OGLE16aaa - a Signature of a Hungry Super Massive Black Hole


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ABSTRACT

We present the discovery and first three months of follow-up observations of a currently on-going unusual transient detected by the OGLE-IV survey, located in the centre of a galaxy at redshift z=0.1655. The long rise to absolute magnitude of -20.5 mag, slow decline, very broad He and H spectral features make OGLE16aaa similar to other optical/UV Tidal Disruption Events (TDEs). Weak narrow emission lines in the spectrum and archival photometric observations suggest the host galaxy is a weak-line Active Galactic Nucleus (AGN), which has been accreting at higher rate in the past. OGLE16aaa, along with SDSS J0748, seems to form a sub-class of TDEs by weakly or recently active super-massive black holes (SMBHs). This class might bridge the TDEs by quiescent SMBHs and flares observed as “changing-look QSOs”, if we interpret the latter as TDEs. If this picture is true, the previously applied requirement for identifying a flare as a TDE that it had to come from an inactive nucleus, could be leading to observational bias in TDE selection, thus affecting TDE-rate estimations.

Key words: galaxies: individual: OGLE16aaa – galaxies: active – black hole physics

1 INTRODUCTION

It has become a paradigm that nearly all galaxies at current times harbour a super-massive black hole (SMBH) in their centre (e.g., Magorrian et al. 1998). In the cold dark matter (ΛCDM) theory of cosmology, current (redshift zero) galaxies are the product of hierarchical mergers of smaller galaxies. These smaller building blocks also host black holes

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in their centres (Kormendy & Richstone 1995, Greene 2012), potentially intermediate-mass black holes (IMBHs), with masses from 100 to 10,000 M☉. After two galaxies merge, the two black holes will merge as well (see Begelman, Blandford, & Rees 1980). Therefore, mergers of black holes may play an important role in building SMBHs (cf. Pelupessy, Di Matteo, & Ciardi 2007). Interestingly, SMBHs with masses of more than a billion M☉ have been found already at redshifts of more than 6 when the Universe was less than 1 Gyr old (see Fan 2006). Such SMBHs may be seeded by 100 M☉ black holes at redshifts z>15 which then grow by
uninterrupted accretion of gas at the Eddington rate with a standard radiative efficiency of 10 per cent (e.g., Haiman 2013). However, quasars grow only for about \(4.5 \times 10^7\) years before feedback stops the gas supply (Silk & Rees 1998). In order to solve this problem one can start with more massive black holes such as IMBHs and/or allow mass to be accreted at a rate higher than the Eddington limit and/or have part of the black hole growth be due to mergers of black holes.

Tidal Disruption Events (TDEs, e.g., Hills 1975, Rees 1988), in which a star is torn apart by the tidal forces of the SMBH, offer a unique opportunity to study the mass distribution of SMBHs. The intrinsic TDE rate should be dominated by the SMBHs with the lowest mass (Wang & Merritt 2004; Stone & Metzger 2016), so volume-complete TDE samples can measure the occupation fraction of SMBHs in small galactic bulges informing SMBH formation theories. However, the inhomogeneous and small sample of TDEs currently available, found either in X-rays (e.g., Bade, Komossa & Dahlén 1996; Nikolaev & Walter 2013) or in the UV/optical (e.g., van Velzen et al. 2011; Gezari et al. 2012; Arcavi et al. 2014; Holoien et al. 2014) prevents us from discriminating between various emission mechanisms of TDEs and hence from conclusions on the SMBH mass function (e.g., Stone & Metzger 2016).

Optical/UV TDEs are relatively luminous (\(M_{peak} \sim -20\)) and blue (\(T \sim \text{few 10}^4\)K) few-months- to years-long transients with broad H and/or He II emission lines (Arcavi et al. 2014; Arcavi et al. 2014; and French, Arcavi & Zabludoff 2010) noted that 75 per cent of optically found TDEs occurred in quiescent Balmer-strong (called E+A by some) galaxies, which account for 2.3 per cent of SDSS galaxies. Such galaxies are thought to be products of a recent merger (within 1 Gyr), which triggered an observed increase in star formation (Balmer absorption series in their spectra are caused by a significant amount of A-type stars). That preference might be due to disturbed dynamics of the nuclear star cluster, or the presence of a coalescing binary black hole, causing nearby stars to go on a collision course with their central black holes (e.g., Wegg & Nate Bode 2011).

