient at cryogenic temperatures, e.g., in dense cold quiescent clouds. The dashed lines represent the reactions that proceed with an activation barrier.

In Fig. 8, three subsequent H-atom addition reactions to form NH2OH from a single NO molecule are highlighted in red double-bordered boxes

\[ \text{NO} + \text{H} \rightarrow \text{HNO}, \]  
\[ \text{HNO} + \text{H} \rightarrow \text{H}_2\text{NO}, \]  
\[ \text{H}_2\text{NO} + \text{H} \rightarrow \text{H}_2\text{NOH}. \]

These three reactions represent the simplest and most efficient formation route of hydroxylamine. In reactions (1a) and (1b), the H atoms are subsequently added to the N atom of the NO molecule, while the last H atom in reaction (1c) is added to the oxygen. The reactions

\[ \text{NO} + \text{H} \rightarrow \text{NOH}, \]

FIG. 8. A schematic representation of the solid-state NO + H reaction network, summarizing the conclusions from this work and Ref. 22. Solid squares indicate species that are detected in the ice while dashed squares indicate intermediate products that are not detected but may be formed in the ice. Solid arrows represent efficient and/or barrierless reactions. Dashed arrows represent reaction pathways with an activation barrier. Only atom and radical addition reactions are presented here (e.g., hydrogen abstraction reactions are omitted for simplicity).
where the addition of H atoms takes place on the oxygen atom side of the NO molecule, are expected to be less efficient. Matrix isolation experiments\textsuperscript{44} and \textit{ab initio} calculations\textsuperscript{52} indeed indicate that these reactions have high activation barriers (>4000 K and >10000 K, respectively)\textsuperscript{52} and are, therefore, less efficient compared to the barrierless reactions (1a) and (1b). However, for the reaction

\begin{equation}
\text{HNO} + \text{H} \rightarrow \text{HNOH}, \tag{2b}
\end{equation}

the presence of a third body can efficiently change the branching ratio between H-atom addition and H-atom subtraction reactions, leading, for instance, to the efficient surface formation of NH\(_2\)OH ice. As in reaction (3a), the hydrogen abstraction reaction

\begin{equation}
\text{H}_2\text{NO} + \text{H} \rightarrow \text{HNO} + \text{H}_2 \tag{3a}
\end{equation}

is also barrierless\textsuperscript{51,52} and should compete with the hydrogen addition reaction (1c). However, as discussed above, the formation of hydroxylamine has a higher effective branching ratio than reaction (3b) in the solid state.

The hydrogenation of NO monomers through reactions (1a)-(1c) takes place only in NO:CO and NO:H\(_2\)O hydrogenation experiments where the CO or H\(_2\)O lattice can keep a fraction of the NO monomers isolated and can prevent their dimerization through reactions (4a) and (4b)

\begin{equation}
\text{NO} + \text{NO} \rightarrow \text{cis-}(\text{NO})_2, \tag{4a}
\end{equation}

\begin{equation}
\text{NO} + \text{NO} \rightarrow \text{trans-}(\text{NO})_2. \tag{4b}
\end{equation}

Formation of \textit{cis-}\(\text{NO}_2\) is the preferred reaction pathway and is confirmed by NO matrix isolation experiments.\textsuperscript{28,35} Under our experimental conditions \textit{trans-}\(\text{NO}_2\) is observed only in a NO:CO:H = 1:6:70 co-deposition experiment where it is trapped by the CO lattice and the products of NO hydrogenation.

Due to a lack of gas-phase kinetic data, the hydrogenation reaction scheme for \(\text{NO}_2\) has so far been unclear. Here, we experimentally show several reaction routes involving the hydrogenation of \(\text{NO}_2\). First, the nitric oxide dimer reacts with a H atom to form HNO and NO,

\begin{equation}
\text{(NO)}_2 + \text{H} \rightarrow \text{HNO} + \text{NO}, \tag{5a}
\end{equation}

or produces a HNO·NO associated complex. The existence of this complex has already been suggested by several groups,\textsuperscript{53-56}

\begin{equation}
\text{(NO)}_2 + \text{H} \rightarrow \text{HNO} \cdot \text{NO}. \tag{5b}
\end{equation}

The addition of another H atom to a nitrogen atom of this associated complex leads to the formation of H\(_2\)NO, HNO, and NO

\begin{align*}
\text{HNO} \cdot \text{NO} + \text{H} & \rightarrow \text{H}_2\text{NO} + \text{NO}, \tag{6a} \\
\text{HNO} \cdot \text{NO} + \text{H} & \rightarrow \text{HNO} + \text{HNO}. \tag{6b}
\end{align*}

All these reactions lead to the formation of hydroxylamine precursors as shown by reactions (1a)-(1c). Our experiments confirm that hydroxylamine will eventually be formed upon hydrogenation of these species.

There is another final product of NO hydrogenation reactions which is always observed in our experiments: N\(_2\)O ice. The reaction pathway to form N\(_2\)O is not yet clear, since all known possible mechanisms involve a high activation barrier. A possible mechanism to form N\(_2\)O is through HNO dimerization and the subsequent dissociation of the formed hyponitrous acid,\textsuperscript{51,53,54}

\begin{equation}
\text{HNO} + \text{HNO} \rightarrow \text{H}_2\text{N}_2\text{O}_2 (\text{OHN} = \text{NHO}), \tag{7}
\end{equation}

\begin{equation}
\text{H}_2\text{N}_2\text{O}_2 (\text{OHN} = \text{NHO}) \rightarrow \text{isomerisation} \rightarrow \text{N}_2\text{O} + \text{H}_2\text{O}. \tag{8}
\end{equation}

This mechanism is rather unlikely to occur at cryogenic temperatures, because reaction (8) involves at least one H-atom migration from a nitrogen atom to an oxygen atom of a OHN = NHO molecule leading to the formation of the intermediate HON(NO)H or HON = NOH which will then dissociate to form nitrous oxide and water.\textsuperscript{51,53} Both dissociation and H-atom migration involve activation barriers as well as the HNO dimerization itself (reaction (7)). However, several of these activation barriers can be overcome by the following mechanism. First, a H atom is added to a N atom of the HNO·NO associated complex. This reaction is exothermic and takes place without an activation barrier. The excess energy of this reaction can then help to overcome the H\(_2\)N\(_2\)O\(_2\) dissociation barrier leading to the formation of nitrous acid by a mechanism similar to reactions (7) and (8)

\begin{equation}
\text{HNO} \cdot \text{NO} + \text{H} \rightarrow \text{H}_2\text{N}_2\text{O}_2 ("\text{hot")}
\end{equation}

\begin{equation}
\rightarrow \text{isomerisation} \rightarrow \text{N}_2\text{O} + \text{H}_2\text{O}. \tag{9}
\end{equation}

The weak point of both suggested mechanisms is the high activation barrier in the last step of reactions (8) and (9). Because of this barrier H\(_2\)N\(_2\)O\(_2\) (OHN = NHO) should be detectable in our spectra under our experimental conditions together with N\(_2\)O. Moreover, the infrared spectrum of H\(_2\)N\(_2\)O\(_2\) and its deuterated analogue have strong absorption features, i.e., two close bands around 1000 cm\(^{-1}\) which are not present in our RAIR spectra (see Figs. 1(c), 1(e), 2(a), and 2(c)).\textsuperscript{57}
The absence of H$_2$N$_2$O$_2$ absorption features indicates that another reaction pathway is likely responsible for the N$_2$O formation.