However, a fraction of the optical/UV TDEs seems to not match such scenario. Host galaxy spectra of Extreme Coronal Line Emitters (ECLE), SDSS J095209.56+214313.3 (Komossa 2008; Palaversa 2016) and SDSS J0748 (Wang et al. 2011), as well as ASASSN-14li (Holoien et al. 2016) and PTF09axc (Arcavi et al. 2014) present weak, narrow emission lines, which could be indicative of a weak Active Galactic Nucleus (AGN) present in the core.

AGNs are known to exhibit photometric variability at a level of few tens of magnitude (e.g., Kozłowski et al. 2016). However, occasionally, flares are observed well above the level of their typical variability, both in X-rays (e.g., Strożyk and Johannsen 2016) and optical (e.g., Tanaka 2013). The reasons for such significant changes in mass accretion rates are still under debate and include binary black hole interactions with the disk as well as stellar disruptions.

Long-term and wide-field spectroscopic and photometric data obtained by the Sloan Digital Sky Survey (SDSS), as well as on-going transient searches (ASAS-SN, PanSTARRS, MASTER surveys), have revealed nearby AGNs (Seyfert galaxies), that changed their spectral characteristics, often accompanied with a temporal increase in observed flux, e.g., Lawrence et al. 2010. In particular, an event observed in the quasar SDSS J0159+0033 (LaMassa et al. 2015), has been recently interpreted as a TDE by a massive SMBH (Merloni et al. 2015). In large fraction of other similar object, so called, “changing-look quasars” (hereafter, CH-L-QSOs) (MacLeod et al. 2016), the appearance of a broad component to the H\(\alpha\) emission line was found to be transient (Baldassare et al. 2016), i.e., the broad line disappeared on the time-scale of years. Moreover, one of such CH-L-QSOs, found in NGC 2617 by ASAS-SN in 2013 (Shappee et al. 2014) has recently been reported to show a re-brightening (Oknyansky et al. 2016a). If those flares are indeed due to TDEs, the stellar disruption must be induced fairly frequently, namely, on a time-scale of years. An X-ray TDE candidate in an AGN IC 3599 was found to show a reoccurrence period of about 9 years (Campana et al. 2015), which could be explained with a binary central black hole.

In this Letter we describe the discovery and first months of follow-up of an unusual transient, OGLE16aaa. We propose that OGLE16aaa is a TDE candidate in a low-luminosity AGN host galaxy and it may be another link in chain of disruption events that occur from quiescent to active galaxy cores.

2 DISCOVERY AND EARLY FOLLOW-UP

OGLE16aaa (Fig. 2) was discovered by the OGLE Transient Detection System (Wyżykowski et al. 2014), a programme within the Optical Gravitational Lensing Experiment (OGLE-IV, Udalski et al. 2015). This transient, located at RA, Dec(J2000.0) = 1:07:20.88, -64:16:20.7, was found during visual inspection of candidates detected by the automated pipeline on January 2nd, 2016 at I-band magnitude \(\sim 20\) mag (Wyżykowski et al. 2016). It was found at the centre of GALEX A spect J010720.81-641621.4 galaxy of 17.00±0.01 mag as measured by OGLE in the I-band and 20.83±0.12 mag and 21.82±0.32 mag in GALEX NUV and FUV, respectively. Fig. 1 shows the host galaxy image from a deep stack of \(\sim 100\) OGLE I-band images (all taken before the event) with the position of the transient.

On 17 Jan 2016 the Public ESO Spectroscopic Survey

\[^{1}\text{http://ogle.astrouw.edu.pl/ogle4/transients}\]
3 HOST AND TRANSIENT CHARACTERISTICS

Location. The offset of the transient from the galaxy’s photo-centre, obtained from the OGLE I-band images using DIA, is less than 160±140 mas (460±400 pc for z=0.1655), as shown in Fig. 1 and is consistent with the position of the nucleus.

Host galaxy. From the light distribution of the host on the deep OGLE I-band image we obtained a Sersi´c index of $n=1.08$, which corresponds to a black hole mass of $\log M_{\text{BH}} = 6.58 \approx 4 \times 10^{6} M_{\odot}$ using the relation in (Savorgnan 2016).

Since there is, to our knowledge, no pre-flare host spectrum available, we have to wait until the event fades out completely to obtain the spectrum of the nucleus of the galaxy alone. However, the La PHARE (Ilbert et al. 2006) best matching template spectrum to the spectral energy distribution (SED) from archival UV, optical, near- and far-infrared observations is found for a galaxy with a stellar mass of $\log M = 10.3 \pm 0.2 M_{\odot}$, with no strong evidence for star formation.