One of these possible formation mechanisms for N$_2$O was suggested earlier in Refs. 58 and 59, where reduction of NO with alkaline hydroxylamine was studied:

\[
\text{HNOH} + \text{NO} \rightarrow \text{HON(NO)H}, \quad (10)
\]

\[
\text{HON(NO)H} \rightarrow \text{N}_2\text{O} + \text{H}_2\text{O}. \quad (11)
\]

Reaction (10) is exothermic$^{51}$ and, as other radical-radical reactions, should proceed without an activation barrier. The product of reaction (10) may be formed with enough internal energy to overcome its dissociation barrier and to produce nitrous oxide and water (reaction (11)). However, the formation of HNOH (reaction (2b)) is thermodynamically unfavourable.$^{51}$ Therefore, another formation mechanism is proposed here, i.e., the direct addition of a H atom to the oxygen atom of a (NO)$_2$ dimer followed by the dissociation of the unstable radical:

\[
\text{ONNO} + \text{H} \rightarrow (\text{ONNOH}) \rightarrow \text{N}_2\text{O} + \text{OH}. \quad (12)
\]

This reaction pathway is exothermic, and according to BAC-MP4 calculations$^{51}$ the second step is barrierless (or at least has a very small barrier). The first step of this reaction is similar to reaction (2a) and, therefore, presents an activation barrier (dashed arrow in Fig. 8). Although our experimental data are not conclusive on the efficiency of the N$_2$O formation mechanisms, reaction (12) is likely the most efficient pathway because of its simplicity and because we do not detect the intermediates formed according to reactions (7)–(11).

In Ref. 22, two reaction pathways are suggested for the formation of N$_2$O ice in a low surface coverage regime (<1 monolayer) on silicate surfaces. The first one is reaction (12) as discussed above. A different formation route is:

\[
\text{HNO} + \text{NO} \rightarrow \text{N}_2\text{O} + \text{OH}. \quad (13a)
\]

BAC-MP4 calculations$^{51}$ show that this formation mechanism involves at a certain stage the migration of one H atom from a nitrogen to an oxygen atom:

\[
\text{HNO} + \text{NO} \rightarrow \text{HN(NO)} \rightarrow \text{ONNOH} \rightarrow \text{N}_2\text{O} + \text{OH}. \quad (13b)
\]

This migration process has a very high activation barrier (122 kJ mol$^{-1}$) and is endothermic.$^{51}$ A possible way to overcome this activation barrier under our experimental conditions is through the participation of a “hot” HNO molecule formed from reaction (1a) before the HNO dissipates its excess energy (203 kJ mol$^{-1}$) (Ref. 51) into the bulk of the ice. However, this reaction is unlikely to occur in a high surface coverage regime, when HNO is always surrounded by other species. Moreover, reaction (13) requires at the same time the presence of NO monomer, which is detected in a high surface coverage regime only if NO is diluted into a CO matrix (see Fig. 2(a)). Therefore, reaction (13) is unlikely to be efficient under our experimental conditions.

The unidentified absorption feature in our IR spectra, here named X–NO, appears at 1829 cm$^{-1}$. Further experimental constraints are needed to assign this feature to the HNO· NO complex or to the HON(NO)H intermediate.

The produced OH radicals from reaction (12) can react with H atoms and NO molecules to form:

\[
\text{OH} + \text{H} \rightarrow \text{H}_2\text{O}, \quad (14a)
\]

\[
\text{OH} + \text{NO} \rightarrow \text{HNO}_2. \quad (14b)
\]

Recently another solid-state reaction pathway involving the OH radical was extensively investigated at cryogenic temperatures in Refs. 60 and 61:

\[
\text{OH} + \text{H}_2 \rightarrow \text{H}_2\text{O} + \text{H}. \quad (14c)
\]

The identification of H$_2$O and HNO$_2$ is not trivial due to the low final yield of these species and the overlap of their absorption features with other molecules present in the ice. Newly formed H$_2$O molecules are reported in the gas phase upon desorption from the ice in Ref. 22. The amount of water ice formed at the end of the experiments described there is estimated to be a fraction of a monolayer (~0.5 ML). Moreover, all the suggested formation routes of H$_2$O ice have N$_2$O as an accompanying product. Since N$_2$O is present only in low abundances in our experiments, the formation of H$_2$O in the present study is expected to be inefficient.