Signatures of the host are already present in the
PESSTO spectrum of the flare. The spectrum shows weak [OII], [OI] and [NII] narrow forbidden lines in emission as well as narrow emission of Hα and Hβ. Assuming the broad features are from the transient and the narrow ones are from the host, the host spectrum resembles that of SDSS J0748 (Wang et al. 2011) which hosted a TDE candidate. The galaxy does not seem to belong to the quiescent Balmer-strong galaxies, which were noted to host most of the optically discovered TDEs so far [French, Arcavi & Zabludoff 2016]. The line ratios (log([OII]/Hα) = −0.43 v s log([OIII]/Hβ) = −0.25) indicate there is a mixture of star formation in the galaxy of OGLE16aaa as well as a weak AGN in its core. The line ratios are also similar to those in the host of SDSS J0748. Moreover, the WISE (Wright et al. 2010) colours for OGLE16aaa and SDSS J0748 place those hosts near the region occupied by AGNs, following the classification method of Assef et al. (2010). However, a spectrum taken after the transient emission fades is still required for a quantitative comparison to other TDE hosts and for disentangling the continuum level of the transient from the host.

The host galaxy is seen face-on and has no obvious signature of spiral structure, apart from a faint hint of an extended structure to the West from the centre (Fig. 1).

Flare spectrum. The flare spectrum, taken by PESSTO at ~3 rest-frame days from the I-band maximum, supplemented with the earliest Swift observation in UVM2, taken at around the peak, is consistent with black-body model with temperature higher than 22,000 K. The host emission is clearly present in the spectrum at wavelengths longer than rest-frame 4000 Å. The residual spectrum, after subtraction of the black-body continuum, shows two broad emission features around HeII and Hα, as seen in all optically–selected TDEs [Arcavi et al. 2014] reported on a continuum of broad HeII and Hα emission lines (after the host-galaxy light has been subtracted) in the spectra of TDEs. OGLE16aaa seems to fit the picture very well, though with somewhat higher velocity dispersion (~23,000 km/s and ~19,000 km/s for HeII and Hα lines, respectively, corresponding to FWHMs of ~850Å (54,000 km/s) and ~970Å (45,000 km/s), respectively).

Light curve. There is no apparent flaring nor variability activity in the historical (3.5 yrs) OGLE I-band light curve of the nucleus of the host galaxy at a level below I ~22 mag (host subtracted), less than 1 per cent of host light. The optical light curve reached the peak in about 30 rest-frame days and then exhibited significant variability, particularly around 15 restframe days after I-band maximum. Comparing our brightest UV measurements to the archival GALEX data we estimate the overall UV amplitude of the flare of about 3 mag. The overall decline of the light curve in both optical and UV is very slow, however, the actual slope of the decline is yet to be determined in the observations in the second half of 2016. No X-ray emission was detected by Swift/XRT at a level above 0.002 counts/s (3σ), corresponding to an upper limit for the unabsorbed luminosity of 5E42 erg/s (0.3-10 keV), for a power-law with photon index of ~2.

Tidal Disruption Model. We fit OGLE16aaa’s photometry with the tidal disruption light curve fitting software TDEFit (Guillochon et al. 2014; Vinkó et al. 2015) (see Fig. 1), a Monte Carlo modelling code. For OGLE16aaa, we presume that the observed light comes from a combination of the light produced by a viscously-driven disk component (Guillochon et al. 2016, in prep), emission from circularisation [Jiang et al. 2016], and reprocessing of light from the debris that ensheaths the accretion disk structure [Guillochon et al. 2014]. We assume a flat prior for Mbh that allows all black hole masses between 10^4 and 10^9 M☉, a prior on impact parameter β that assumes pinhole scattering (P(β) ∝ β−2), a Kroupa stellar mass function [Kroupa 2001], and flat priors on all other parameters. We find that the most likely combination of disruption parameters is Mbh = 10^6−2×10^7 M☉, M* = 10−0.5×10^4 M☉ (between 0.1 and 0.8 M☉), median ~0.3 M☉) and β = 1.77−0.53, with a degeneracy between a sub-solar star suffering a full disruption and a solar star suffering a partial disruption. The total observed energy emitted in the event is about 5e52 ergs. For assumed (median) 8 per cent of efficiency the total accreted mass is therefore about 0.3 M☉, indicating either complete or partial disruption of the star and suggesting the mass of the star was probably higher than 0.6 M☉.