As previously mentioned H$_2$CO is not formed during low H-atom flux experiments, but it is present in the ice only if all NO is converted to its hydrogenation products. Both gas-phase experimental work$^{62}$ and calculations$^{51}$ show that the following reaction:

\[
\text{HCO} + \text{NO} \rightarrow \text{HNO} + \text{CO} \quad (15)
\]

is exothermic and likely barrierless. If reaction (15) takes place under our experimental conditions, the presence of NO in the ice can efficiently reduce the final H$_2$CO yield in favour of the NH$_2$OH formation. This is consistent with our experimental results. However, there is no direct evidence for this reaction pathway in our experiments. In particular, the low efficiency of H$_2$CO formation may also be explained by a short residence time of the H atoms on the ice surface combined with the competition between the CO + H reaction, which has a barrier, and the barrierless NO + H reaction (1a).

2. **UV processing of the ice**

As opposed to the neutral-neutral non-energetic surface reactions discussed in Sec. III G 1, UV processing of ices provides an energetic input that leads to the dissociation of molecular bonds and the formation of “hot” fragments, i.e., radicals, atoms, and ions. These “hot” fragments can in turn react with the surrounding molecules and easily overcome activation barriers, desorb from the ice surface or diffuse into the bulk of the ice and kick other molecules out.$^{12,26,27}$ Here, we focus on the UV photolysis of pure solid NO and NO mixed with CO and H$_2$O ices. Although a detailed quantitative analysis of the UV-induced reaction network is challenging without knowing photodissociation and photoionisation cross sections for all the molecules present and/or formed in the ice...
upon UV photolysis, we can still draw important conclusions on the efficiency of several reaction pathways.

UV processing of pure NO ice gives two final products: N₂O and N₂O₃. These species can be formed through the following reaction:

\[ 2(\text{NO})_2 + h\nu \rightarrow \text{N}_2\text{O} + \text{N}_2\text{O}_3. \] (16)

First, a N-O bond from a cis-(NO)₂ is broken upon UV photolysis (reaction (17a)). Then, the free oxygen atom recombines with a nearby nitrogen atom of another (NO)₂ dimer to form dinitrogen trioxide (reaction (17b)):

\[ (\text{NO})_2 + h\nu \rightarrow \text{N}_2\text{O} + \cdot\text{O}, \] (17a)

\[ (\text{NO})_2 + \cdot\text{O} \rightarrow \text{N}_2\text{O}_3. \] (17b)

This result is consistent with Ref. 63 where the photolysis of cis-(NO)₂ is studied at 13 K in an Ar matrix. The formation of nitrous oxide via direct photodissociation of nitric oxide and the consequent N-atom addition reaction to a NO dimer

\[ (\text{NO})_2 + h\nu \rightarrow \text{NO} + \cdot\text{N} + \cdot\text{O}, \] (18a)

\[ (\text{NO})_2 + \text{N} \rightarrow \cdot\text{NCO}. \] (19)

UV processing of NO:CO ice mixtures yields two additional products: NO monomers and CO₂. The formation of NO monomers can be explained by the reaction

\[ (\text{NO})_2 + h\nu \rightarrow \text{NO} + \text{NO}^*. \] (20)

and the subsequent isolation of “hot” monomers in a CO lattice.

The other product is CO₂ ice that can be formed through “hot” oxygen atom addition to carbon monoxide

\[ \text{CO} + \cdot\text{O} \rightarrow \text{CO}_2. \] (21)

This reaction pathway is discussed in several experimental studies. According to Ref. 64, reaction (21) takes place even at cryogenic temperatures, although it has a significant activation barrier. The transfer of O atoms from cis-(NO)₂ molecules is also observed in Ref. 63 after the exposure of Ar/NO/CO samples to photons at \( \lambda = 220–330 \) nm.

Our experiments show that photolysis of nitric oxide in a matrix of carbon monoxide and water produces an even larger amount of CO₂. As recently shown in independent laboratory experiments, CO₂ can be efficiently formed through the reactions

\[ \text{H}_2\text{O} + h\nu \rightarrow \cdot\text{OH} + \cdot\text{H}, \] (22)

\[ \text{CO} + \cdot\text{OH} \rightarrow \text{HOCO} \rightarrow \text{CO}_2 + \cdot\text{H}, \] (23)

where direct dissociation of the HOCO complex leads to the formation of CO₂ ice.

Photodissociation of water ice (reaction (22)) can also explain the formation of the final products only detected when a NO:H₂O ice was exposed to UV photons: H₂O₂ by recombination of two OH radicals, HNO by H-atom addition to NO (reaction (1a)), and possibly HNO₂ by recombination of nitric oxide with a hydroxyl radical

\[ \text{NO} + \cdot\text{OH} \rightarrow \text{HONO}. \] (24)

Additionally, water can react with a “hot” oxygen atom to form hydrogen peroxide

\[ \text{H}_2\text{O} + \cdot\text{O} \rightarrow \cdot\text{OH} + \cdot\text{OH} \rightarrow \text{H}_2\text{O}_2. \] (25)

This reaction pathway competes with the formation of N₂O₃ in water-rich matrices (reactions (17a) and (17b)) by consuming free O atoms and may be the cause of the low final yield of N₂O₃ detected in water-rich ices (see Fig. 4).

The newly formed products may themselves participate in UV induced chemistry. For example, UV photodissociation of CO₂ and N₂O leads to the formation of free O atoms

\[ \text{CO}_2 + h\nu \rightarrow \text{CO} + \cdot\text{O}, \] (26)

\[ \text{N}_2\text{O} + h\nu \rightarrow \text{N}_2 + \cdot\text{O}. \] (27)

These reactions can play a role in oxygen transfer, but become important only after a long photolysis period with a total UV exposure of at least \( 10^{17} \) photons cm⁻². This flux corresponds to more than hundred incident UV photons per adsorption spot.

The low ionisation potential of cis-(NO)₂ can lead to the formation of NO⁺ and (NO)₂⁺ in the ice upon absorption of Ly-α radiation. However, it is not clear from our experimental data to what extent ion chemistry is involved in the O atom transfer and the subsequent formation of N₂O₃ and CO₂.

The non-detection of NH₂OH in the NO:H₂O ice photolysis experiments indicates that there is a lack of free H atoms in the bulk of the ice. Therefore, reaction (22) has a low efficiency and/or recombination of hydrogen atoms is an efficient mechanism under our experimental conditions. Another mechanism which could reduce the final NH₂OH yield is the efficient photodissociation of the intermediate HNO. Recently, we tried to form NH₂OH by UV photolysis of NH₃:H₂O and NH₂:O₂ ice mixtures. These experiments show that the formation of NH₂OH in the bulk of NH₃:H₂O ice is not efficient either, while UV photolysis of NH₃:O₂ ice mixtures gives uncertain results.