4 DISCUSSION

The characteristics of OGLE16aaa resemble those of other optically found Tidal Disruption Events. First, the flare’s location coincides with the nucleus of the host galaxy. Also, the derived photospheric black-body temperature remains high (~20,000 K) throughout the available data for this event, significantly higher than in typical supernovae. Moreover, the temperature seems to rise at 35 rest-frame days from I-band maximum, as seen before in ASASSN-14ae TDE [Holoien et al. 2014]. Both the optical light curve and the presence of very broad HeII and Hα in the spectra resemble those of other optical/UV TDEs [Arcavi et al. 2014]. OGLE16aaa seems, therefore, to be a TDE. Lack of any variability in 3.5 years prior to the flare and large amplitude of the flare strongly disfavours regular AGN flaring.

However, the underlying host galaxy is somewhat different in OGLE16aaa than in most of known TDEs so far [French, Arcavi & Zabludoff 2016] have shown that most optically-found TDEs detected so far occur in quiescent, Balmer-strong galaxies, which are considered post-mergers. The Balmer absorption line series are not present in case of OGLE16aaa, however, a deep post-flare spectrum is still needed to verify it. Another TDE candidate, SDSS J0748, also does not show Balmer series and is an outlier on Figure 2 of [French, Arcavi & Zabludoff 2016].

Among the X-ray-detected TDEs, IGR J12580, interpreted as due to a flare due to disruption of a Jupiter-mass planet was also detected in a weak AGN/Seyfert galaxy NGC4845 [Nikolajuk & Walter 2013]. The narrow emission lines ratios in all those three hosts, as well as their WISE colours indicate that the host contains weak-AGN.

For most TDEs found so far, the black hole was assumed dormant, since there was no evidence to the contrary. Here we propose that OGLE16aaa and several other TDE candidates are due to a stellar disruption in a weak-AGN or Seyfert 1-type galaxy, where narrow emission lines originate from the circumnuclear material, photoionised by X-ray photons generated due to accretion. [Bennert et al. 2006] showed that for several Seyfert II galaxies the projected distance of the narrow-line region extends to hundreds and even
thousands of parsecs. Such accretion is likely to have been due to regular AGN accretion, but it could also have been due to a previous TDE. TDEs are expected to repeat on timescales of 10^8 years (Wang & Merritt 2004) and if the paths from the lines of sight from the narrow line region to us represent a broad delay function, we could still see an echo of previous TDEs in the spectra (e.g., Wegg & Nate Bode 2011).

Moreover, the TDEs in narrow emission line hosts seem to be bridging stellar disruptions by dormant SMBHs and those in much more active AGN, which exhibit a transient to be bridging stellar disruptions by dormant SMBHs and an echo of previous TDEs in the spectra (e.g., Wegg & Nate Bode 2011).

In the course of OGLE-IV search for extragalactic transients, we have discovered a new candidate for a TDE of a 0.1-0.8 M_☉ star by a 10^6-2 M_☉ SMBH.

OGLE16aaa event, along with SDSS J095209.56+214331.3, SDSS J0748, ASASSN-14li and IGR J1258 seem to form a sub-class of TDEs in galaxies hosting a weak AGN, with weak narrow emission lines. Moreover, the existence of such TDEs supports one of the explanations for so called “changing-Look QSOs”, where persistent strong narrow emission lines get super-imposed with variable broad emission lines.

A possible explanation of the observed variability in the light curve of OGLE16aaa is that it is induced by a binary black hole on a tight orbit or due do disk precession or circularisation on a timescale of about a month, however, further multi-messenger follow-up is required to understand this.

The fact that TDEs could also be found in (weak) AGNs is important for determining TDE rates, since currently there is likely an observational bias against selection of optically found TDEs. Whereas only 10 per cent of galaxies host an AGN this bias could be larger given that AGN activity is triggered or enhanced by recent mergers and milliparsec-scale binary SMBHs could strongly enhance the TDE rate.

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REFERENCES

Fan X., 2006, NewAR, 50, 665
Firth R., et al., 2016, ATel, 8559.
Greene J. E., 2012, NatCo, 3, 1304
Greiner J., et al., 2016, ATel, 8579
Haiman Z., 2013, ASSL, 396, 293
Komossa S., 2015, JHEAp, 7, 148
Oknyansky V. L., et al., 2016, ATel, 9015
Rees M. J., 1988, Natur, 333, 523
Udalski A., Szymański M. K., Szymański G., 2015, AcA, 65, 1
Wyrzykowski Ł., et al., 2014, AcA, 64, 197
Wyrzykowski Ł., et al., 2016, ATel, 8577
Zhang W., Yu W., Yan Z., 2016, ATel, 8644