IV. ASTROCHEMICAL IMPLICATIONS

In quiescent dark clouds, grains provide surfaces on which species can accrete, meet, and react and to which excess reaction energy can be donated. Grain surface chemistry is governed by the accretion rate of the gas-phase species and the surface migration rate. The timescale at which gas-phase species deplete onto grains is about \( 10^6 \) years in dense clouds with lifetimes of dense cores between \( 10^5 \) and \( 10^6 \) years. Early during dense cloud formation a H₂O-rich (polar) ice containing mainly CO₂ and traces of CH₄ and NH₃ forms. Under these conditions, NO can be also formed through surface reactions. In prestellar cores, where densities are as high as
10^5 cm^{-2} and temperatures are 10–20 K, gas-phase CO molecules freeze out onto the water ice layer forming a non-polar ice which is the site for additional and more complex reaction pathways. Under these conditions, gas-phase formed NO can also deposit onto icy grains. Species like N_2O, HNO, and NO itself can be formed through surface reactions involving H-, N-, and O-atoms. These species are, therefore, expected to be present in both polar and non-polar ices during the prestellar core-phase. In a recent astronomical study, we discussed the efficiency of the formation of hydroxylamine starting from the hydrogenation of a pure NO ice. In this work, by studying the hydrogenation of NO in H_2O-, CO-, N_2-ice mantles, a full reaction scheme in polar and non-polar environments has been derived.

Due to the low abundance of interstellar solid NO compared to water ice the upper part of the chemical network presented in Fig. 8 is not likely to be astronomically relevant. The formation of a (NO)_2 dimer in space is indeed not expected to happen due to the low probability of finding NO molecules in this environment. However, the hydrogenation of NO molecules and the subsequent formation of NH_2OH is efficient (through a barrierless mechanism) at cryogenic temperatures (15 K) in both water- and carbon monoxide-rich ices. Therefore, according to our experimental results we expect that in quiescent dark clouds NO is efficiently converted to hydroxylamine. However, special care is needed to extrapolate the laboratory data to astronomical timescales and fluxes, as our experiments, for obvious reasons, are not fully representative for interstellar conditions.

To show that the presented reaction scheme (Fig. 8) can indeed lead to the formation of NH_2OH in dense cores, we have performed a gas-grain model using a full gas and grain chemical network of more than 6000 reactions between 460 species of which 195 are involved in both gas-phase and grain-surface reactions. For this we used the OSU gas-grain code, first described by Ref.66 with some recent modifications, which are described in Ref. 67. The initial abundances and chemical network are taken from a benchmark study of chemical codes.66 The physical conditions are chosen similar to their “TMC 1” model which represents a prototypical dense cloud. A fraction of the formed surface species is allowed to desorb into the gas phase according to the Rice-Ramsperger-Kassel theory with an efficiency of $\alpha = 0.01$.

Most surface reactions in chemical networks are based on chemical intuition or are included by analogy to gas-phase reactions. The reaction between H and HNO in standard chemical networks leads to H_2 and NO, while NH_2OH is usually formed through NH_2 + OH. Therefore, in standard chemical networks, no NH_2OH is formed and NO ice is mostly converted into HNO or N_2O under dark cloud conditions. This is in agreement with Garrod et al. who performed chemical models of star formation and who found that NH_2OH is only abundantly formed in the “hot-core” phase when the grain is lukewarm and NH_2 and OH become mobile.

Since in cold dense clouds the visual extinction is around 10 mag and photodissociation events are rare, the surface abundance of radical species like OH and NH_2 is low. The low grain-surface temperature of 10 K reduces the mobility of these species, while hydrogen atoms are already mobile at this temperature. Moreover, UV photolysis experiments show that hydroxylamine is not formed by irradiating NO:CO:H_2O nor NH_3:H_2O ice mixtures at low temperatures. The NO hydrogenation scheme as experimentally investigated in the present study is, therefore, much more promising as a NH_2OH formation reaction route. In our model, that includes the NO hydrogenation scheme, NH_2OH is indeed abundantly present in both the solid and the gas phase. The hydroxylamine abundance peaks at 10^5 years. Photodissociation into OH and NH_2, of which the latter converts into NH_3 upon hydrogenation, starts to become important at this time. Snow et al.72 showed that gas-phase NH_2OH can react with CH_3^+ to form NH_3OH^+ that can in turn react with different carboxylic acids to form protonated amino acids. These reactions are not included in the networks used. Our model, however, predicts that hydroxylamine has a relative abundance of at least $7 \times 10^{-9}$ nH on the grain surface and $3 \times 10^{-11}$ nH in the gas phase at the end of the lifetime of a dense cold cloud (1–10 Myr), before the core collapses.

To date, and to the best of our knowledge, interstellar hydroxylamine has not been observed in space, and only upper limits have been found for a number of sources.73 These correspond to fractional abundances of <10^{-11} in the gas phase. This is around the same order of magnitude as our gas-phase model predictions for dense clouds. We therefore expect that in the near future the faint radio signal from NH_2OH gas-phase molecules will be recorded by the receivers of ALMA that will investigate the inner parts of proto-planetary disks and hot-cores with relatively high temperatures.74 The hydroxylamine locked onto the grains at the beginning of the core collapse becomes accessible for further reaction at later stages when a protostar forms and UV irradiation and thermal processing start. The astrobiological implications of our study are constrained by the results of Ref. 15. They proposed a route for the synthesis of amino acids in interstellar environments starting from desorption of NH_2OH. In their study, NH_2OH is formed through energetic processing at a late stage of star formation. Here, we have shown how NH_2OH is already formed in the solid phase though a non-energetic route in the early stages of star formation. Figure 9 shows

![FIG. 9. Proposed formation route of amino acids in star-forming regions based on the hydroxylamine surface formation route investigated in the present study combined with the experimental work of Ref. 15.]
the link between the updated solid-phase NH₂OH formation scheme and the gas-phase formation routes of amino acids in space. Here, the key stage of this suggested mechanism is that in which reactions between protonated hydroxylamine and carboxylic acids occur. Although these reaction pathways were already proved experimentally to be efficient at room temperature, recent calculations show that they may involve activation barriers and, therefore, that they may not be efficient under astronomically relevant conditions. Thus, there is now great demand for experimental verification of the latter results.

Finally, acetic acid, another main component of this reaction network, has already been detected towards several astronomical objects (mainly hot-cores) and Bennett et al. showed successfully the laboratory synthesis of acetic acid by UV photolysis of CO₂:CH₄ ices under UHV conditions and cryogenic temperatures.

V. CONCLUSIONS

A systematic study of NO hydrogenation pathways under UHV conditions and for cryogenic temperatures is presented. In particular, the influence of sample temperature, NO deposition rate, H-atom flux, and UV photolysis on the formation of different hydrogenation products is studied. Special attention is given to the hydrogenation of NO in CO- and H₂O-rich ices and the astronomical implications of our results are discussed. The main conclusions of this work are:

1. In CO- and H₂O-rich interstellar ice analogues NO is efficiently converted to hydroxylamine upon exposure to thermal H atoms.
2. The formation of hydroxylamine takes place through a fast and barrierless mechanism by subsequent addition of three H atoms to a single NO molecule.
3. HNO is an intermediate product of hydroxylamine formation and is observed in our experiments only under low H atom flux conditions.
4. The final HNO and NH₂OH formation yield is higher at 15 K than at 25 K.
5. In water-rich interstellar ice analogues, the final NH₂OH formation yield is higher than in CO- rich ices.
6. In CO-rich ices hydroxylamine is formed well before CO molecules are hydrogenated to form formaldehyde and methanol.
7. N₂O is a side product of the NO hydrogenation reaction scheme.
8. Solid NH₂OH is not formed upon UV (Ly-α) photolysis of NO:CO:H₂O ices. The end products are N₂O, N₂O₃ (or NO₂), HNO, H₂O₂, and CO₂ ice.
9. According to our laboratory and modelling results, we expect that NH₂OH is already present in the solid phase in dark molecular clouds and mainly in the gas phase in protostar regions. We therefore expect that in the near future gas-phase NH₂OH will be detected by ALMA.
10. Hydroxylamine formed efficiently in the solid phase provides a gas-phase reservoir upon desorption as a starting point in the formation of glycine and β-alanine.

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